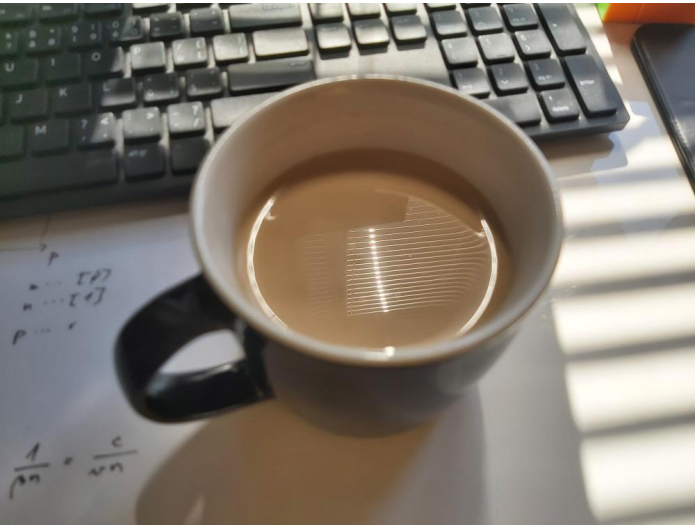
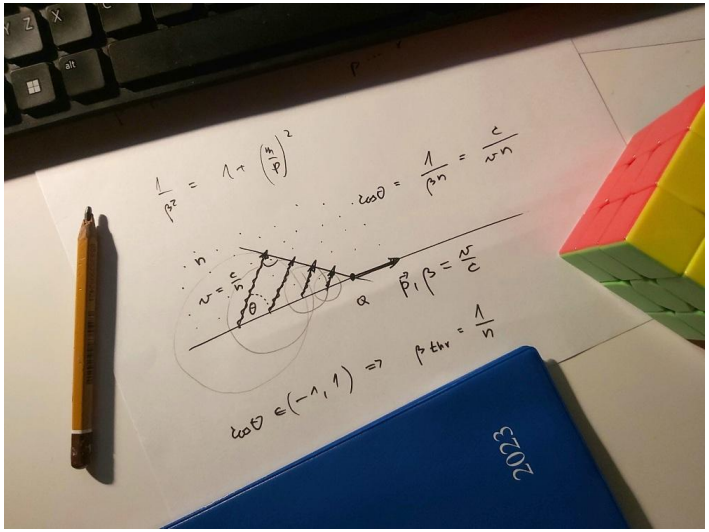
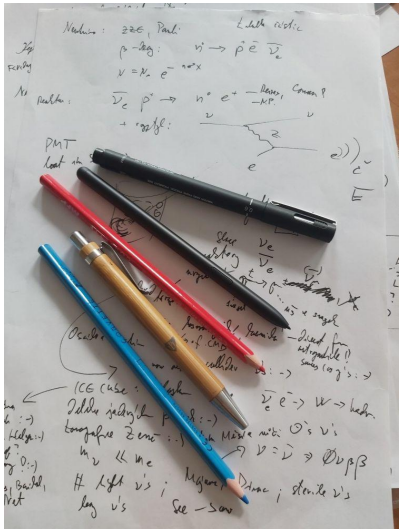






Jak se chytají částice

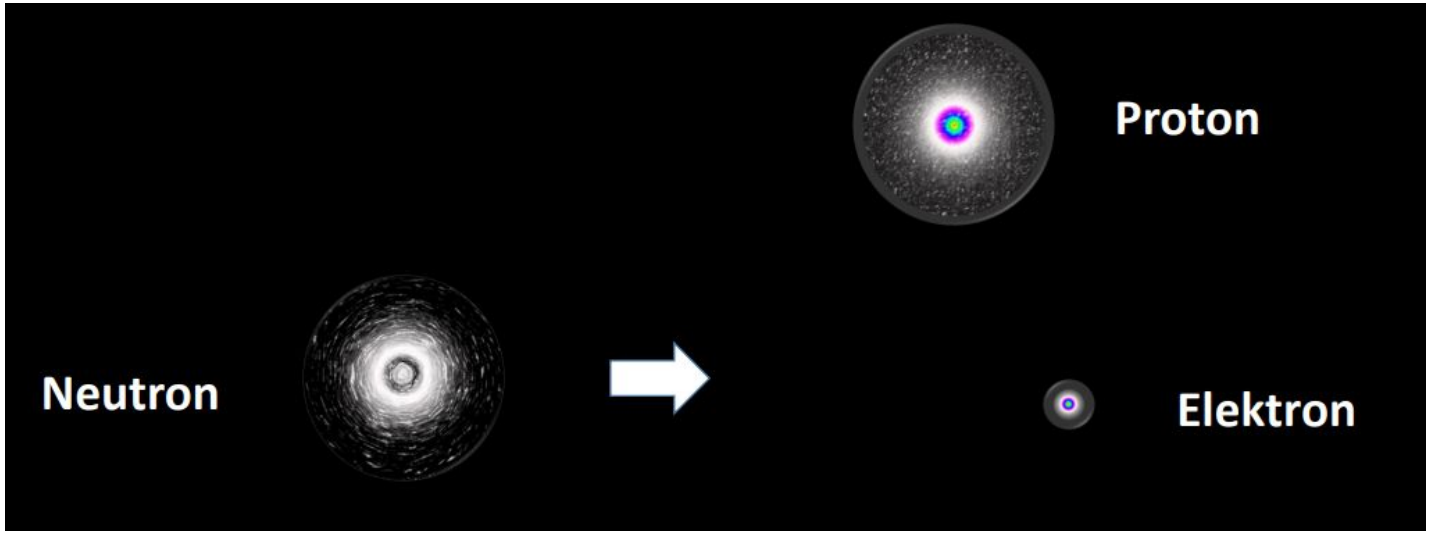
PALACKÝ UNIVERSITY OLOMOUC

Mgr. Jiří Kvita, Ph.D.

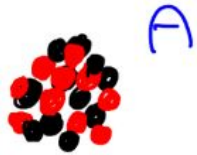
Badatelský víkendový seminář 2023
Společná laboratoř optiky UP

Částice

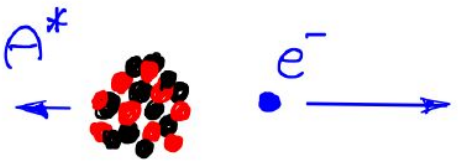
Rozpad neutronu



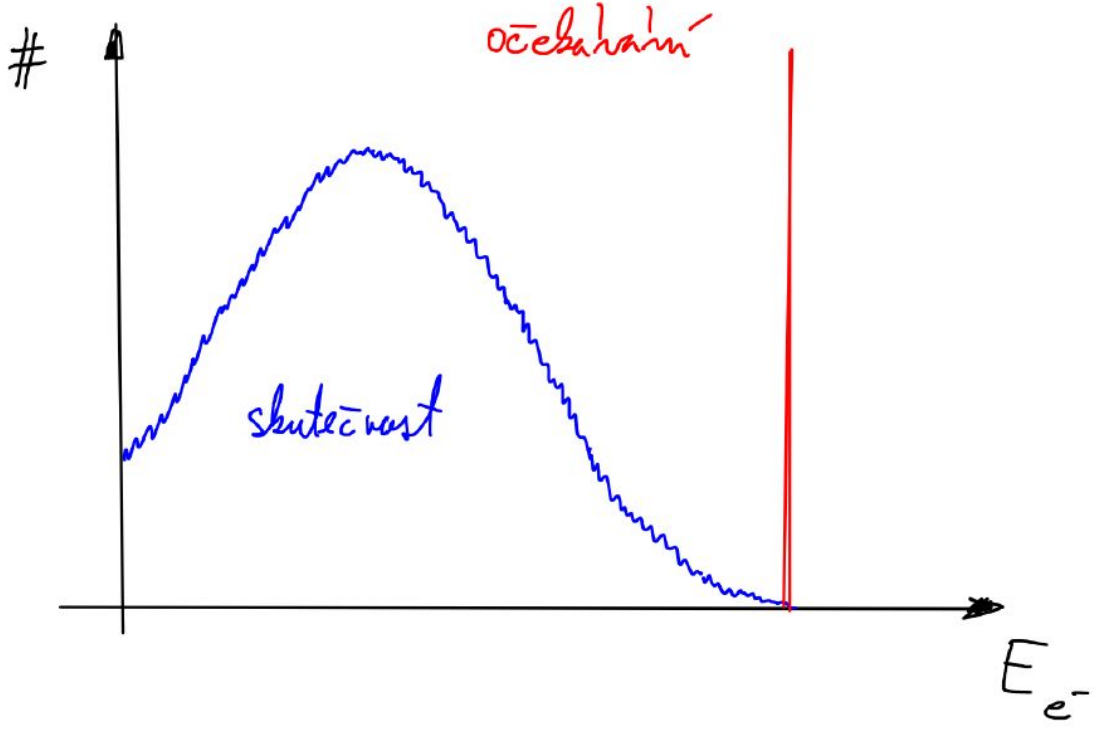
Beta rozpad



Beta rozpad



- Problém: pozorované spektrum energií elektronů z beta rozpady bylo široké, neodpovídalo očekávání jedné energie, kdyby šlo o dvojčásticovy rozpad $A \rightarrow A^{*}\, e^{-}$.
- Dáno zákonem zachování hybnosti a energie.



Neutrino – teoretické předpověď 1930

Wolfgang Pauli

Offener Brief an die Gruppe der Radioaktiven bei der Gauvereins-Tagung zu Tübingen.

Abschrift

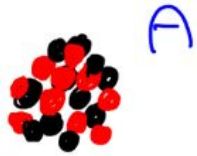
Physikalisches Institut
der Eidg. Technischen Hochschule
Zürich

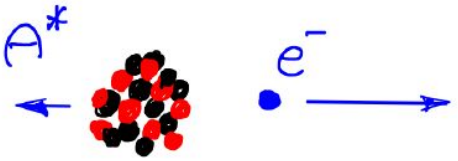
Zürich, 4. Dez. 1930
Gloriastrasse

Liebe Radioaktive Damen und Herren,

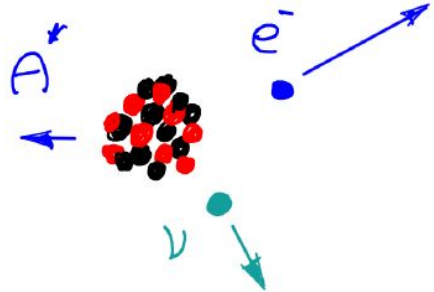
Wie der Ueberbringer dieser Zeilen, den ich huldvollst anzuhören bitte, Ihnen des näheren auseinandersetzen wird, bin ich angesichts der "falschen" Statistik der N- und Li-6 Kerne, sowie des kontinuierlichen beta-Spektrums auf einen verzweifelten Ausweg verfallen um den "Wechselsatz" (1) der Statistik und den Energiesatz zu retten. Nämlich die Möglichkeit, es könnten elektrisch neutrale Teilchen, die ich Neutronen nennen will, in den Kernen existieren, welche den Spin 1/2 haben und das Ausschliessungsprinzip befolgen und

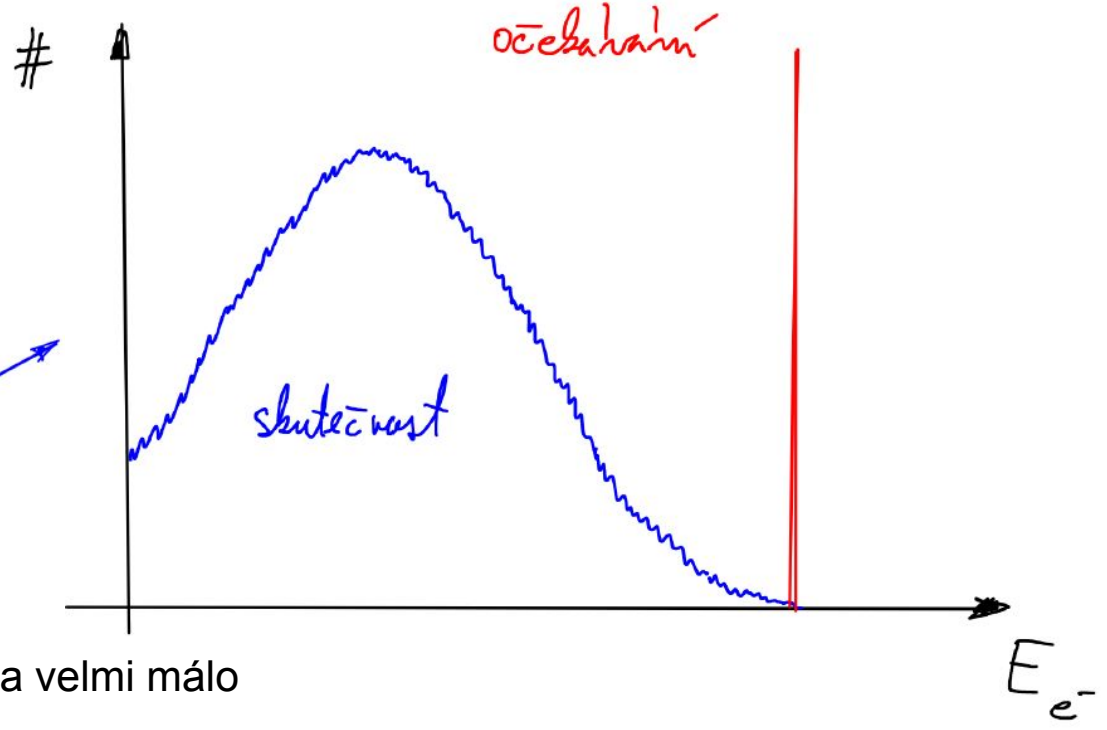
Beta rozpad





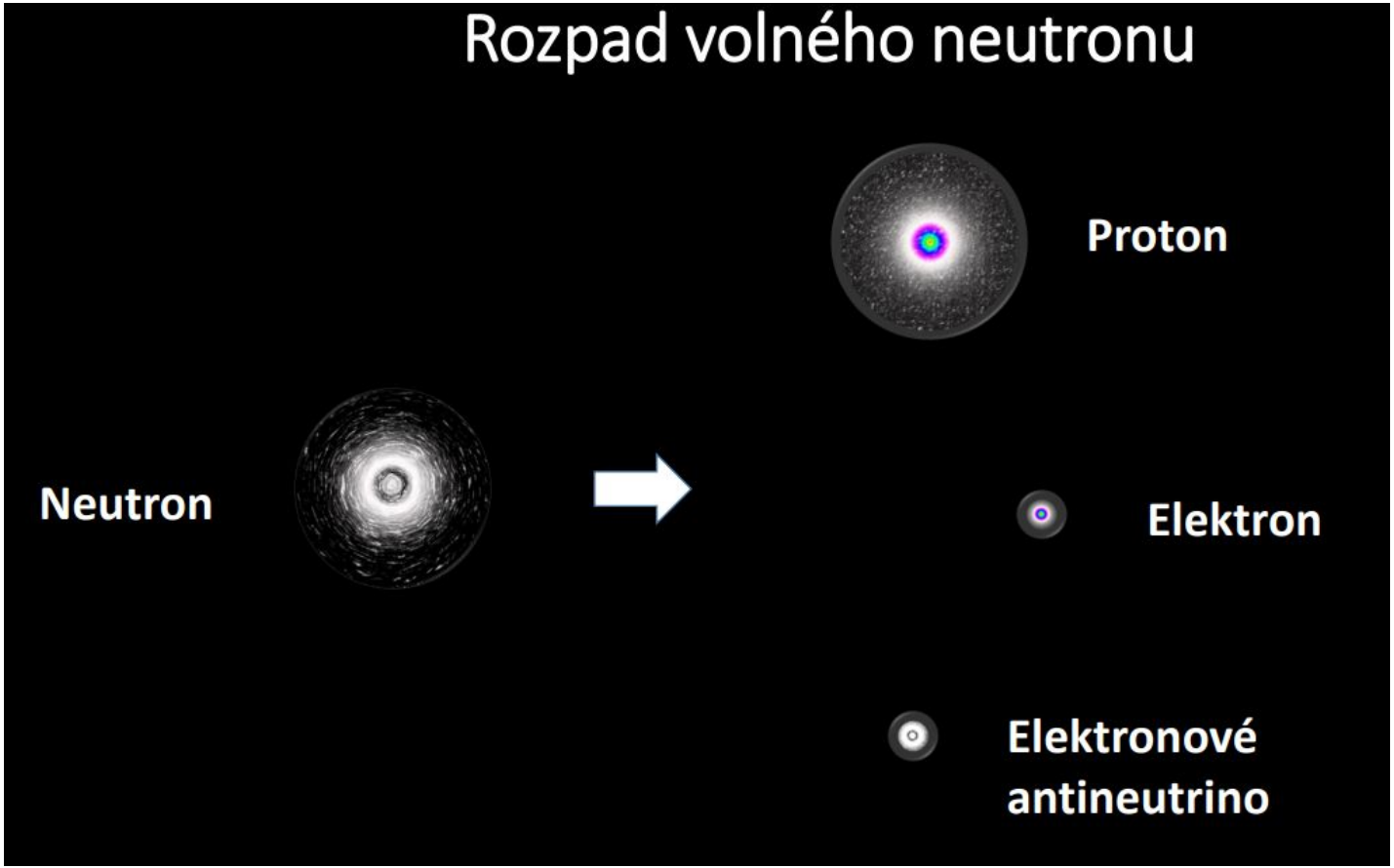
- Problém: pozorované spektrum energií elektronů z beta rozpady bylo široké, neodpovídalo očekávání jedné energie, kdyby šlo o dvojčásticovy rozpad $A \rightarrow A^{*}\, e^{-}$





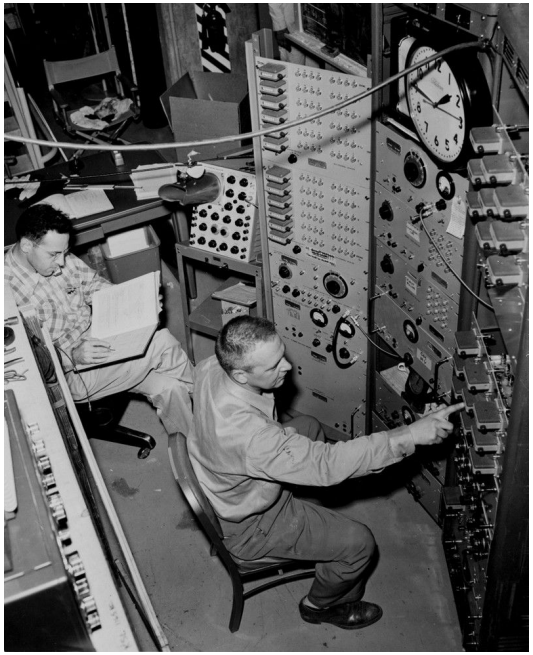
- Neutrino: neutrální, velmi lehká a velmi málo interagující částice.

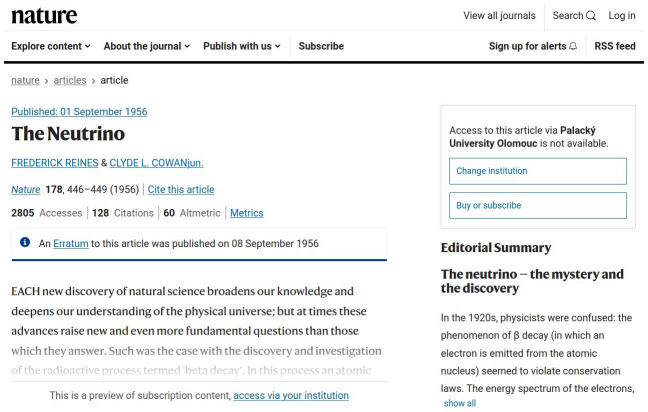
Beta rozpad


Rozpad volného neutronu
Proton
Neutron
Elektron
Elektronové
antineutrino

Objev neutrina :: 1956

- Reaktor Savannah river
- 5×10^{13} neutrin / s / cm^2
- Clyde Cowan, Frederic Reines.





nature

View all journals | Search | Log in

Explore content | About the journal | Publish with us | Subscribe | Sign up for alerts | RSS feed

nature › articles › article

Published: 01 September 1956

The Neutrino

FREDERICK REINES & CLYDE L. COWANjun.

Nature **178**, 446–449 (1956) | Cite this article

2805 Accesses | **128** Citations | **60** Altmetric | Metrics

An Erratum to this article was published on 08 September 1956

EACH new discovery of natural science broadens our knowledge and deepens our understanding of the physical universe; but at times these advances raise new and even more fundamental questions than those which they answer. Such was the case with the discovery and investigation of the radioactive process termed 'beta decay'. In this process an atomic

This is a preview of subscription content, access via your institution

Access to this article via **Palacký University Olomouc** is not available.

Change institution

Buy or subscribe

Editorial Summary

The neutrino – the mystery and the discovery

In the 1920s, physicists were confused: the phenomenon of β decay (in which an electron is emitted from the atomic nucleus) seemed to violate conservation laws. The energy spectrum of the electrons,

show all

The Nobel Prize in Physics 1995

Photo from the Nobel Foundation archive.

Martin L. Perl

Prize share: 1/2

© University of California Regents

Frederick Reines

Prize share: 1/2

https://www.symmetrymagazine.org/article/neutrino-turns-60

https://en.wikipedia.org/wiki/Cowan%E2%80%93Reines_neutrino_experiment

Objev neutrina :: 1956

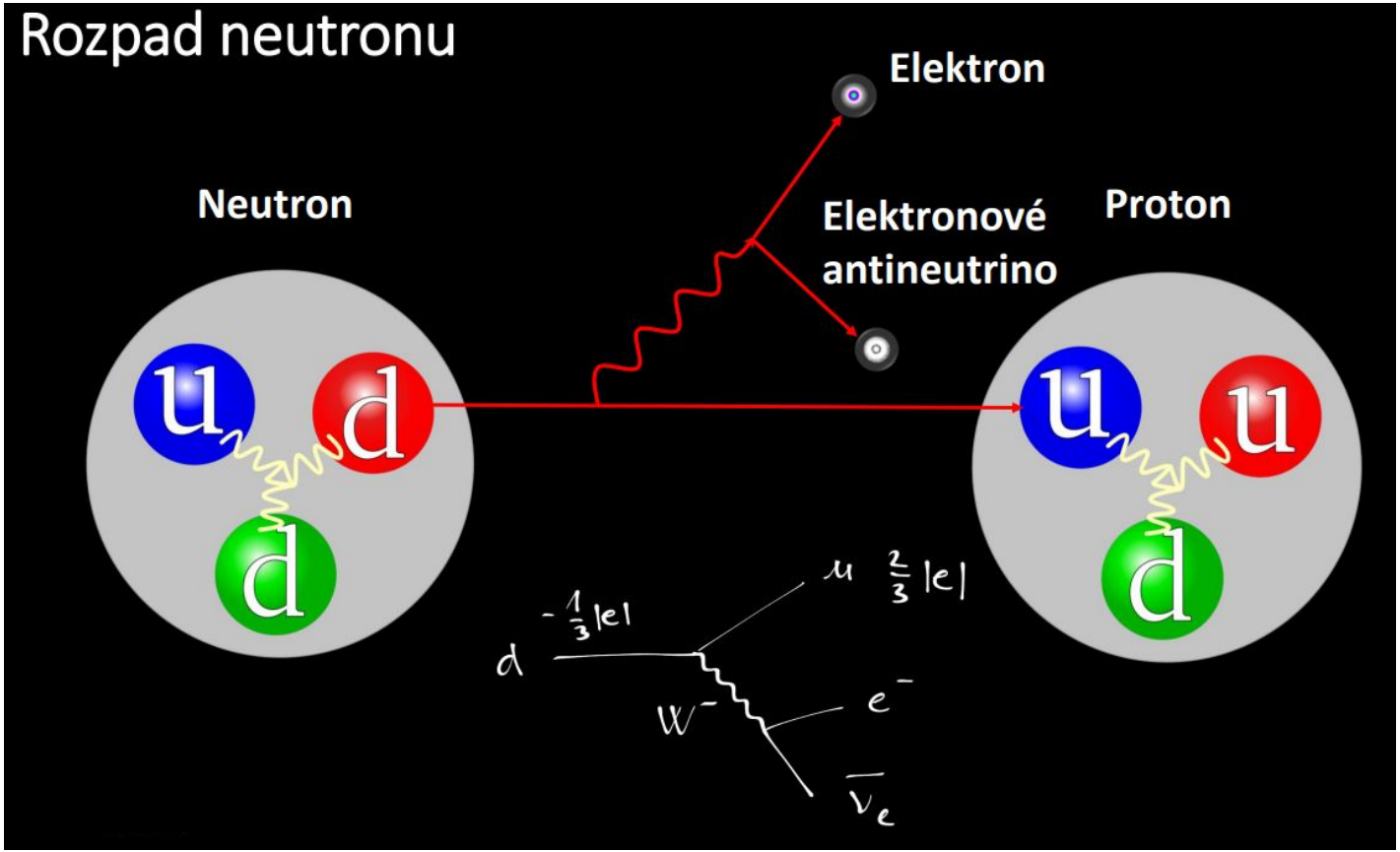
$n^0 \rightarrow p^+ e^- \bar{\nu}_e$

$\bar{\nu}_e p \rightarrow n^0 e^+$ $\qquad$ $e^+ e^- \rightarrow \gamma\gamma$

n $^{108}Cd \rightarrow {}^{109m}Cd \rightarrow {}^{109}Cd + \gamma$

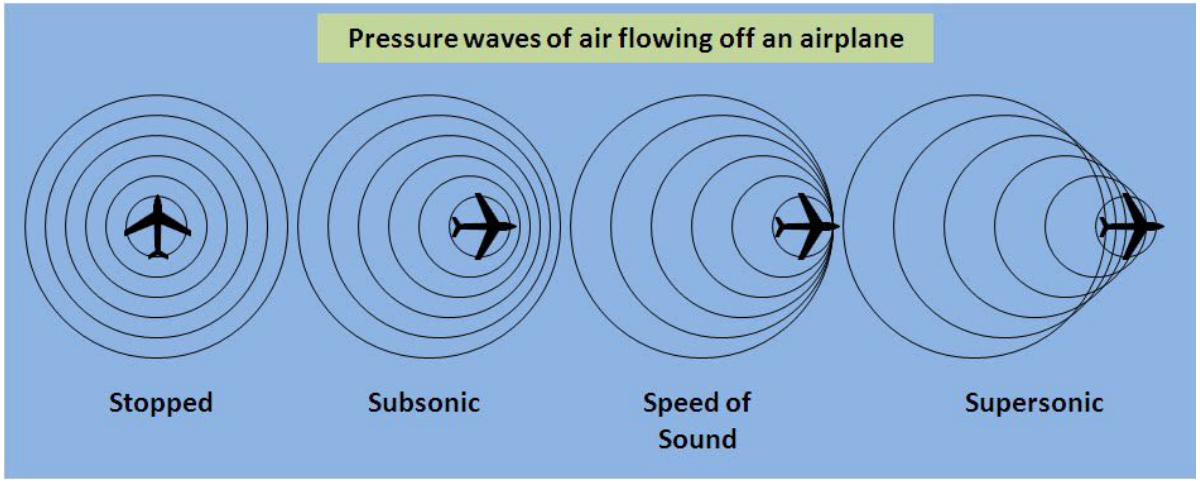
$CdCl_2$

Neutrina z beta rozpadu



Odbočka: Rázová akustická vlna

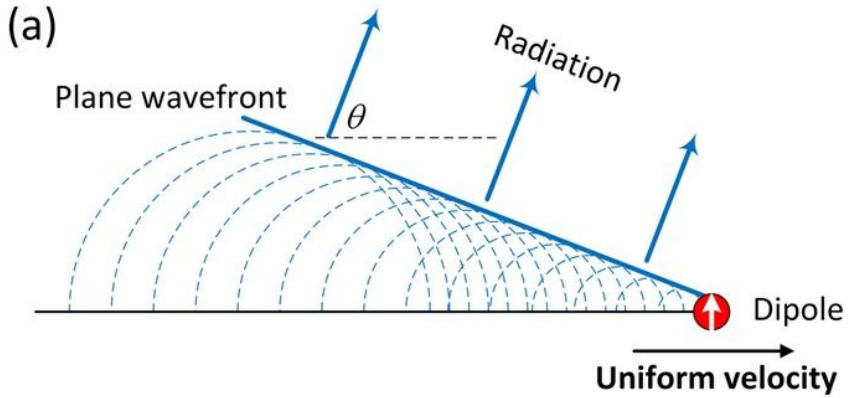
- Letadlo překračuje rychlost zvuku

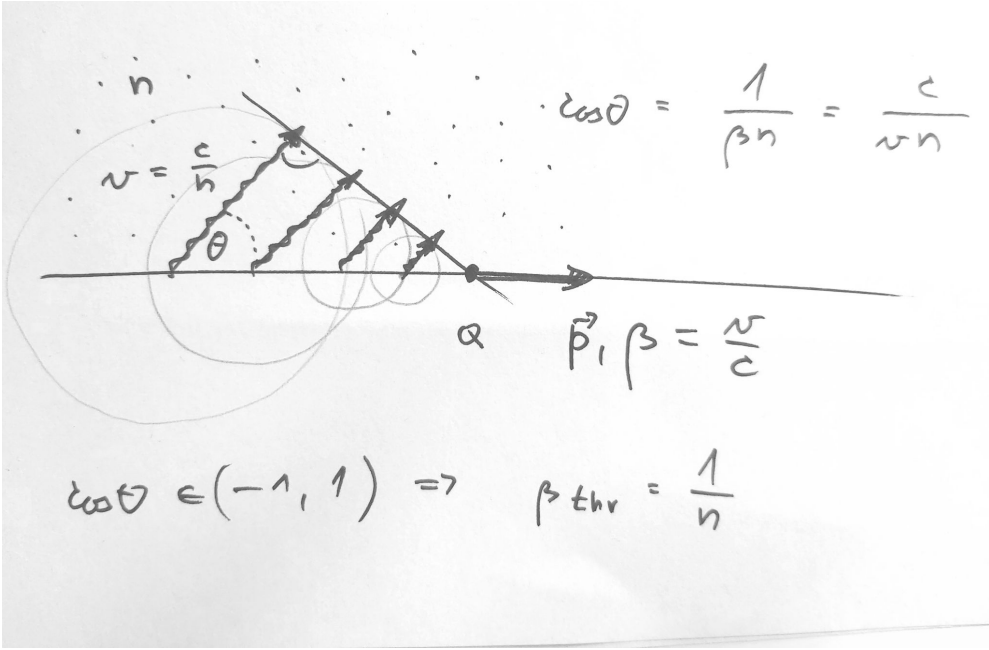


https://physics.stackexchange.com/questions/282353/how-does-particle-speed-affect-cherenkov-radiation
https://en.wikipedia.org/wiki/Sound_barrier

Čerenkovovo záření

- Nabitá částice překračuje rychlost světla v hmotném prostředí
  - ideálně transparentní médium
- $v_{světla}$ = c / index lomu = c / n

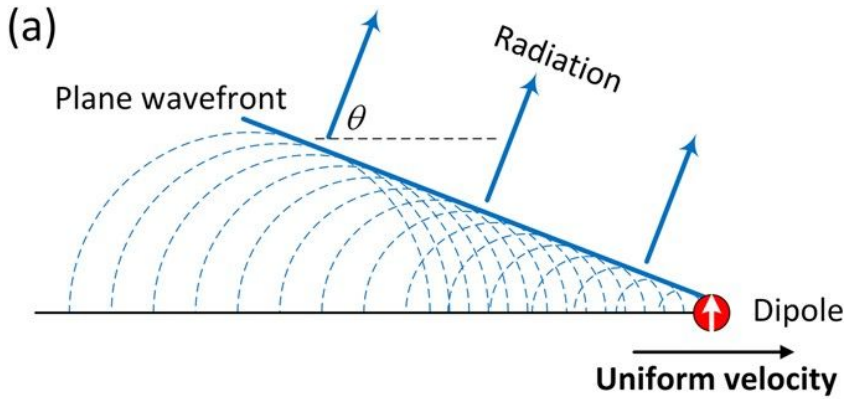




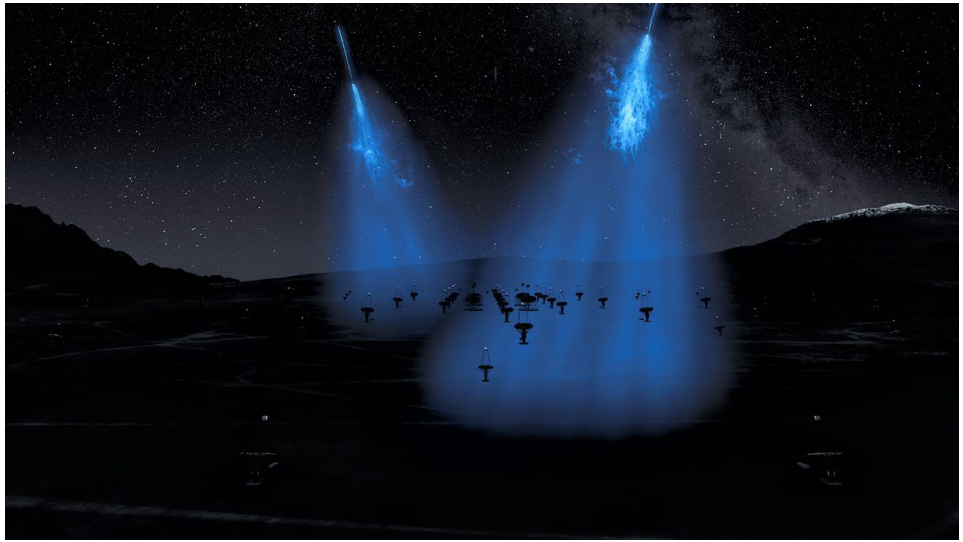
https://www.nature.com/articles/s41598-017-08705-4

Čerenkovovo záření

- Nabitá částice překračuje rychlost světla v hmotném prostředí
  - ideálně transparentní médium
- $v_{světla}$ = c / index lomu = c / n

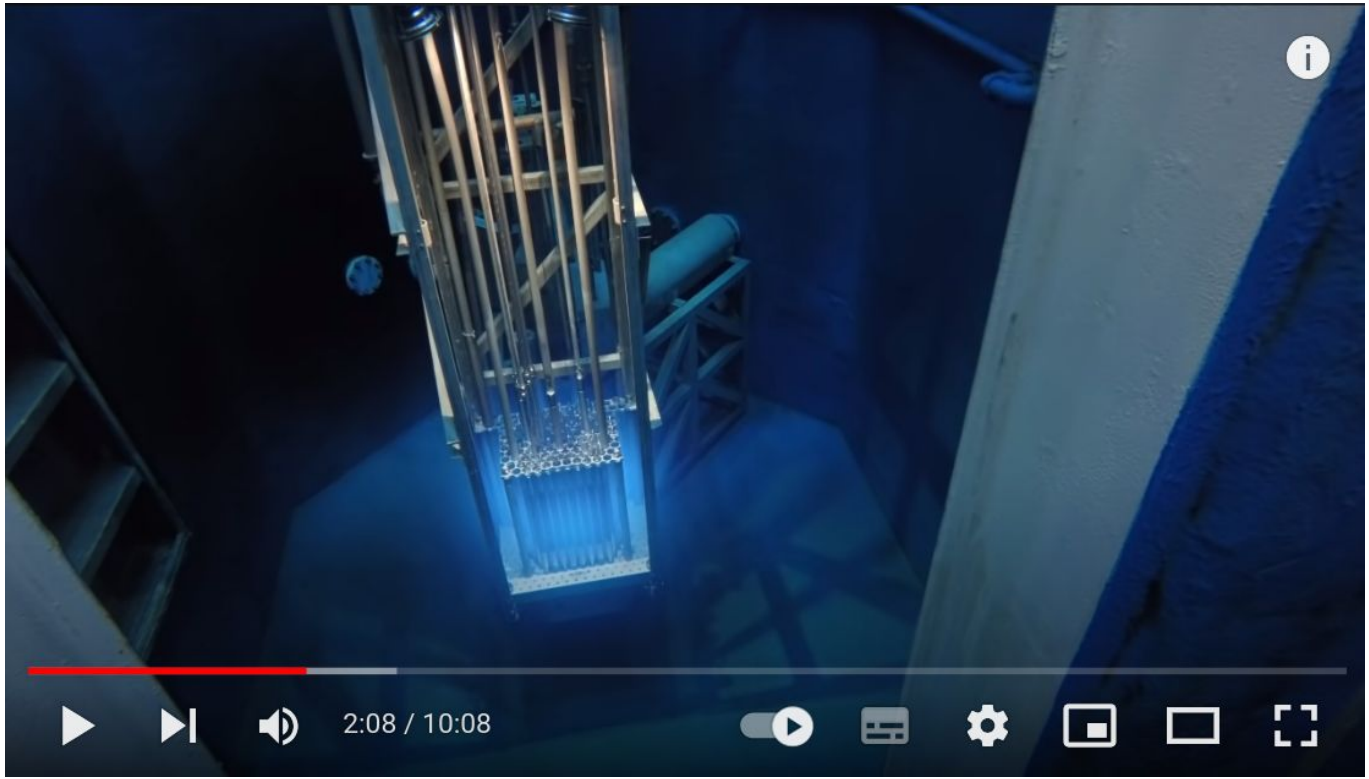


https://www.nature.com/articles/s41598-017-08705-4



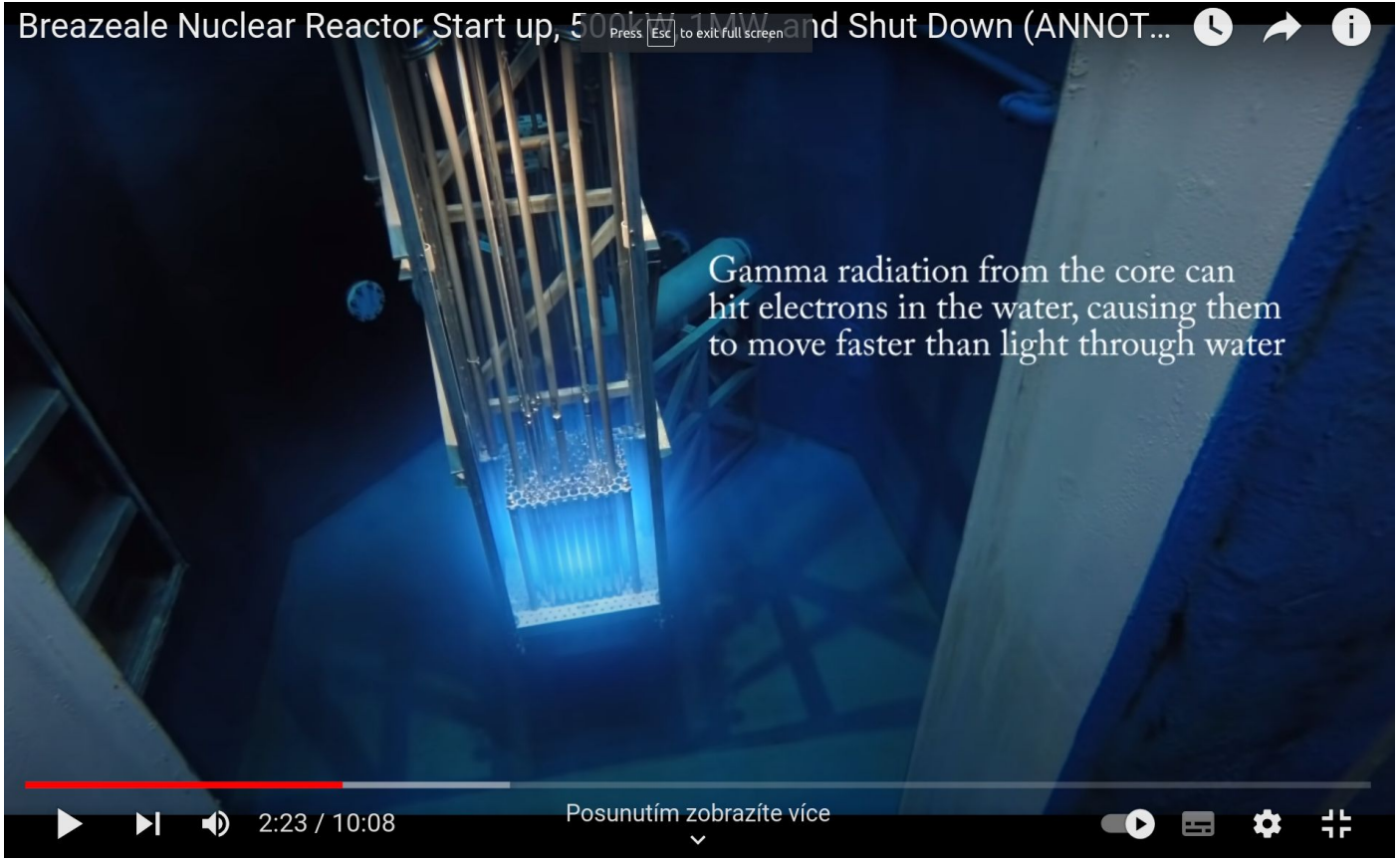
https://www.eso.org/public/czechrepublic/images/eso1841i/

Čerenkovovo záření



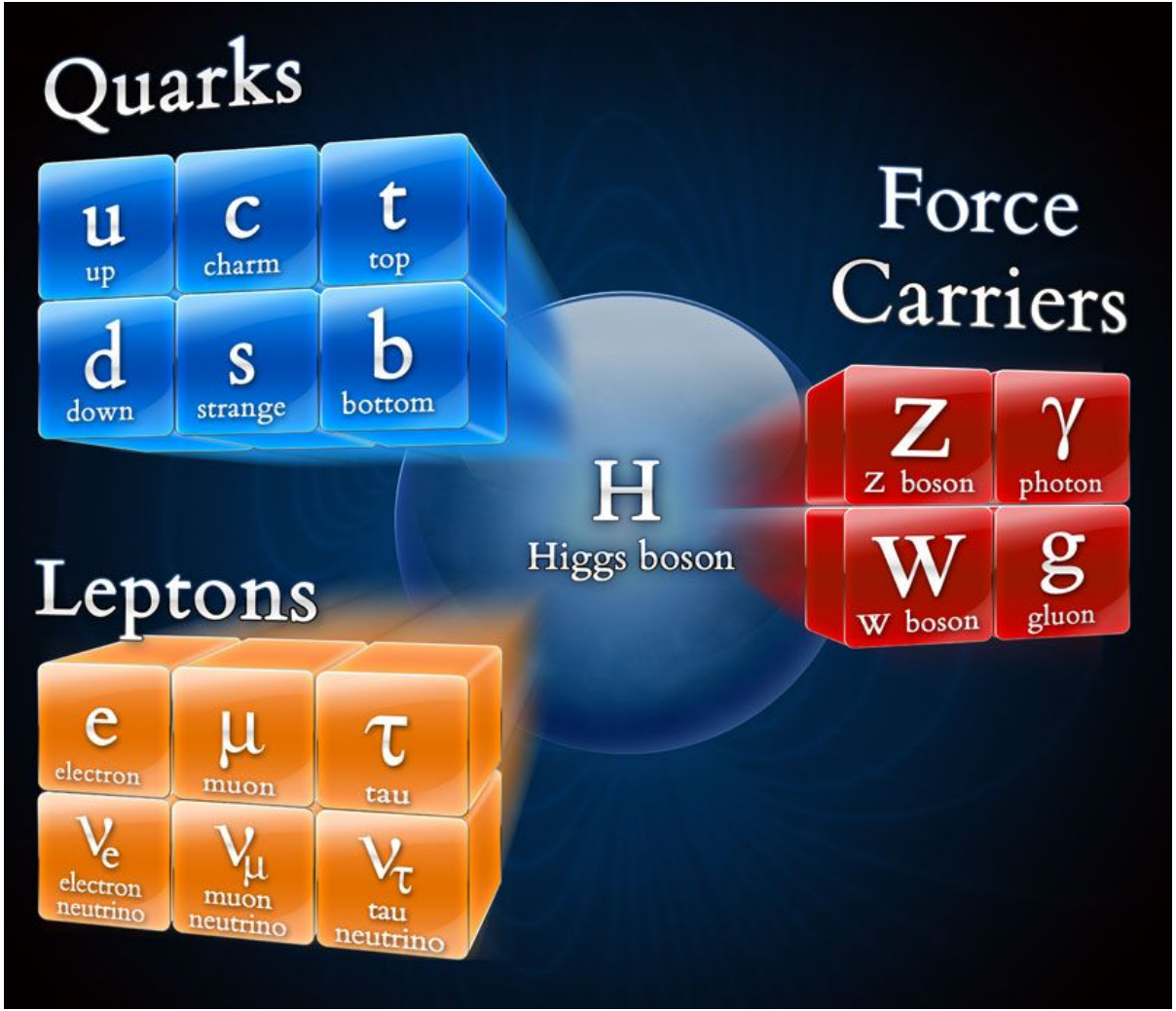
https://www.youtube.com/watch?v=uYrhWO_ZLYw&t=516s&ab_channel=AlexLandress

Čerenkovovo záření



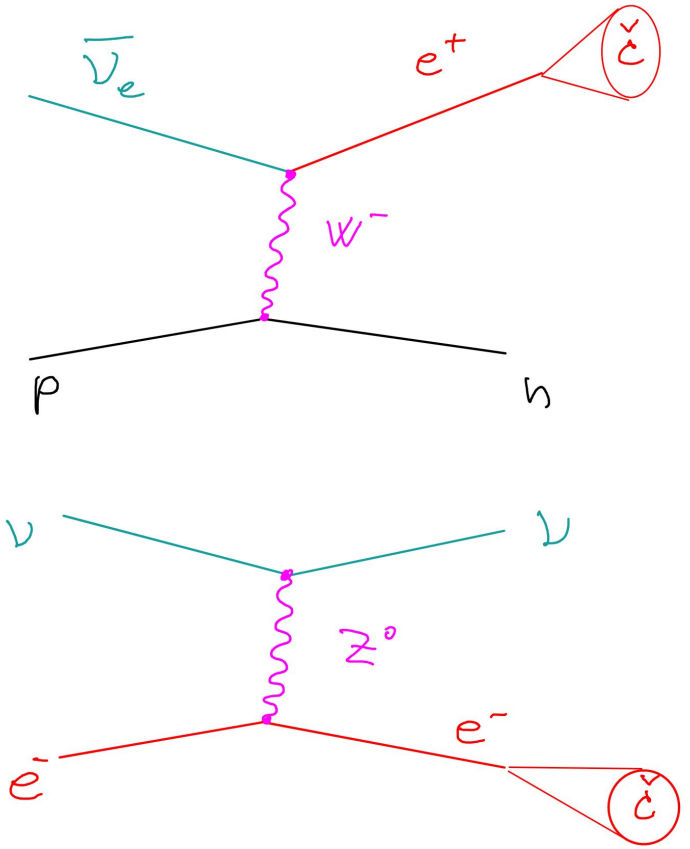
https://www.youtube.com/watch?v=uYrhWO_ZLYw&t=516s&ab_channel=AlexLandress

Částice

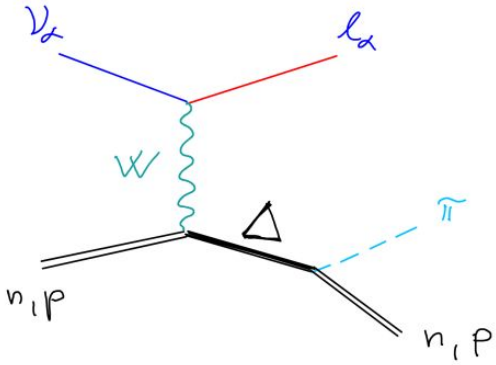

Quarks
u
up
c
charm
t
top
d
down
s
strange
b
bottom
Force Carriers
Z
Z boson
γ
photon
W
W boson
g
gluon
H
Higgs boson
Leptons
e
electron
μ
muon
τ
tau
νe
electron neutrino
νμ
muon neutrino
ντ
tau neutrino

Interakce neutrin

- Měříme částice, které v interakcích vzniknou
- Druh neutrina je spjat s druhem nabitého leptonu



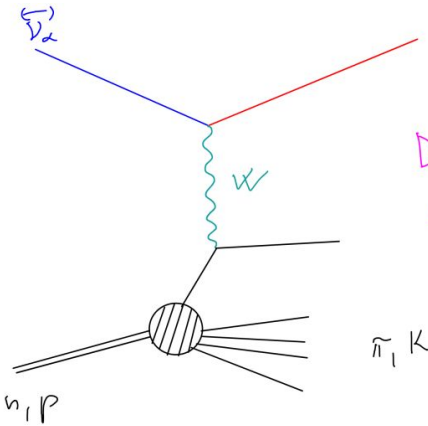
Interakce neutrin



Resonant pion production

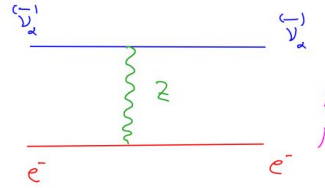
CC 1π

Charged Current + 1π Production



π, K, n, ...

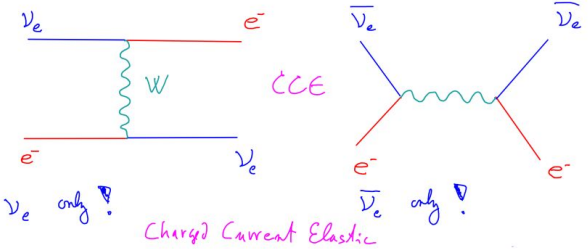
Deep Inelastic Scattering (DIS)



$\alpha = e, \mu, \tau$

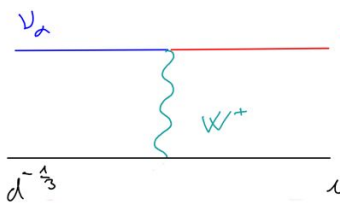
NC e^- ... in matter

Neutral Current



CCE

Charged Current Elastic



$\ell_\alpha^- = e^-, \mu^-, \tau^-$

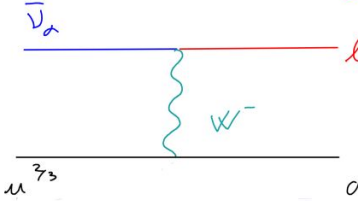
Č Effekt in H_2O

CCQE

$\nu_\alpha n \rightarrow \ell_\alpha^- p^+$

inverse β-decay

Charged Current Quasi-Elastic



$\bar{\nu}_\alpha$

ℓ_α^+ $\alpha = e, \mu, \tau^+$

CCQE

$\bar{\nu}_\alpha p \rightarrow \ell_\alpha^+ n$

$n\ Gd \rightarrow Gd^* \rightarrow Gd\ \gamma$

Experiment Kamiokande(-II) :: 3kt H_20

$p^+ \rightarrow e^+ \pi^0$

$p^+ \rightarrow \mu^+ \pi^0$

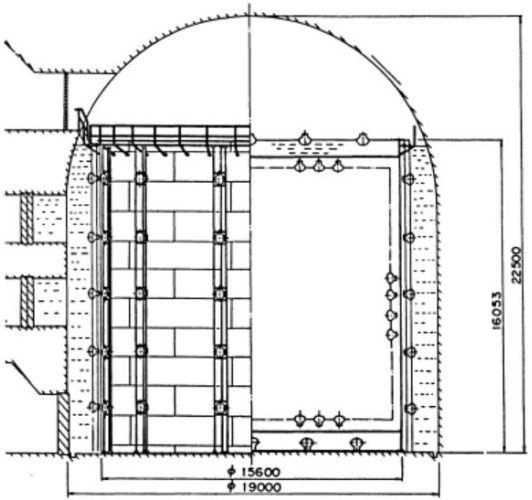
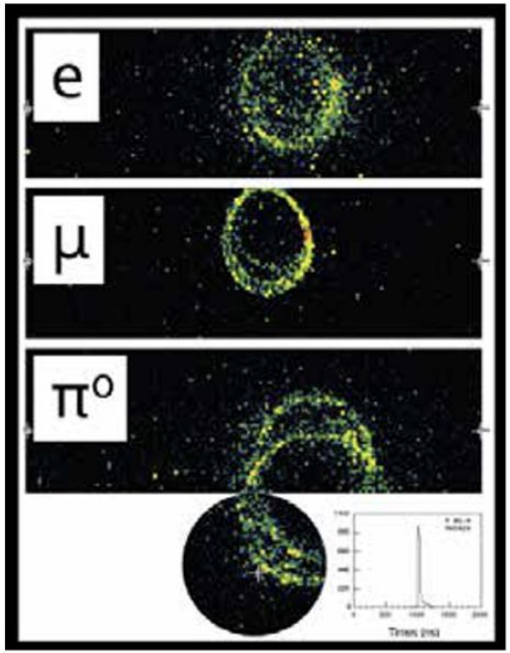


Figure 2: The detector of KAMIOKANDE -II. The dimensions are given in millimeters.



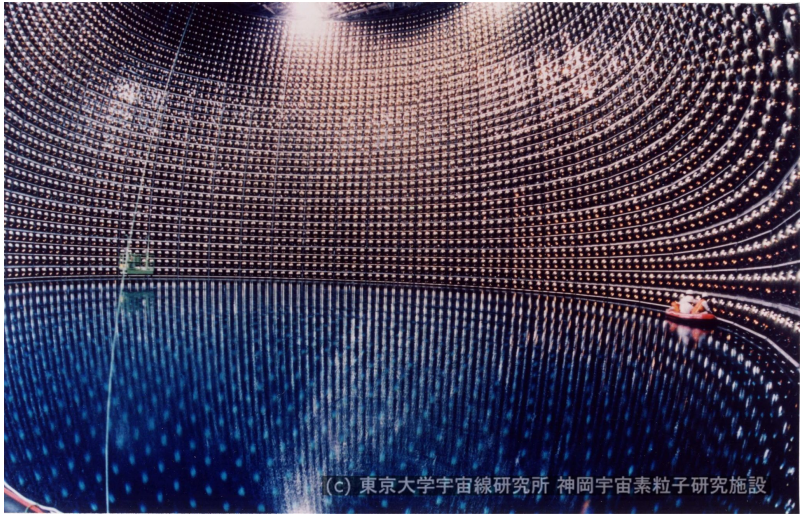
Neutrino identification

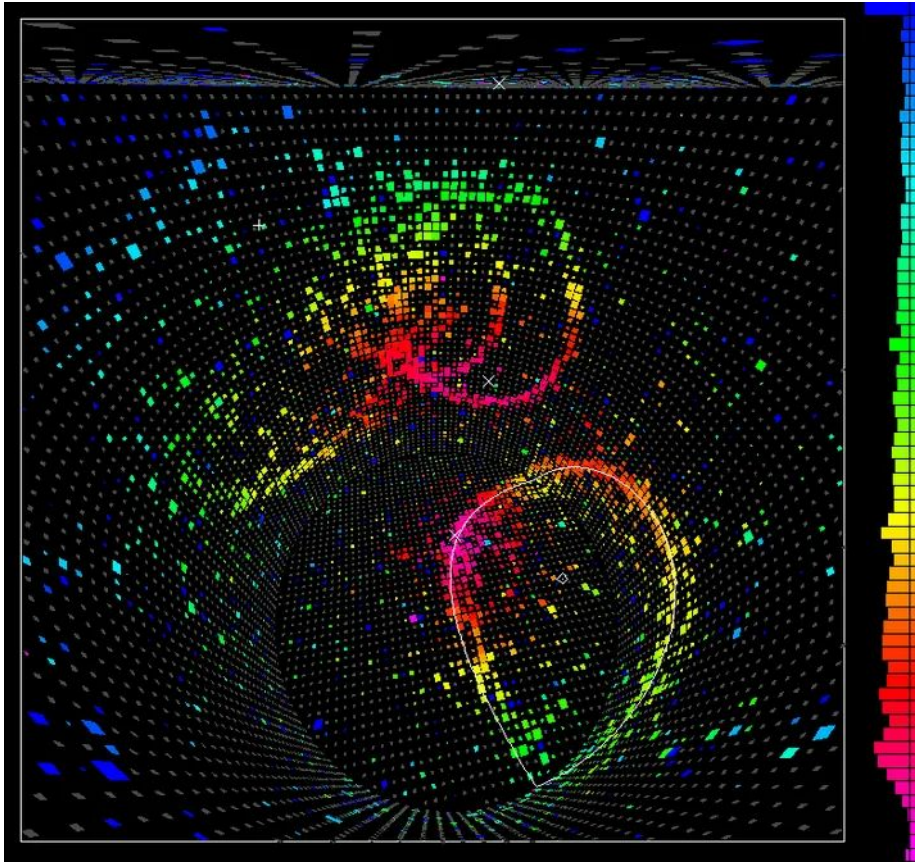
https://articles.adsabs.harvard.edu//full/1987ESOC...26..219K/0000220.000.html

Super Kamiokande :: 50kt H_20
Kamioka Nucleon Decay Experiment

https://www-sk.icrr.u-tokyo.ac.jp/en/sk/about/history/

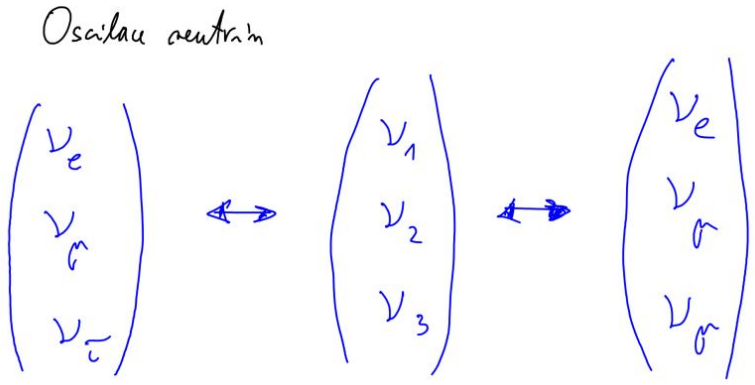
https://www-sk.icrr.u-tokyo.ac.jp/en/sk/experience/gallery/
https://www.forbes.com/sites/startswithabang/2018/12/05/is-there-really-a-fourth-neutrino-out-there-in-the-universe/

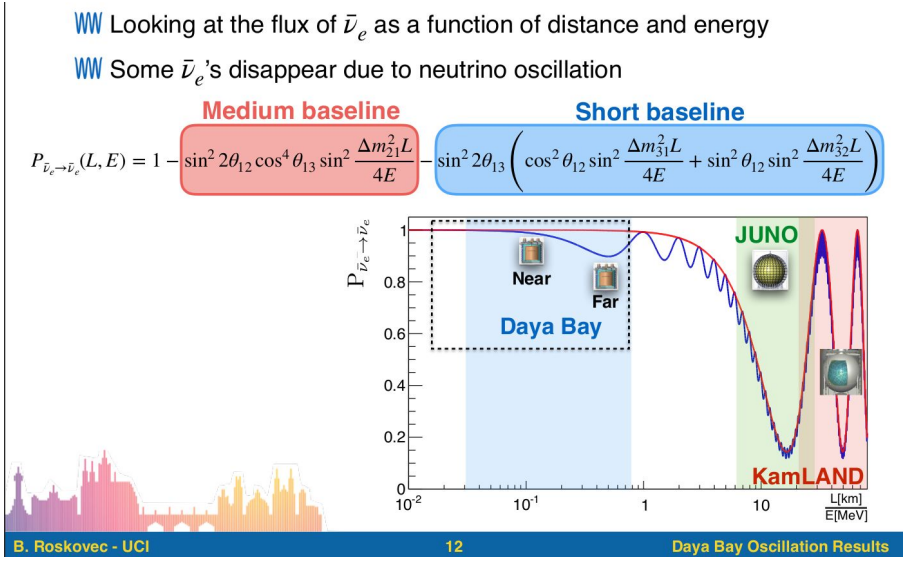




Oscilace neutrin

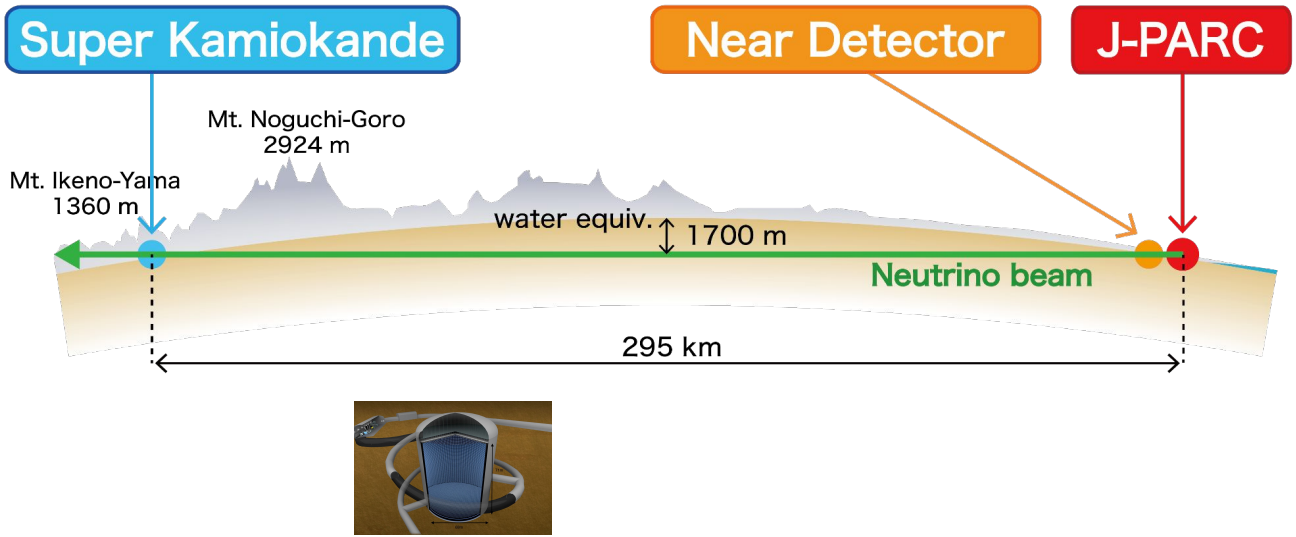
- Neutrino původního druhu s “letem” a interakcí v hmotě mění svůj typ.
- Neutrinová přecházejí v mionová apod
- Druh neutrina poznáme v interakci
- S elektronovým se pojí elektronové apod
- Oscilace tak lze pozorovat a měřit
- Důležitá část Standardního modelu
- Dochází k oscilacím podle současné teorie?





Super Kamiokande a neutrina z urychlovačů

- Oscilace neutrin zkoumány i s uměle vytvořenými neutriny ze svazků částic na urychlovačích.



Super Kamiokande

https://www-sk.icrr.u-tokyo.ac.jp/en/sk/experience/gallery/
https://www.forbes.com/sites/startswithabang/2018/12/05/is-there-really-a-fourth-neutrino-out-there-in-the-universe/

The Nobel Prize in Physics 2002

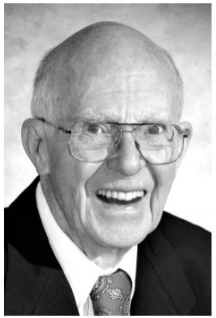


Photo from the Nobel Foundation archive.
Raymond Davis Jr.
Prize share: 1/4

Photo from the Nobel Foundation archive.
Masatoshi Koshiba
Prize share: 1/4

Photo from the Nobel Foundation archive.
Riccardo Giacconi
Prize share: 1/2

The Nobel Prize in Physics 2002 was divided, one half jointly to Raymond Davis Jr. and Masatoshi Koshiba "for pioneering contributions to astrophysics, in particular for the detection of cosmic neutrinos" and the other half to Riccardo Giacconi "for pioneering contributions to astrophysics, which have led to the discovery of cosmic X-ray sources"

The Nobel Prize in Physics 2015

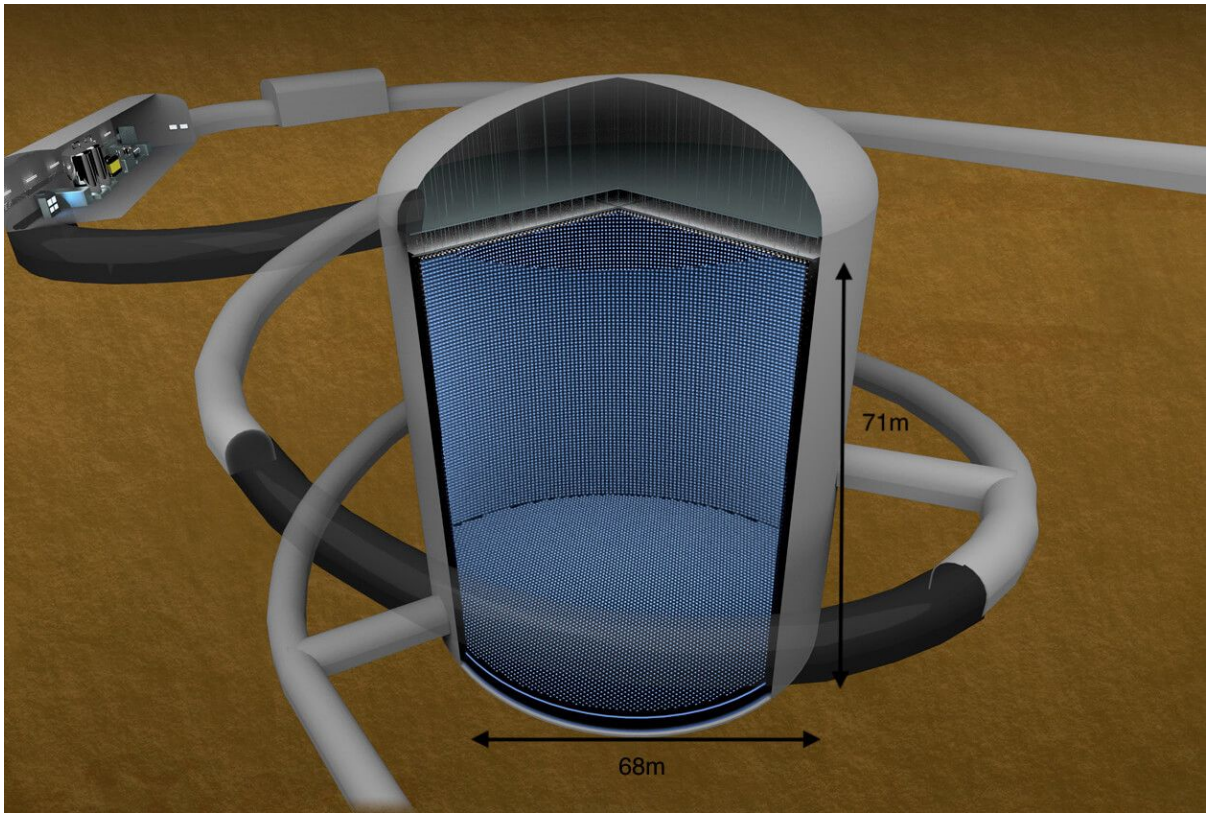
© Nobel Media AB. Photo: A. Mahmoud
Takaaki Kajita
Prize share: 1/2

© Nobel Media AB. Photo: A. Mahmoud
Arthur B. McDonald
Prize share: 1/2

The Nobel Prize in Physics 2015 was awarded jointly to Takaaki Kajita and Arthur B. McDonald "for the discovery of neutrino oscillations, which shows that neutrinos have mass"

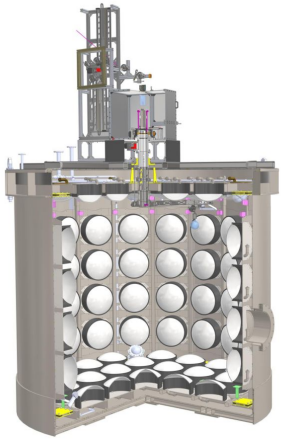
Hyper Kamiokande :: 250kt H_20

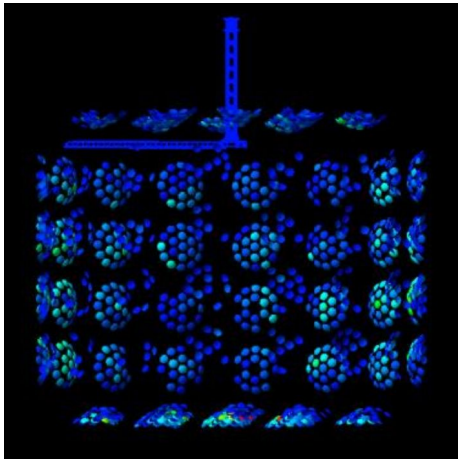
https://www-sk.icrr.u-tokyo.ac.jp/en/hk/about/detector/
https://www-sk.icrr.u-tokyo.ac.jp/en/hk/



YT :: Hyper-Kamiokande
https://www.youtube.com/watch?v=JFOE3D2z7LM&t=12s&ab_channel=Hyper-Kamiokande

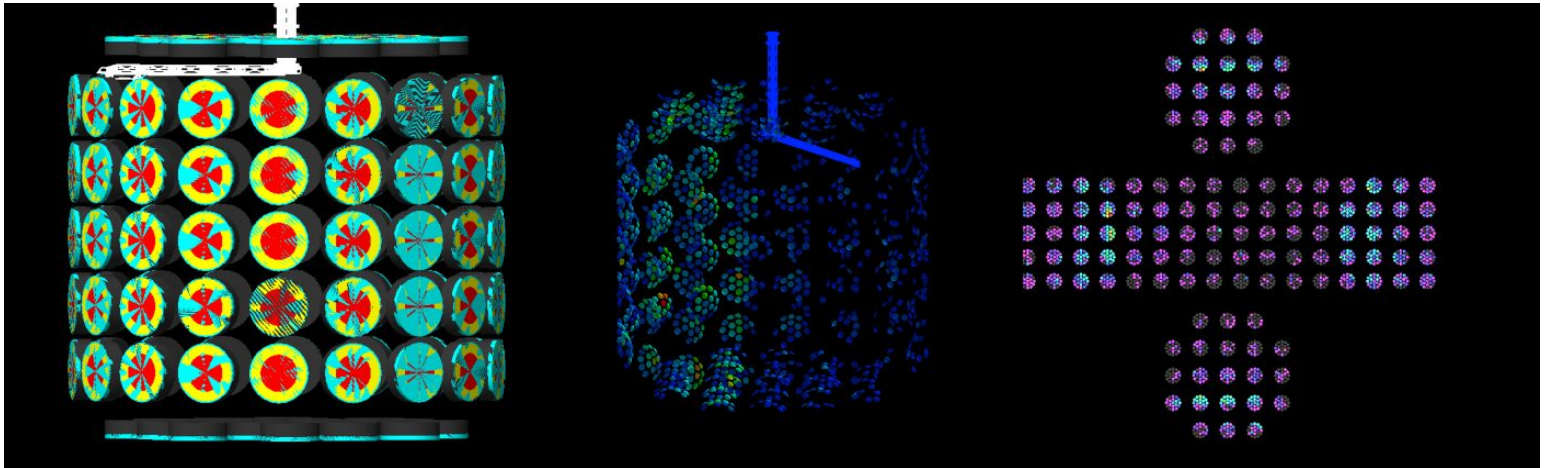
Water Cherenkov Test Experiment :: 50t H_2O @ CERN



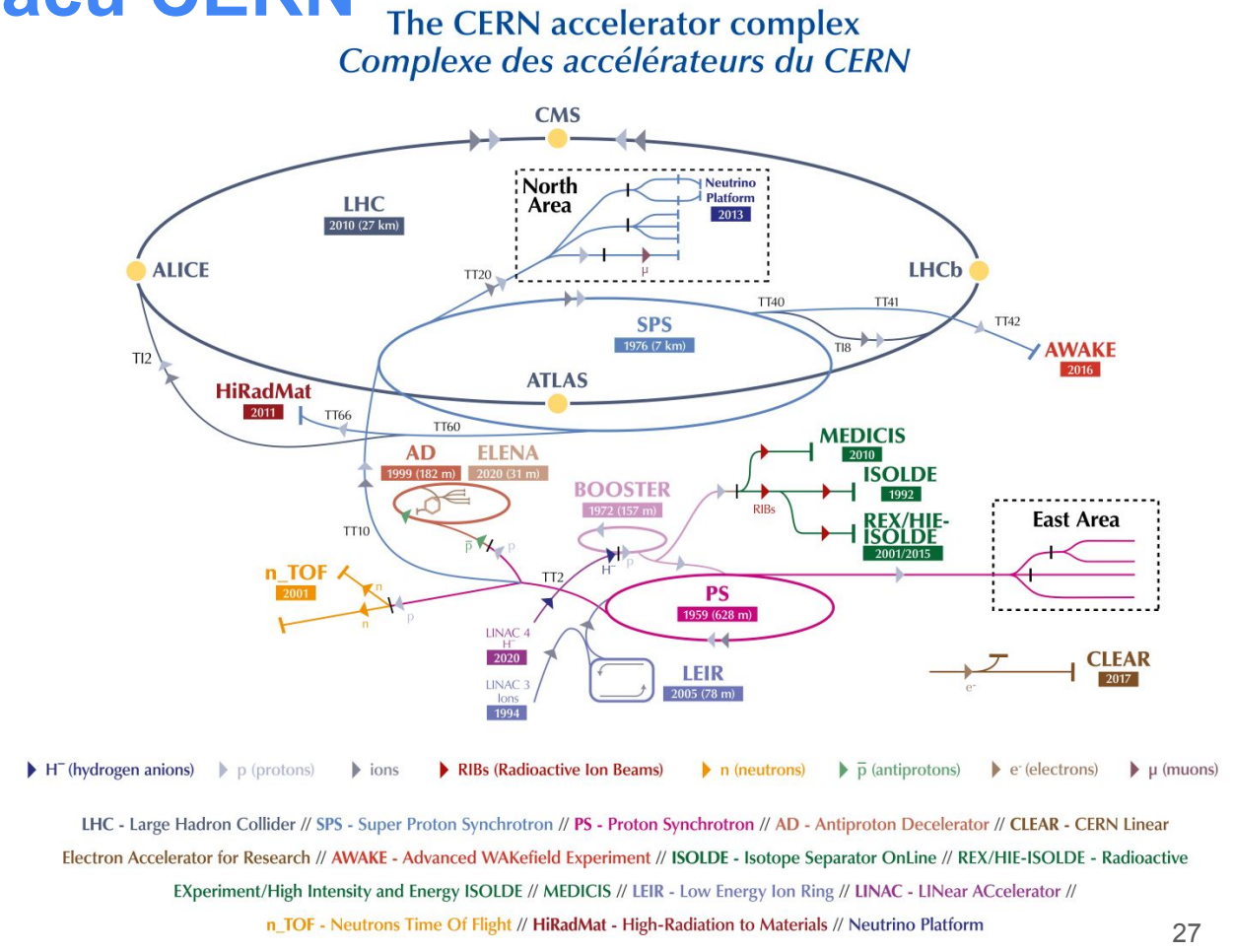


- Cíl: porozumět interakcím částic vznikajících v neutrinových reakcích.
- Pro kalibrace velkých neutrinových experimentů.

https://indico.cern.ch/event/855372/contributions/4452155/attachments/2304577/3921520/WCTE@CERN_NuFACT-2021-09-08_LAnthony.pdf

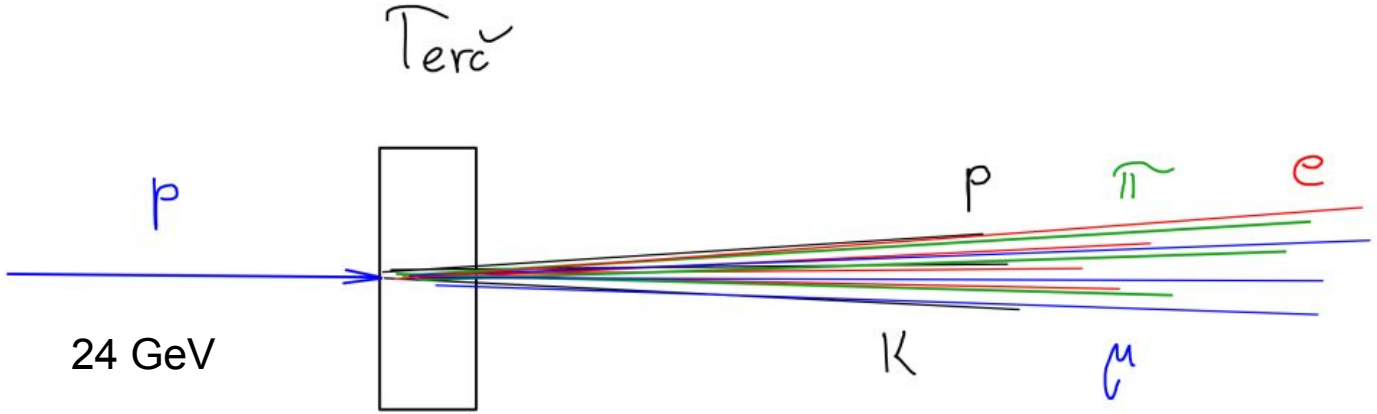


Komplex urychlovačů CERN



https://home.cern/science/accelerators/accelerator-complex

East Area Secondary Test Beams



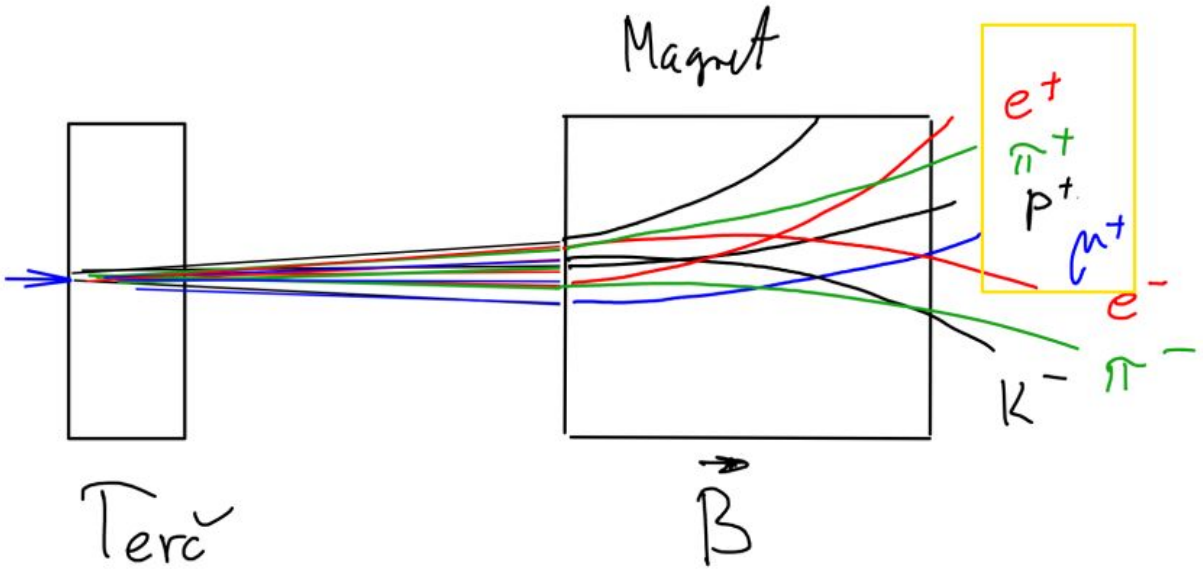
$|\vec{p}|$ 200 MeV ≃ 1 GeV

Analýza složení svazku pro WCTE

$$p = m v \qquad \beta \equiv \frac{v}{c}$$

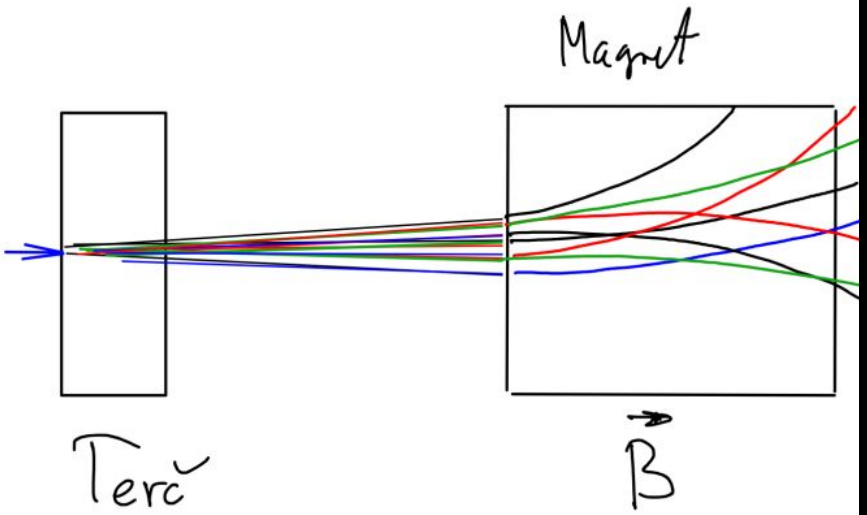
$$p = \gamma m v = \frac{1}{\sqrt{1 - \frac{v^2}{c^2}}} m v$$

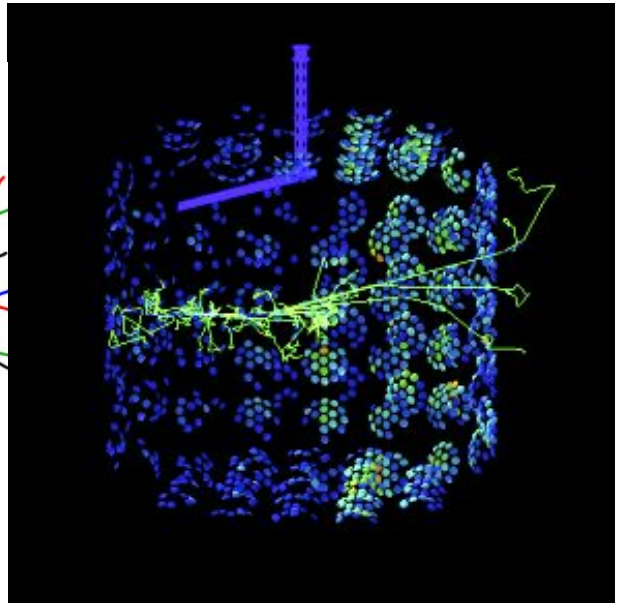
$$p = \beta \gamma mc = \frac{1}{\sqrt{1 - \beta^2}} \beta mc$$



Analýza složení svazku pro WCTE

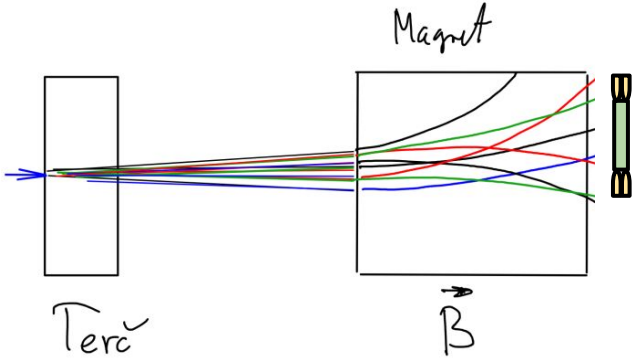
- Bude ale v sekundárním svazku dost pionů/mionů na kalibraci odezvy detektoru na tyto částice?
- Potřeba nejprve změřit složení svazku!
- Identifikace částic.



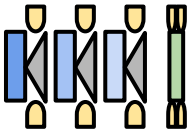


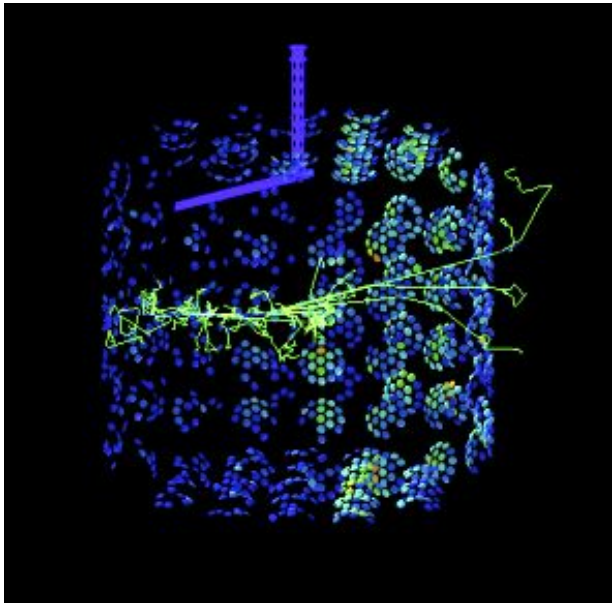
Analýza složení svazku pro WCTE

- Bude ale v sekundárním svazku dost pionů/mionů na kalibraci odezvy detektoru na tyto částice?
- Potřeba nejprve změřit složení svazku!
- Identifikace částic.

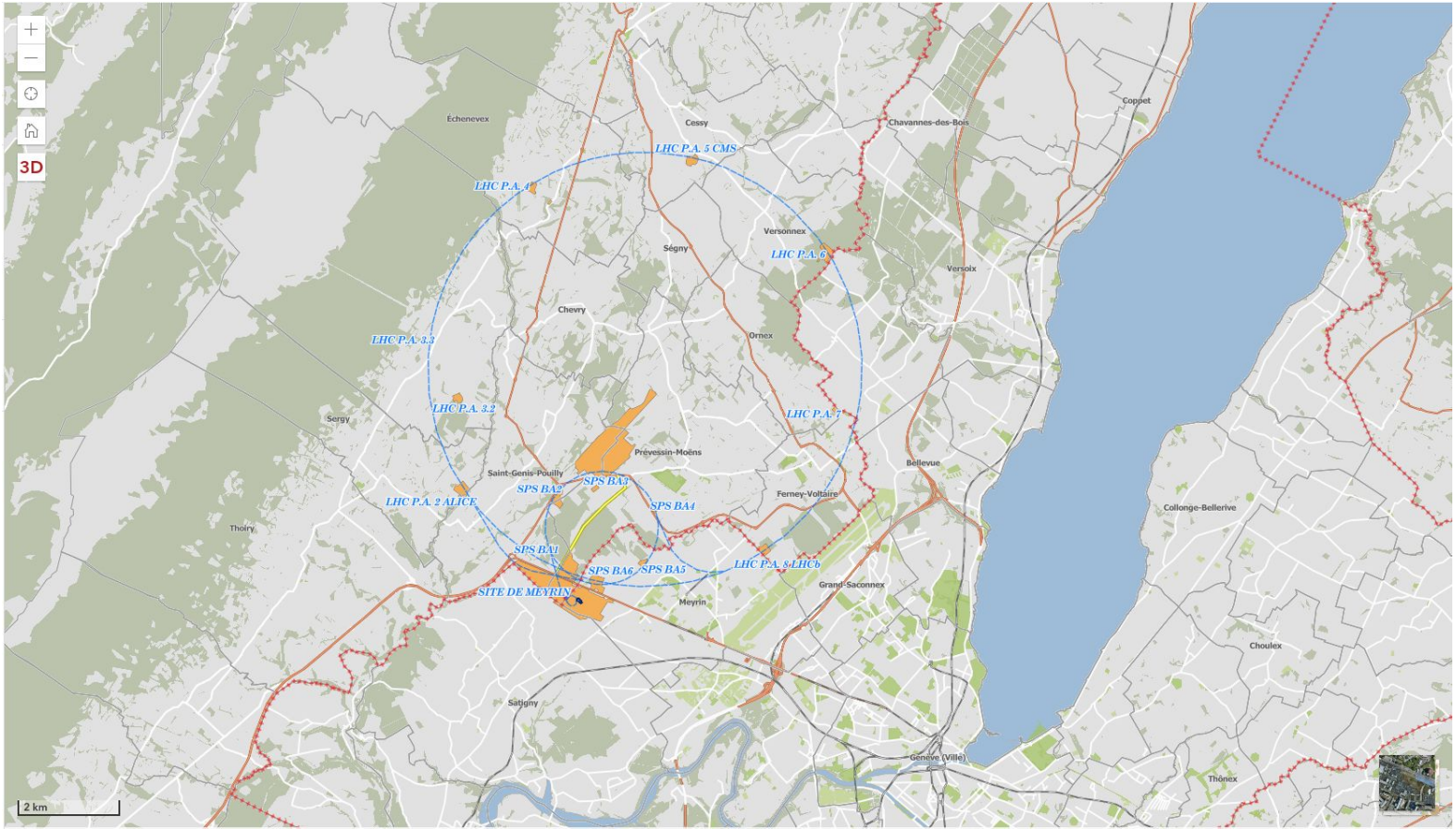


Monitor svazku





CERN :: https://maps.web.cern.ch/



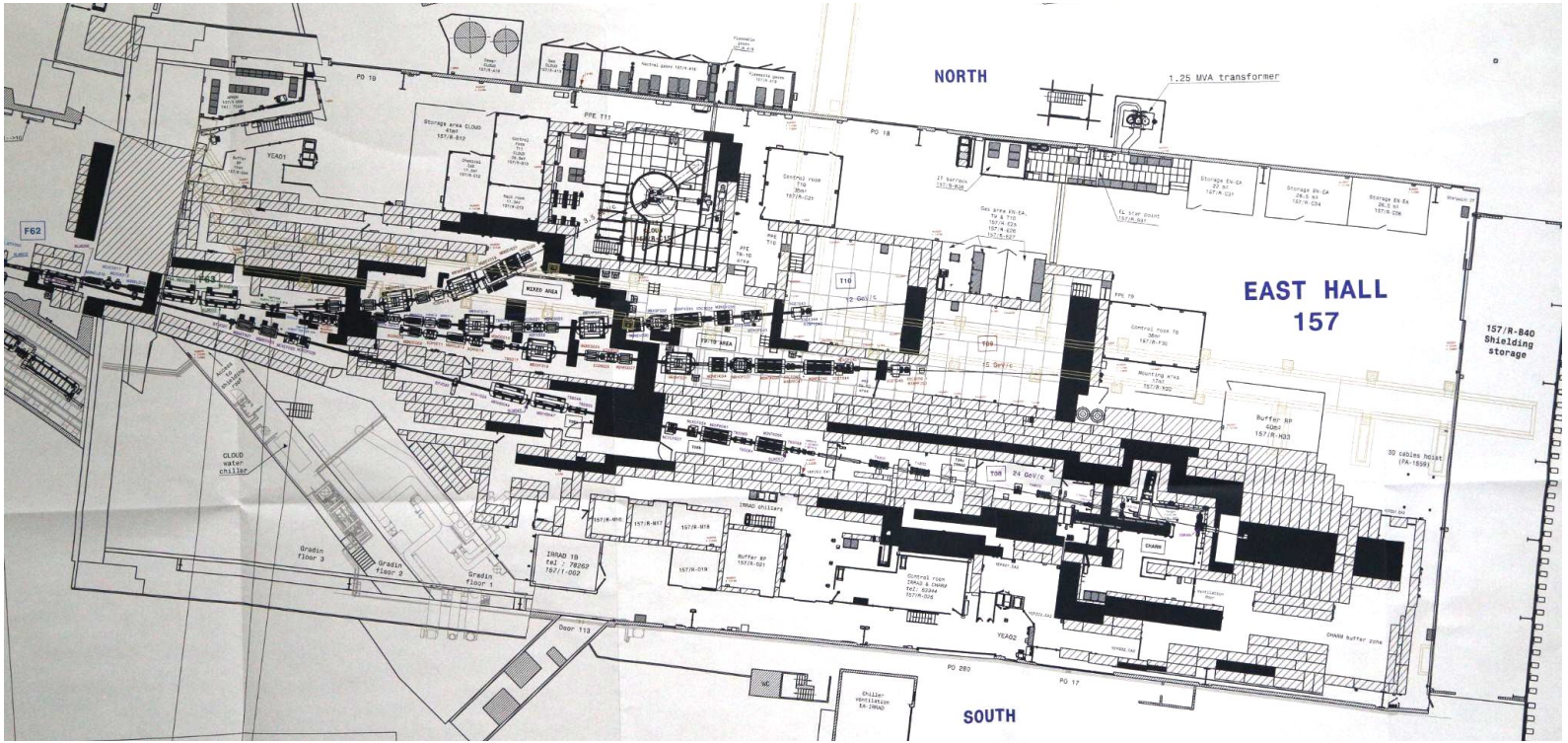
CERN :: https://maps.web.cern.ch/

East Area 2022

- Building 157 at CERN's Meyrin site, large experimental hall housing also the CLOUD experiment, T8, T9 and T10 beam lines.

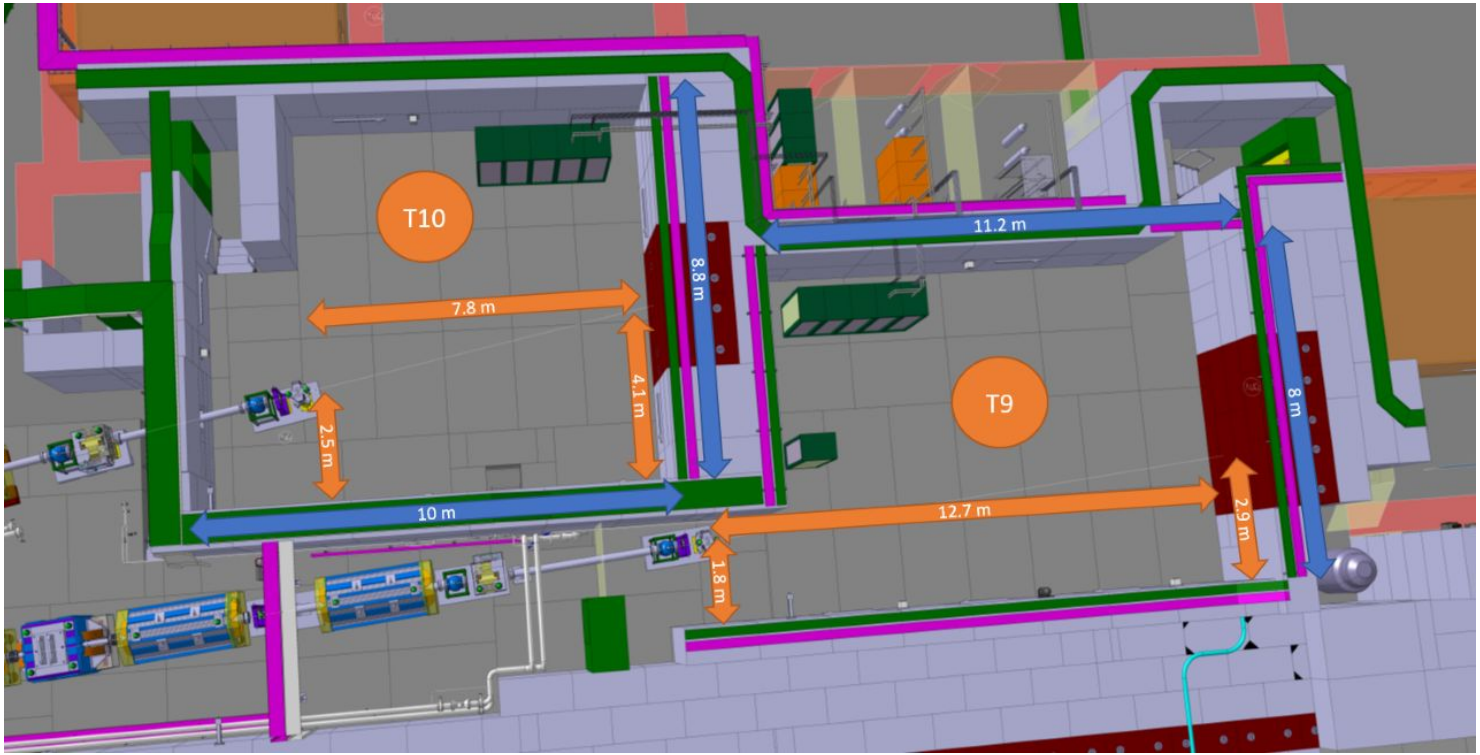
East Area

- Building 157 at CERN's Meyrin site, large experimental hall housing also the CLOUD experiment, T8, T9 and T10 beam lines.



East Area

- Building 157 at CERN's Meyrin site, large experimental hall housing also the CLOUD experiment, T8, T9 and T10 beam lines.



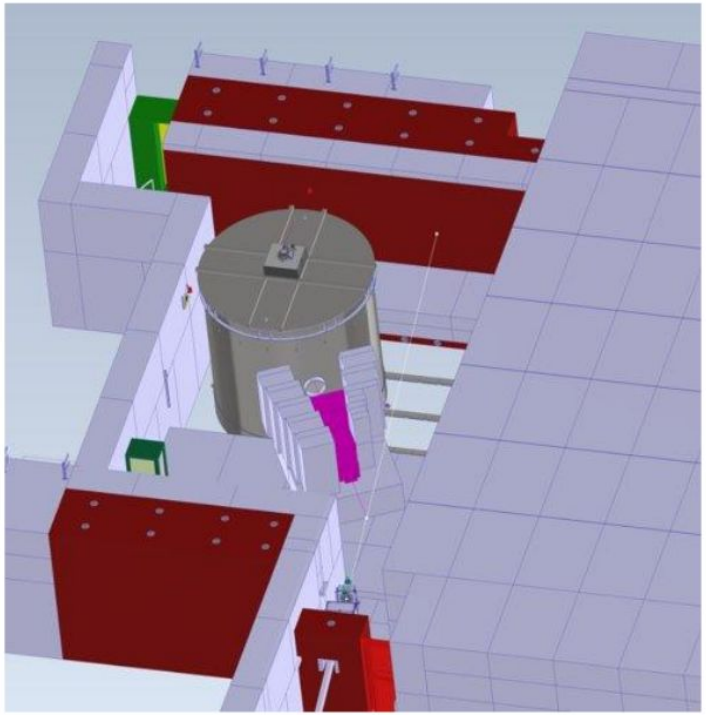
East Area

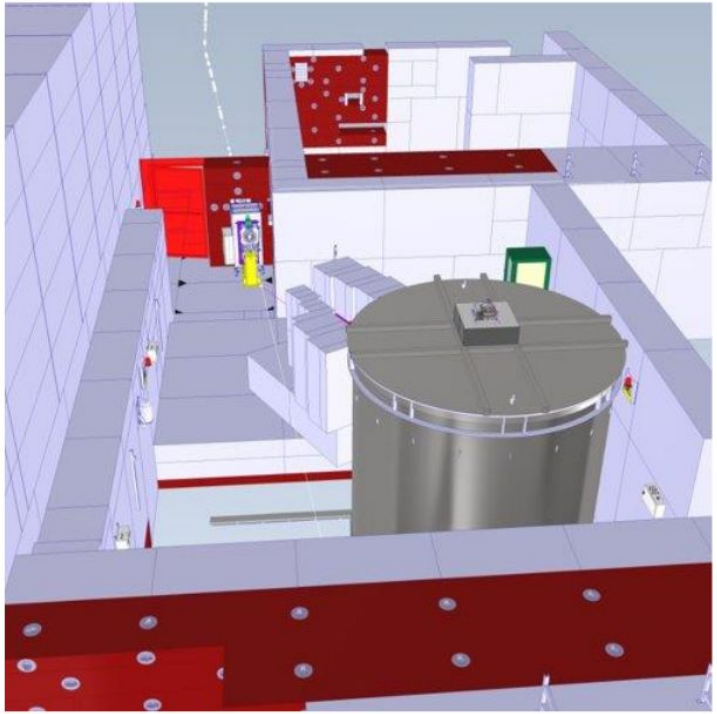
- Building 157 at CERN's Meyrin site, large experimental hall housing also the CLOUD experiment, T8, T9 and T10 beam lines.

East Area

East Area

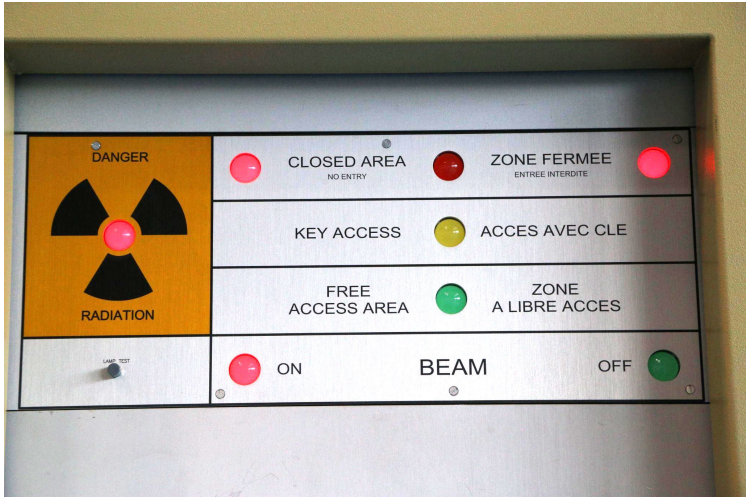
- Building 157 at CERN’s Meyrin site, large experimental hall housing also the CLOUD experiment, T8, T9 and T10 beam lines.





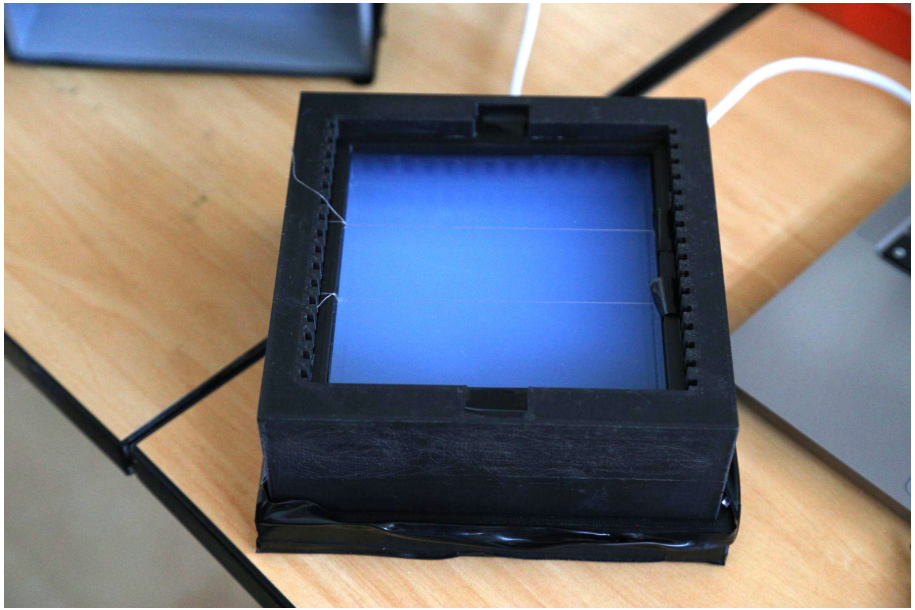
East Area 2022

- Building 157 at CERN's Meyrin site, large experimental hall housing also the CLOUD experiment, T8, T9 and T10 beam lines.



Installation :: Aerogel Cherenkov Counters

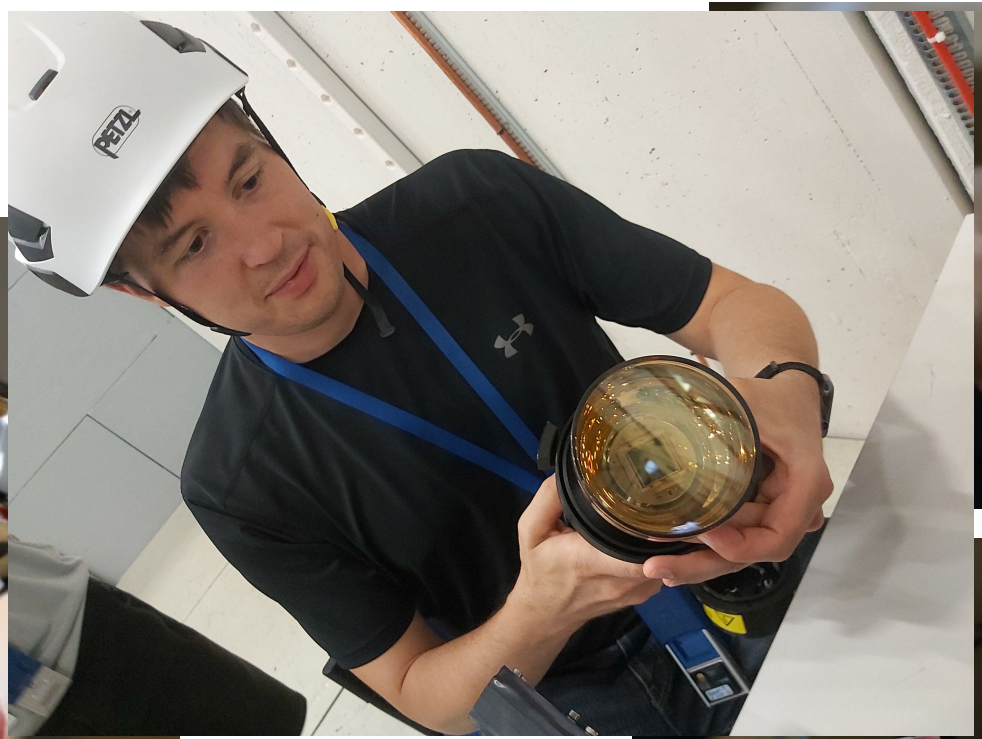
- Transparent aerogel produces Cherenkov radiation when a particle with velocity above the speed of light passes through.
- **World's Lightest Solid!**
  - https://www.youtube.com/watch?v=AeJ9q45PfD0&ab_channel=Veritasium
- Light is reflected by thin foil mirrors to PMT on both sides.

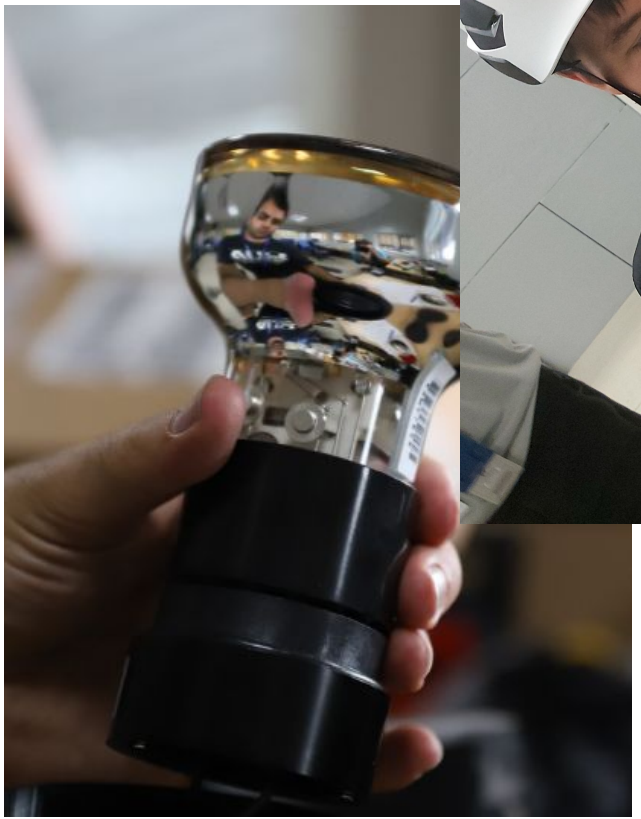


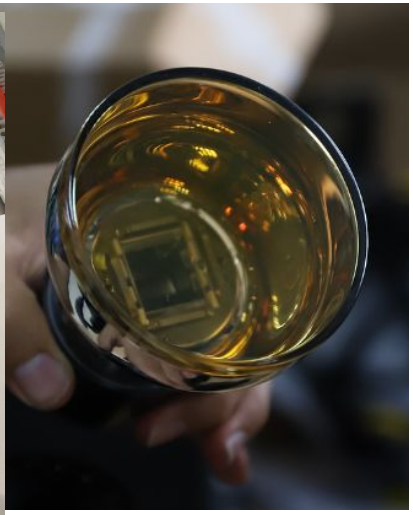
Aerogel

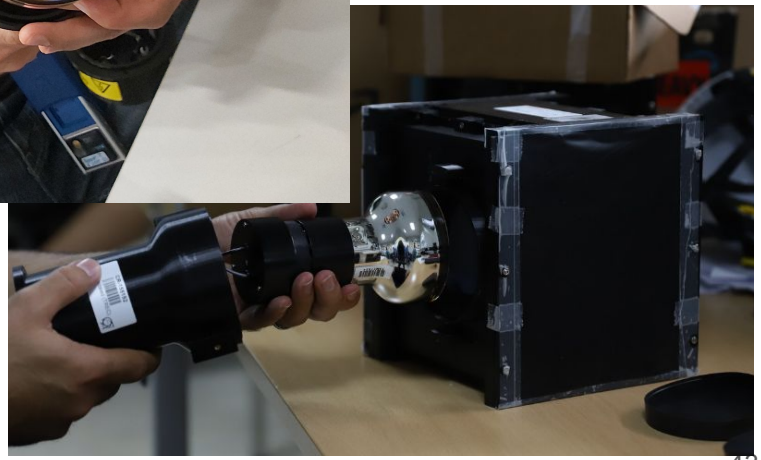
- **World's Lightest Solid!**
  - https://www.youtube.com/watch?v=AeJ9q45PfD0&ab_channel=Veritasium

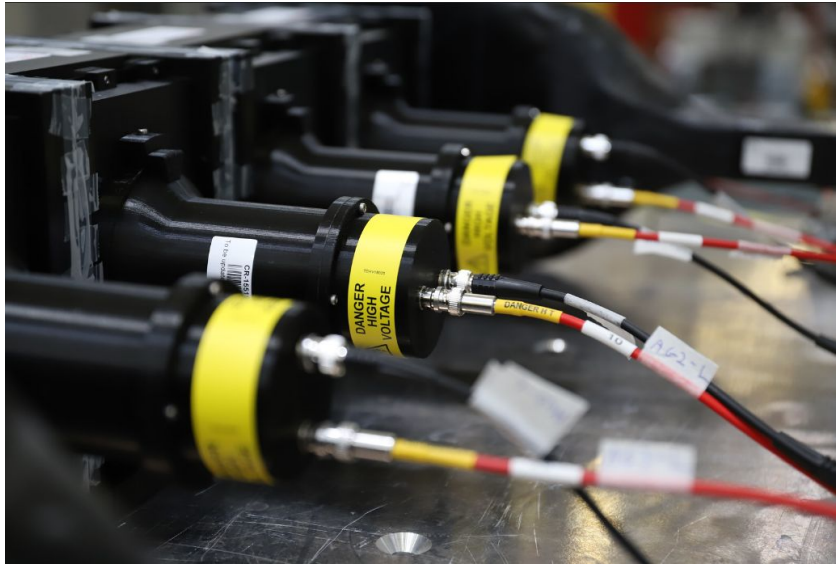
Fotonásobič

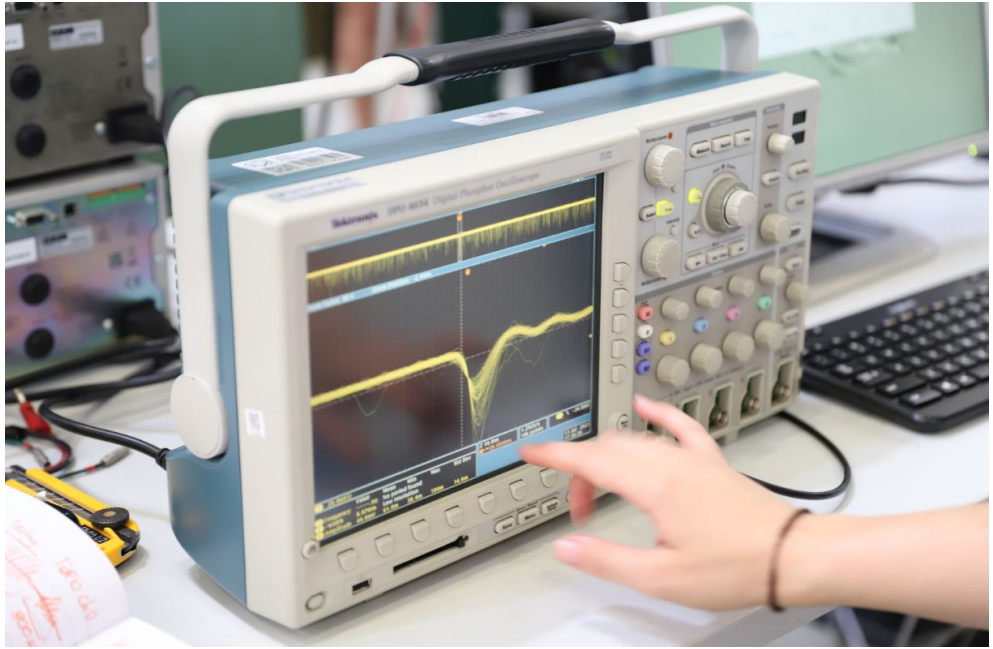


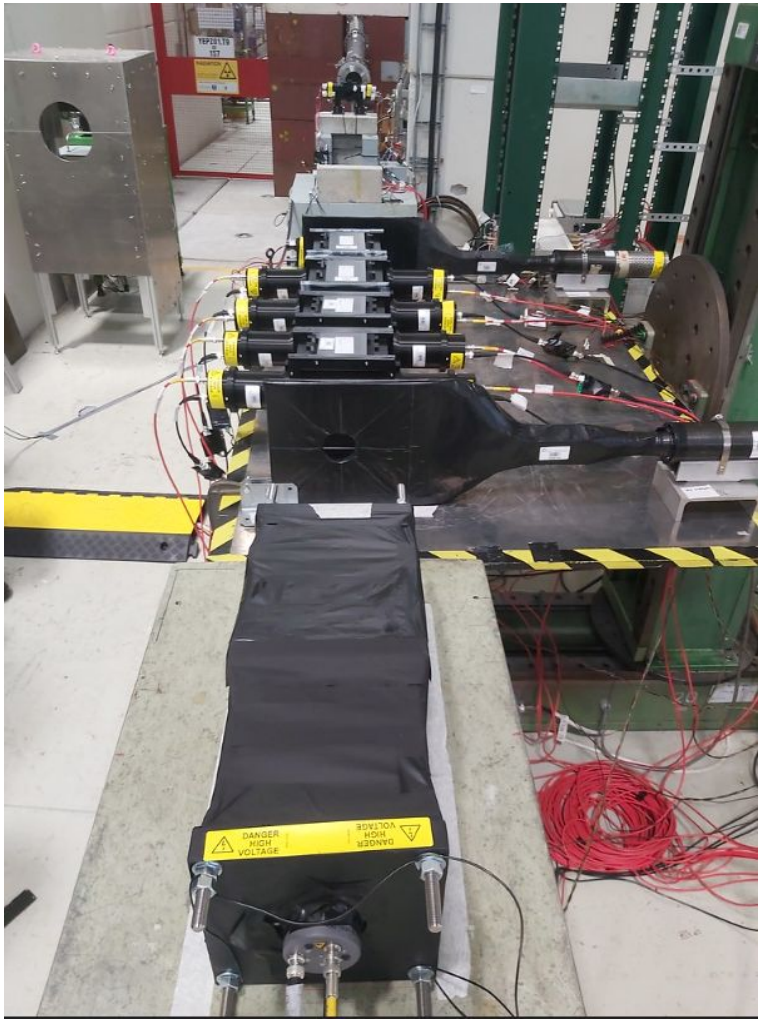







DANGER
HIGH
VOLTAGE



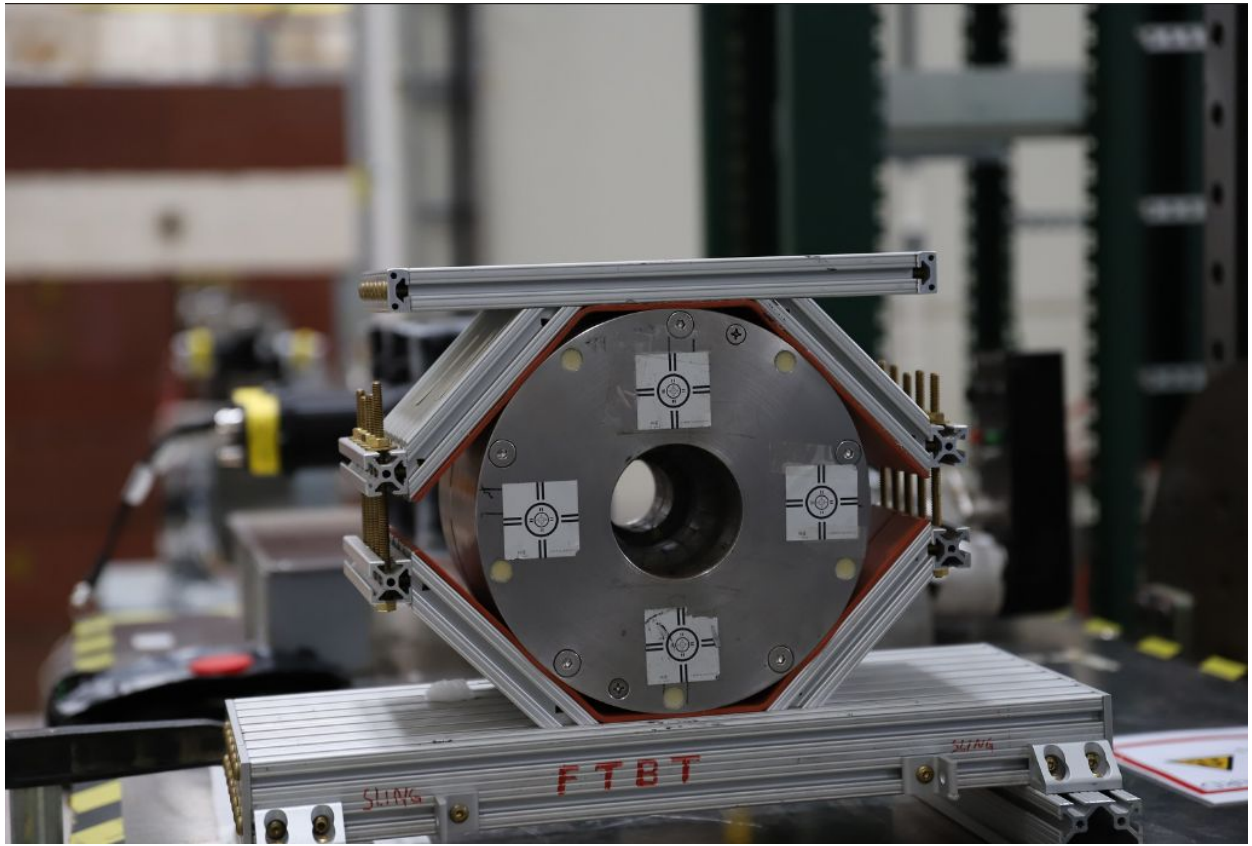

DANGER
HIGH
VOLTAGE


DANGER H T
15
Pb glass

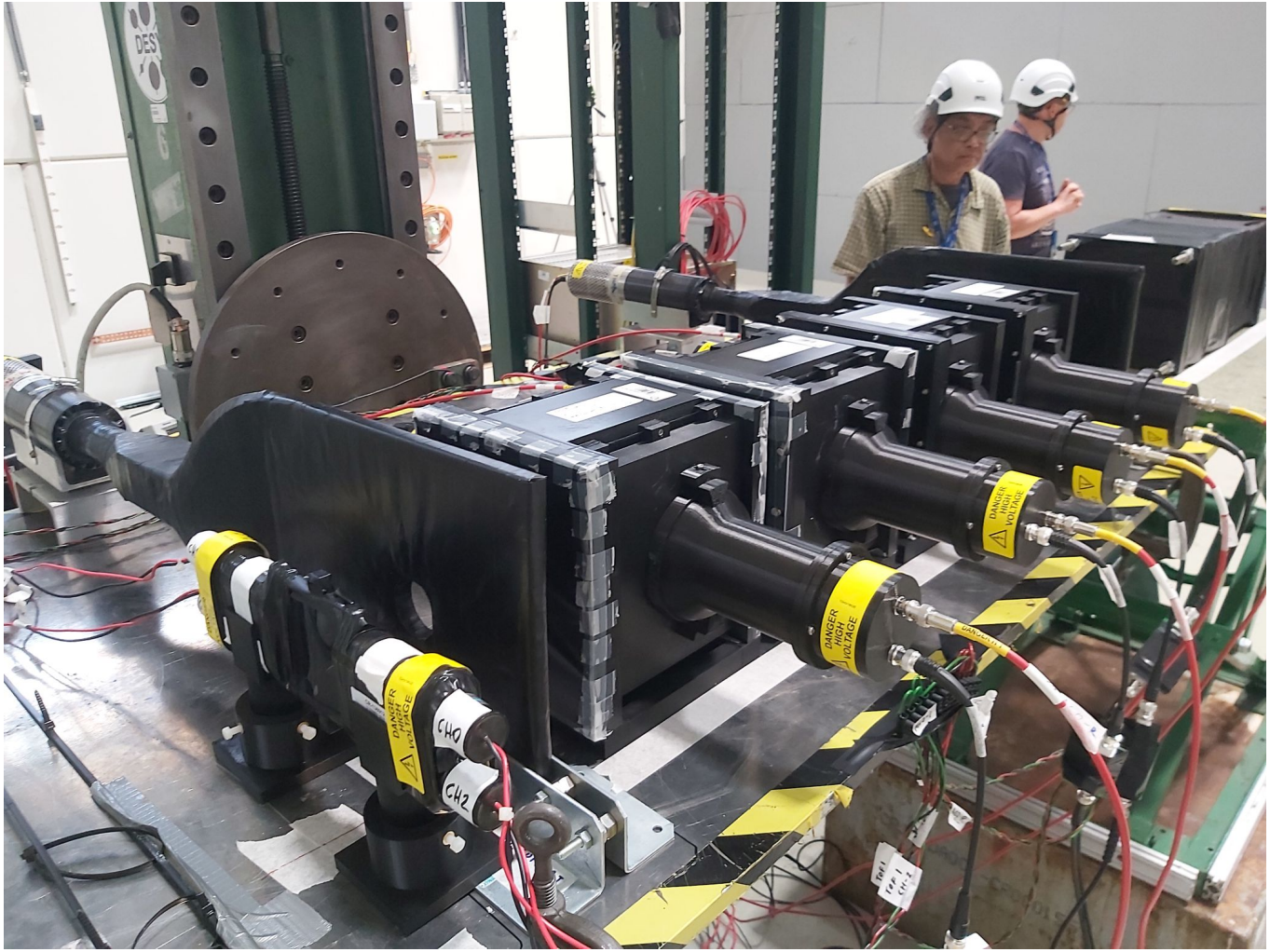
Fotonásobiče….A jejich výměna:)

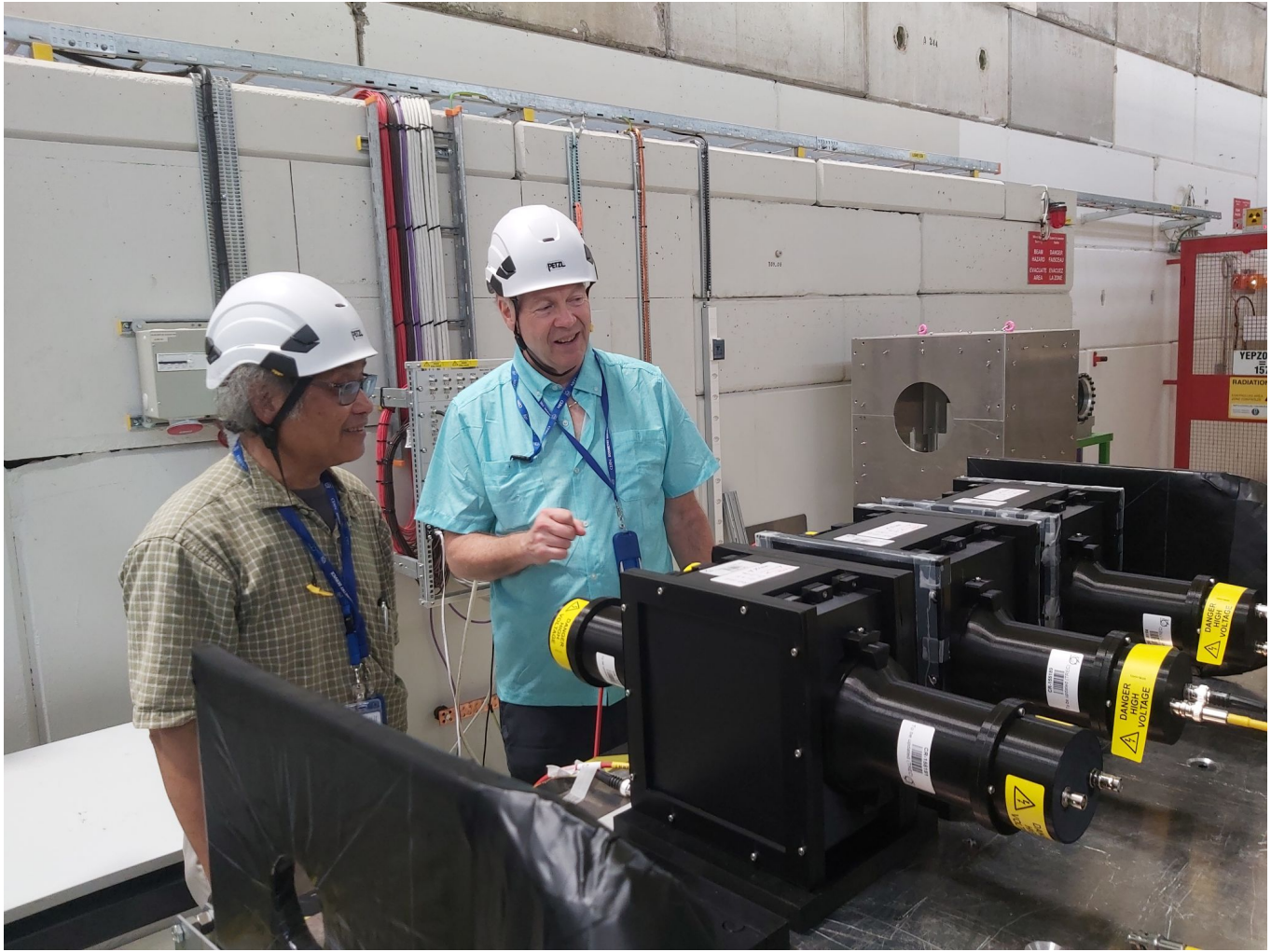
https://www.youtube.com/watch?v=AUMeRy8Z7OU&ab_channel=JiriKvita

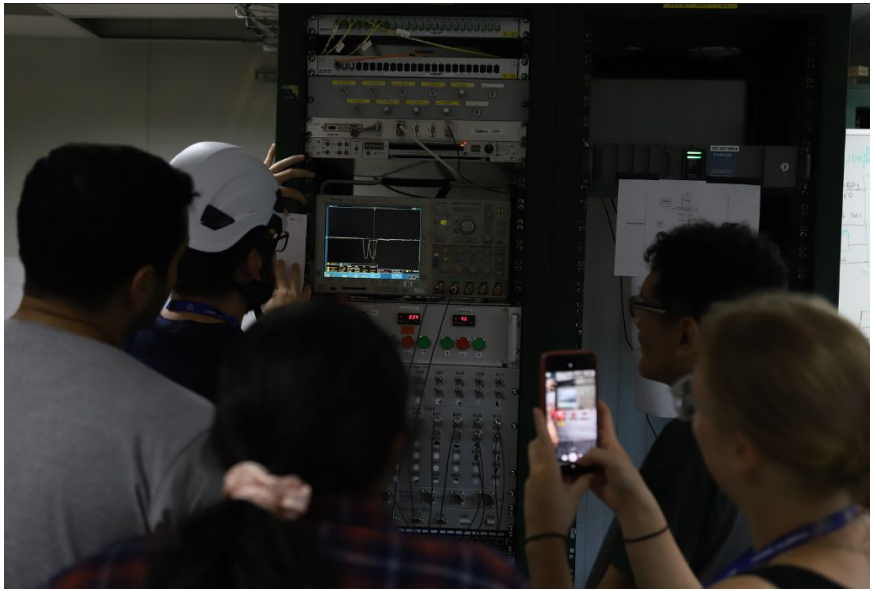
Meet Mr. Permanent Magnet of 1.7T!

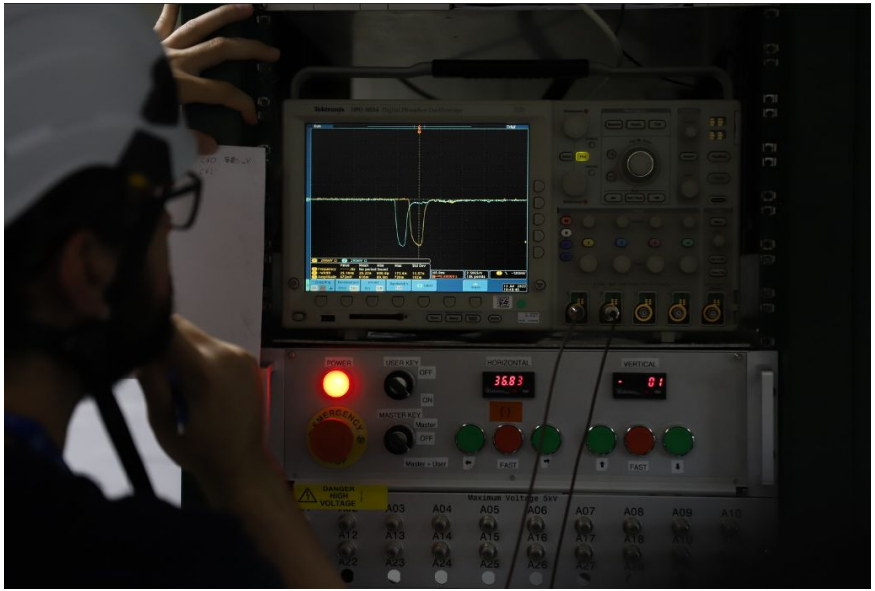
- 


SEARCH
BUA
INCINERABLE RADIOACTIVE WASTE
DECHET INCINERABLE RADIOACTIF
ELECTRONIC & CABLE RADIOACTIVE WASTE
DECHET ELECTRONIQUE + CABLE RADIOACTIF
METALLIC RADIOACTIVE WASTE
DECHET METALLIQUE RADIOACTIF
CONVENTIONAL WASTE
DECHETS CONVENTIONNELS
72504


DANGER HIGH VOLTAGE
DANGER HIGH VOLTAGE
DANGER HIGH VOLTAGE
CH0
CH2

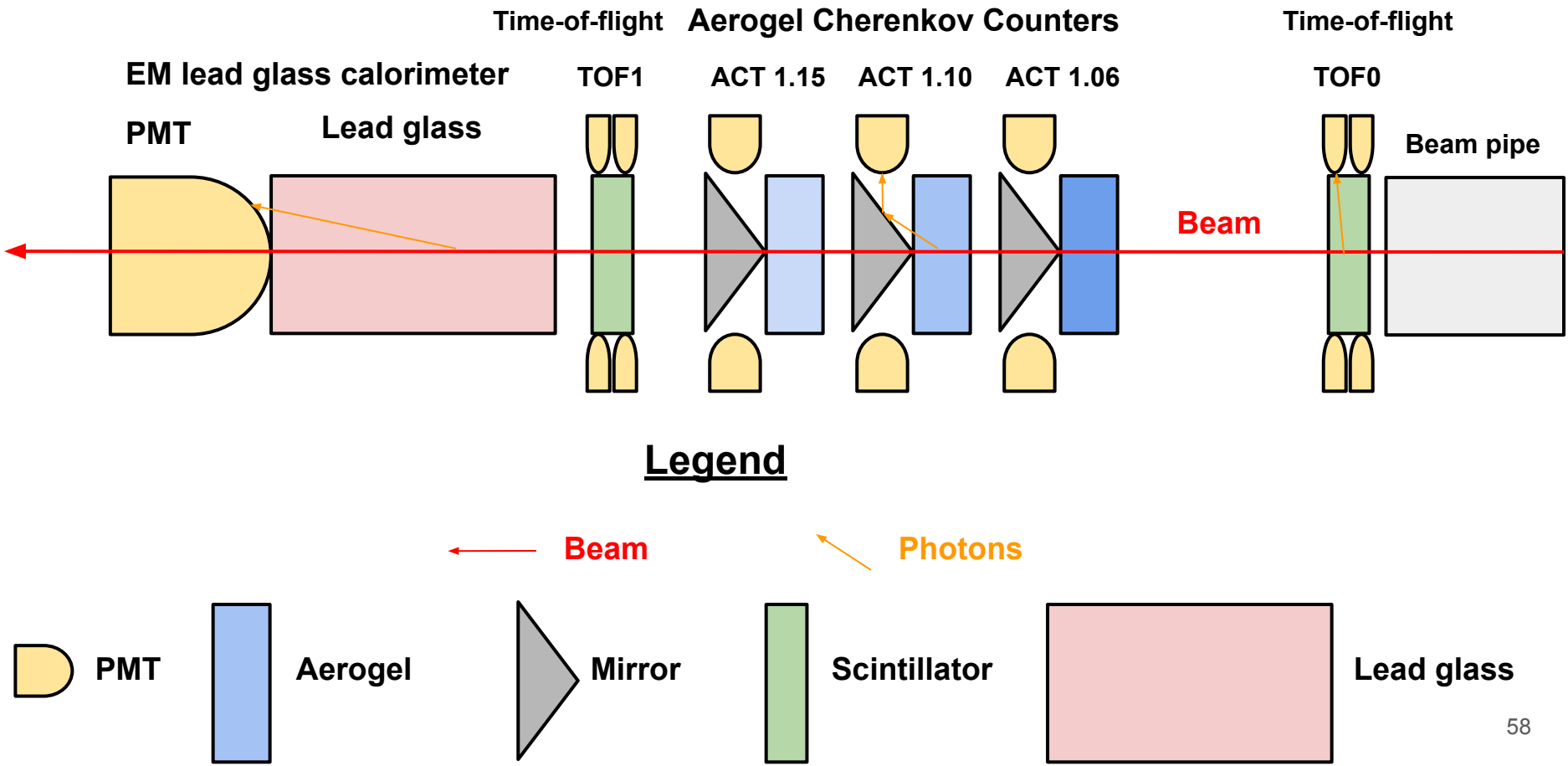

PETZL
DANGER HIGH VOLTAGE
DANGER HIGH VOLTAGE



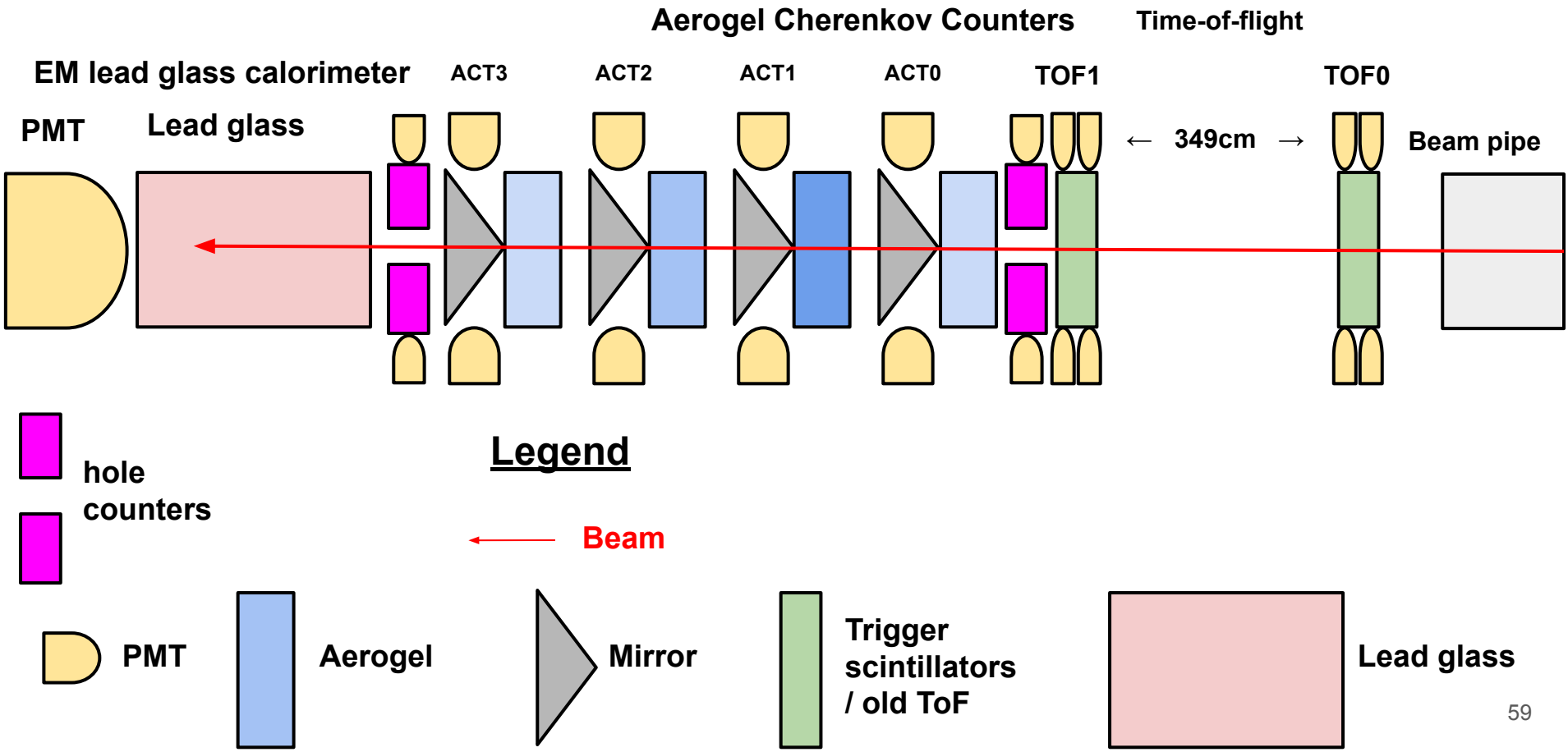

POWER
USER KEY
OFF
ON
MASTER KEY
Master
OFF
EMERGENCY
HORIZONTAL
3683
VERTICAL
01
FAST
FAST
DANGER
HIGH
VOLTAGE
Maximum Voltage 5kV

https://flic.kr/s/aHBqjAM8AQ

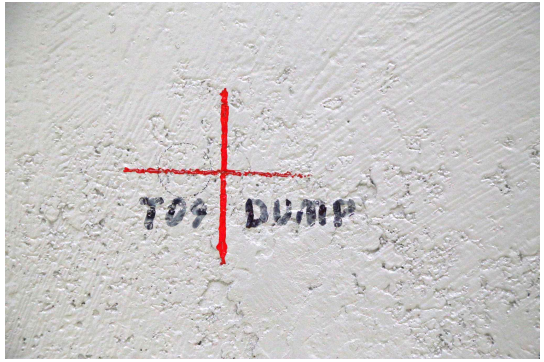
Setup overview 2022

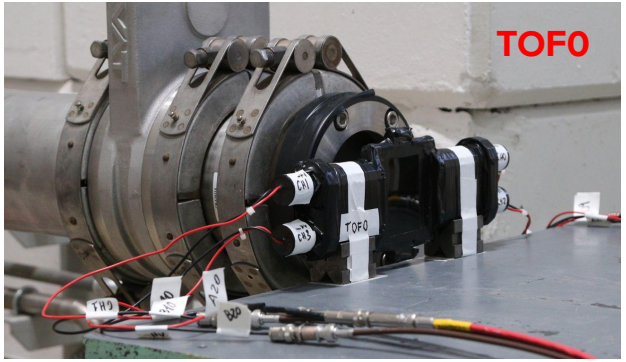


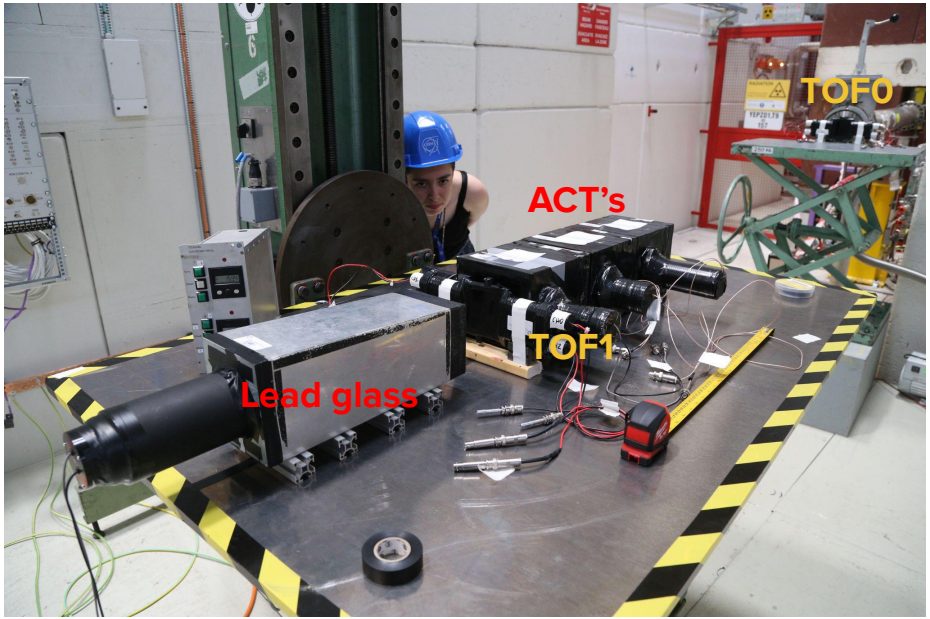
Setup overview updated 2023

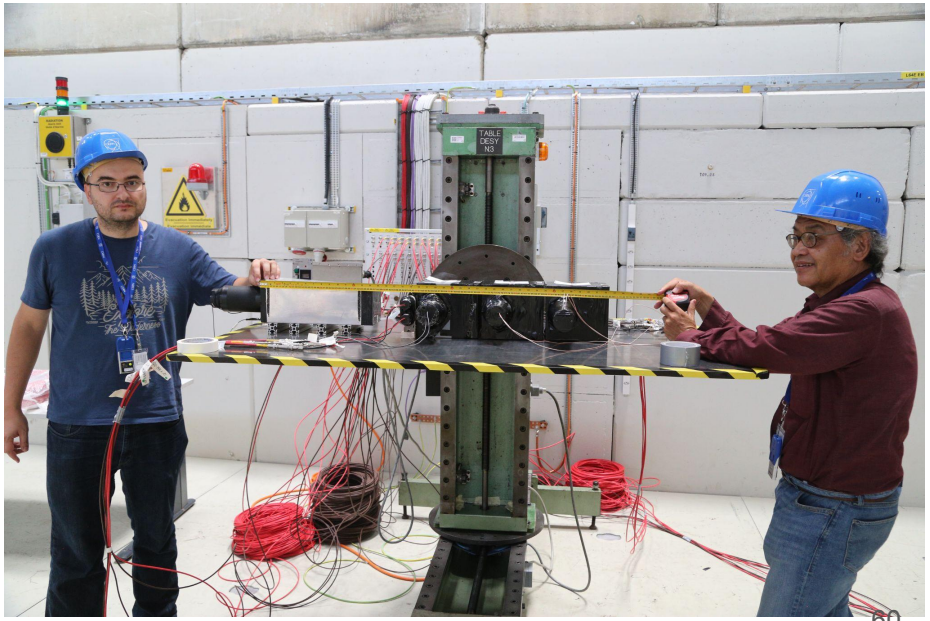


Real Experiment 2022





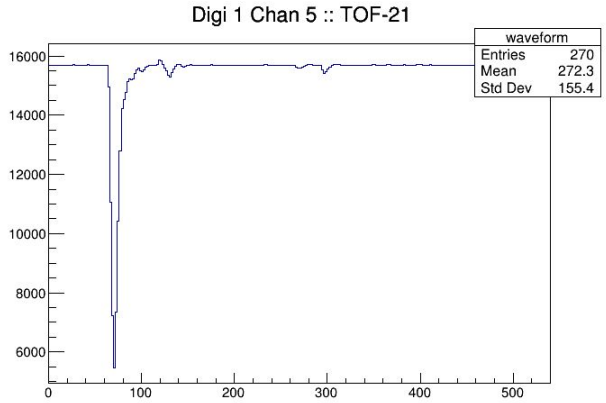


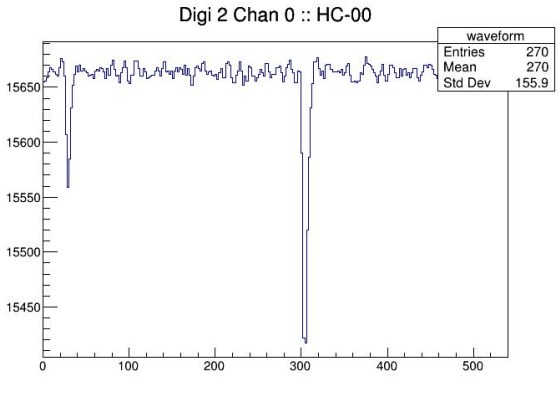


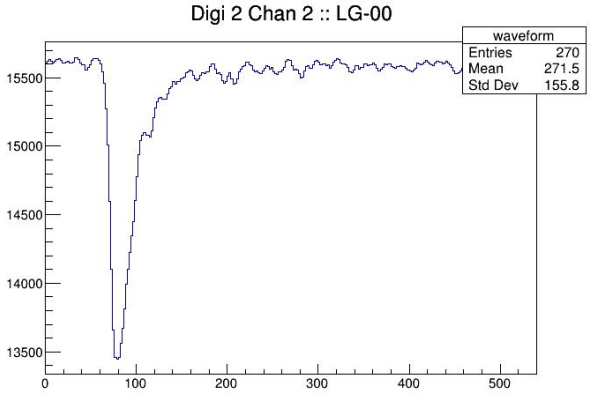
T9 Control Room 2022

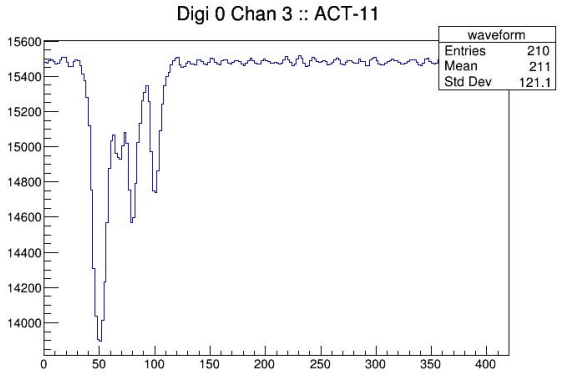
- Incoming spill sound;)
  - https://jointlab.upol.cz/kvita/TB_T9_beep.wav

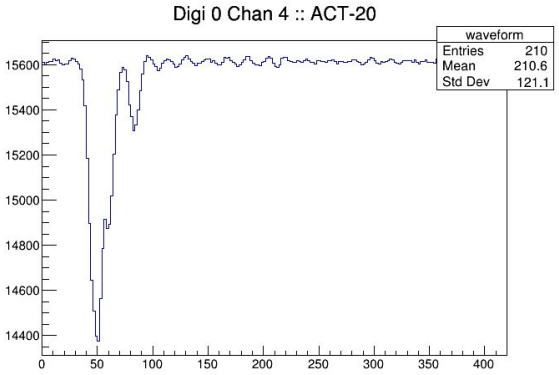
Výstup z fotonásobičů 2023

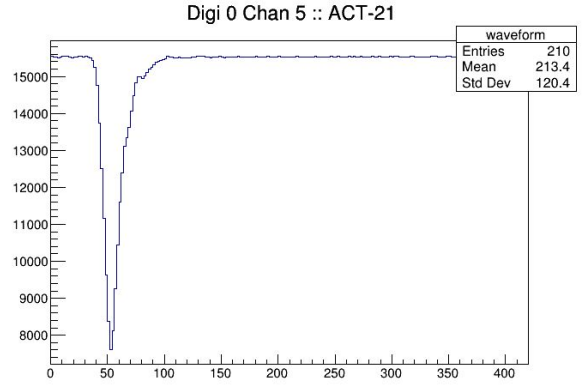


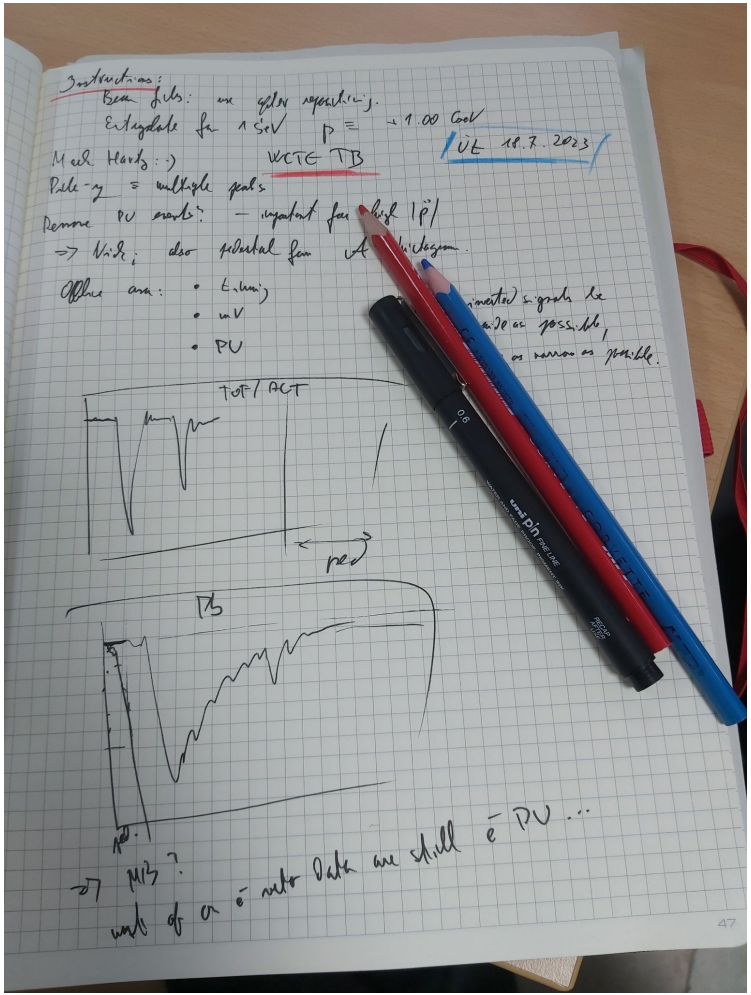


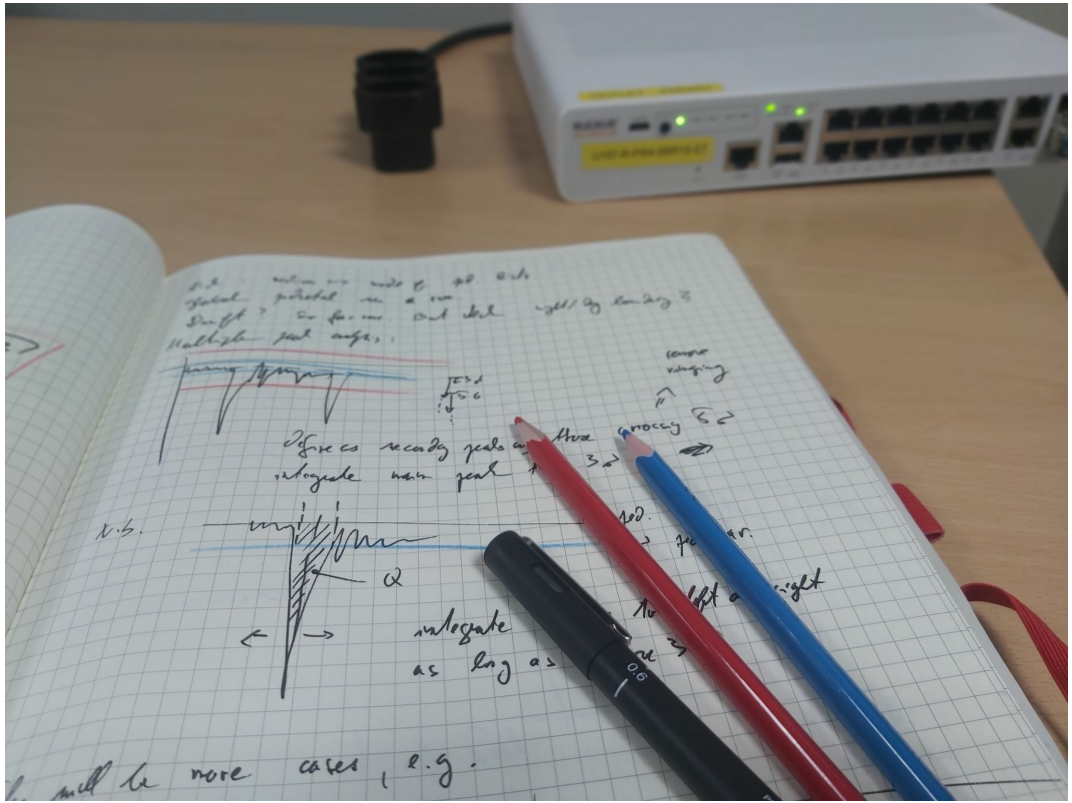




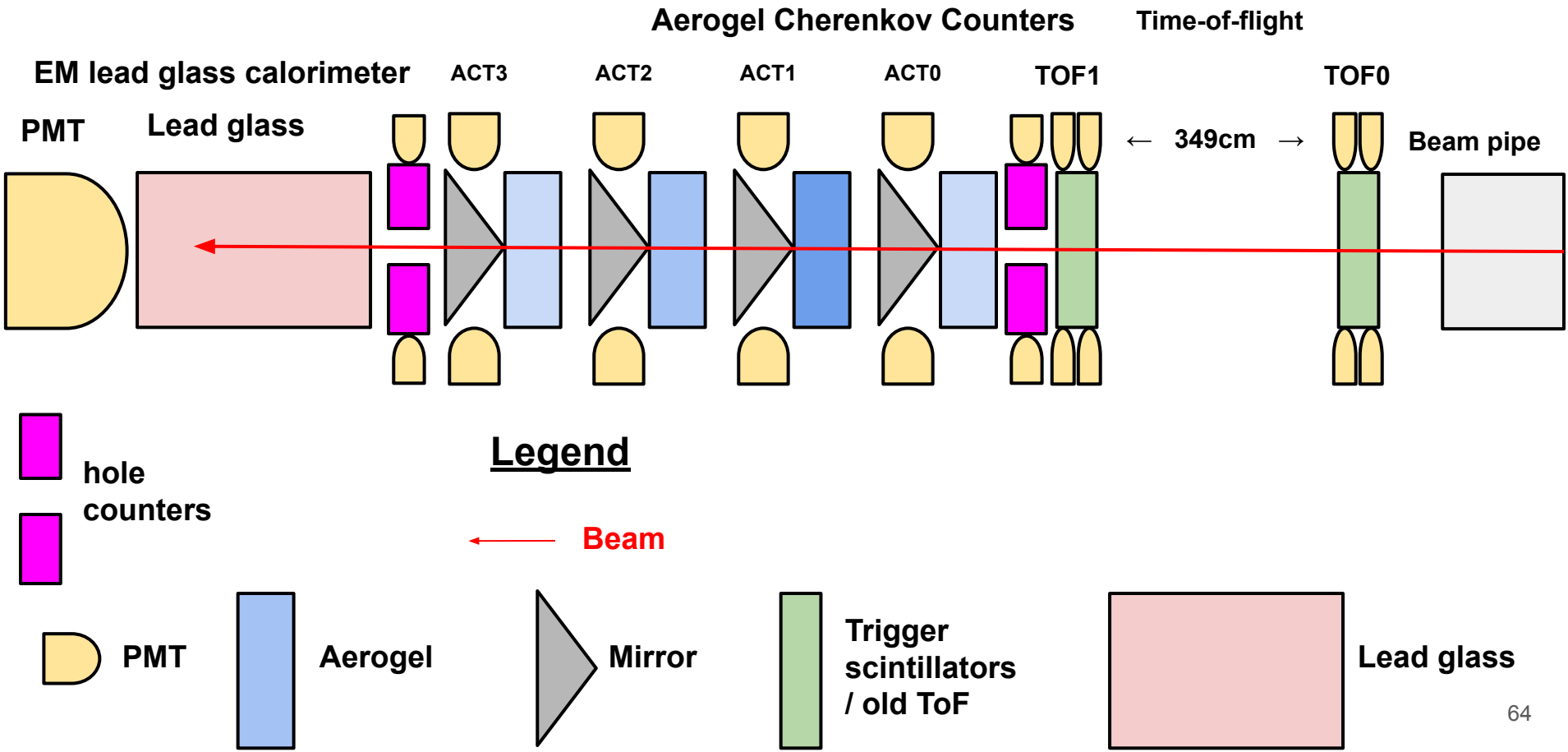




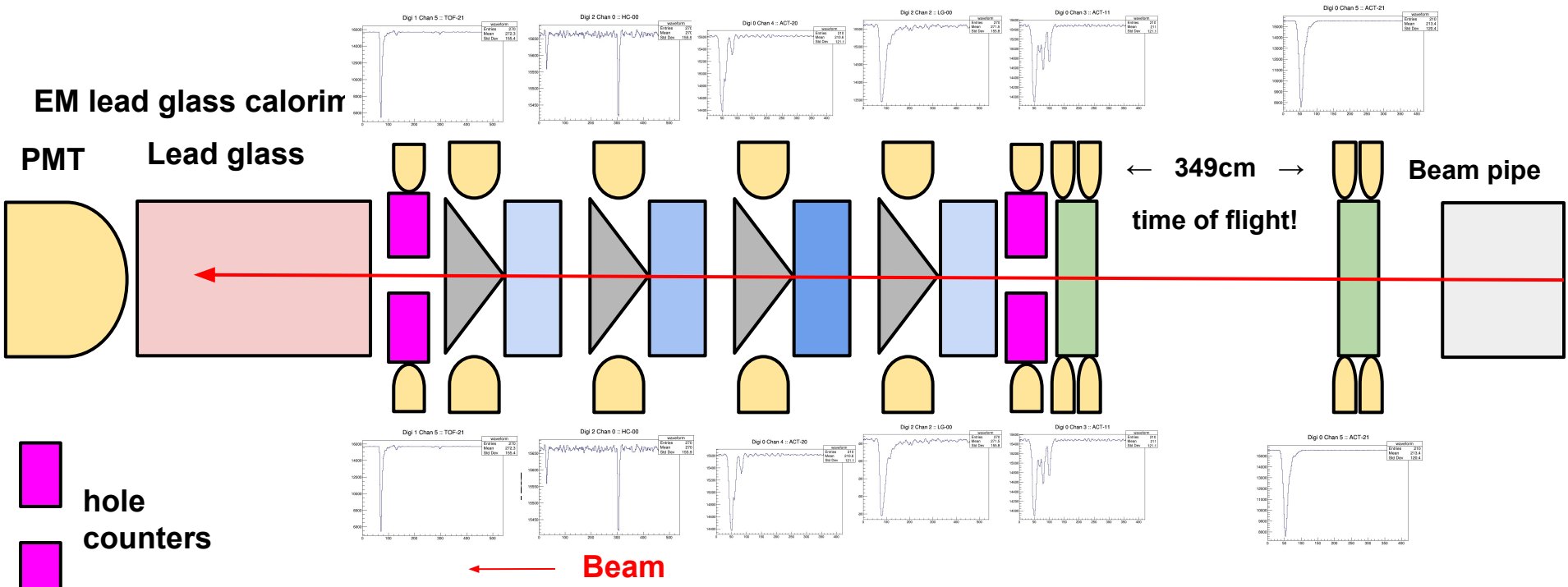

Instructions:
Beam jitter: one pulse separating.
Integrate for 1 sec
ÜL 18.7.2023
WETE TB
Remove PV events:
• Timing
• mV
• PV
TOT / ACT
ped
PS
ped.


integrate
as long as
be more cases, e.g.

Setup overview updated 2023

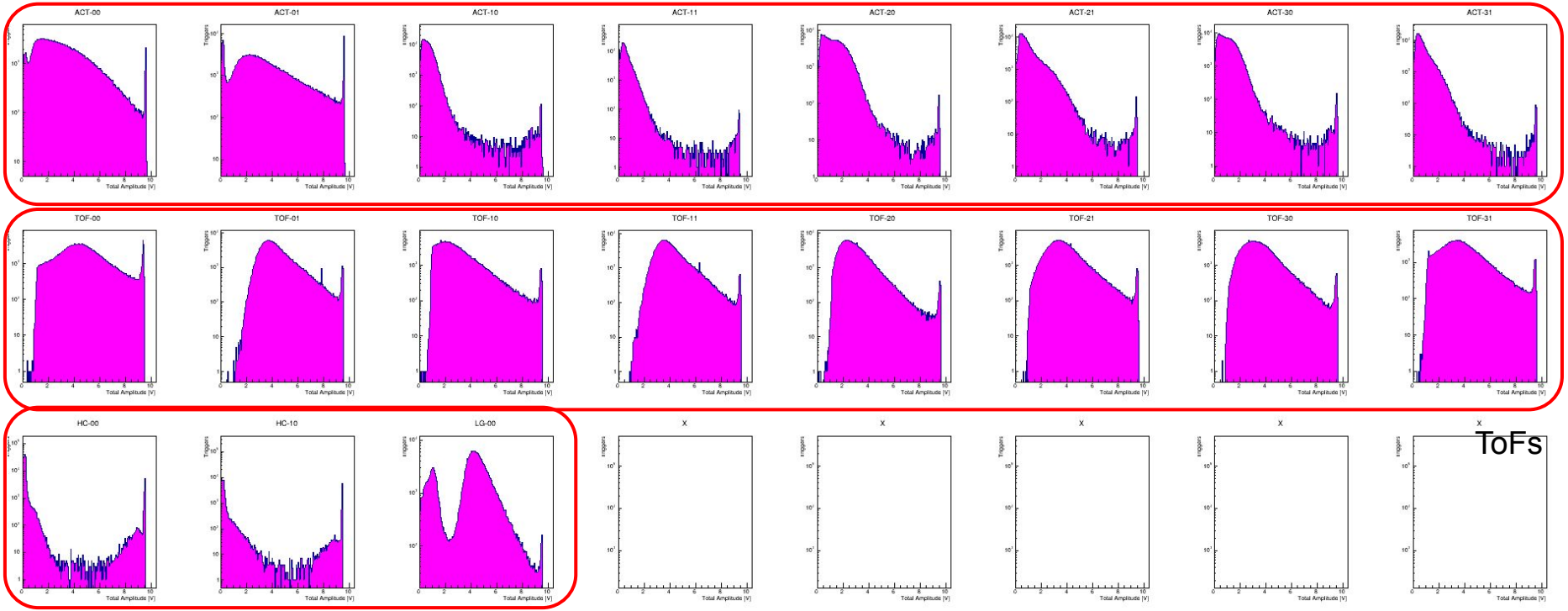


Setup overview updated 2023



Example amplitudes, run 251 :: 2023

Jiri



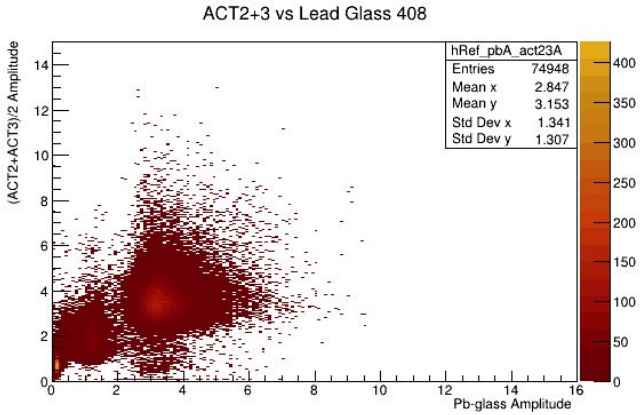
ACTs

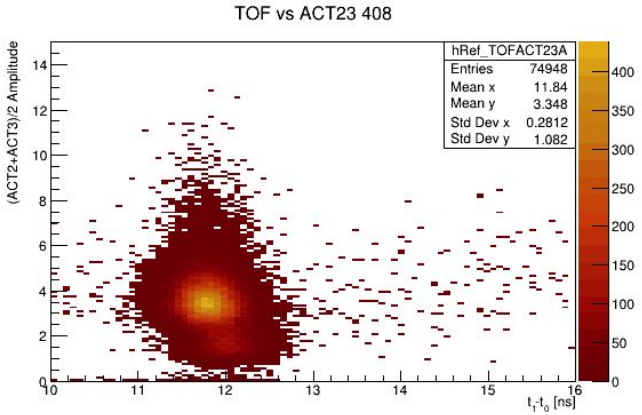
ToFs

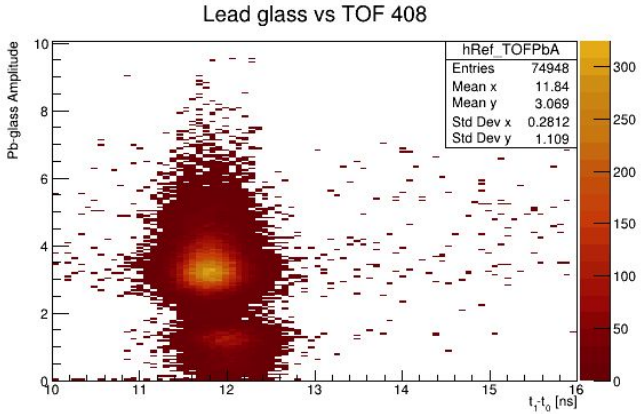
Hole counters and Pb glass

ToF, no cuts on peaks 2023

Jiri

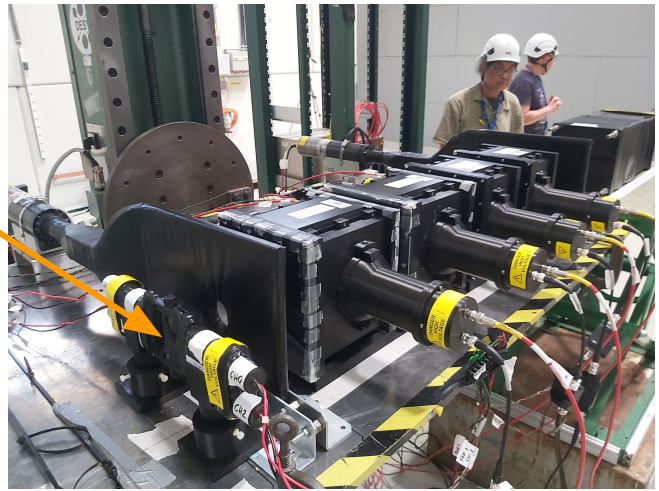






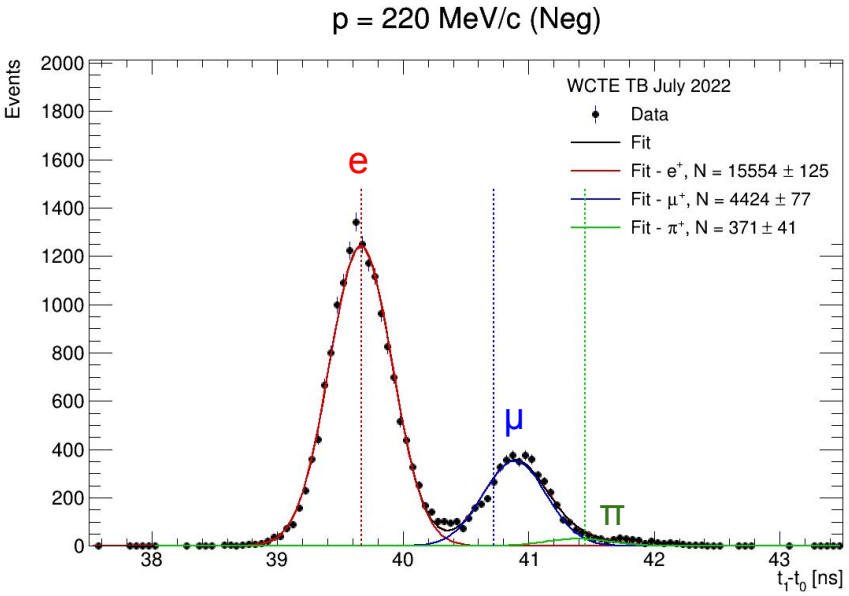
Doba letu = time of flight!

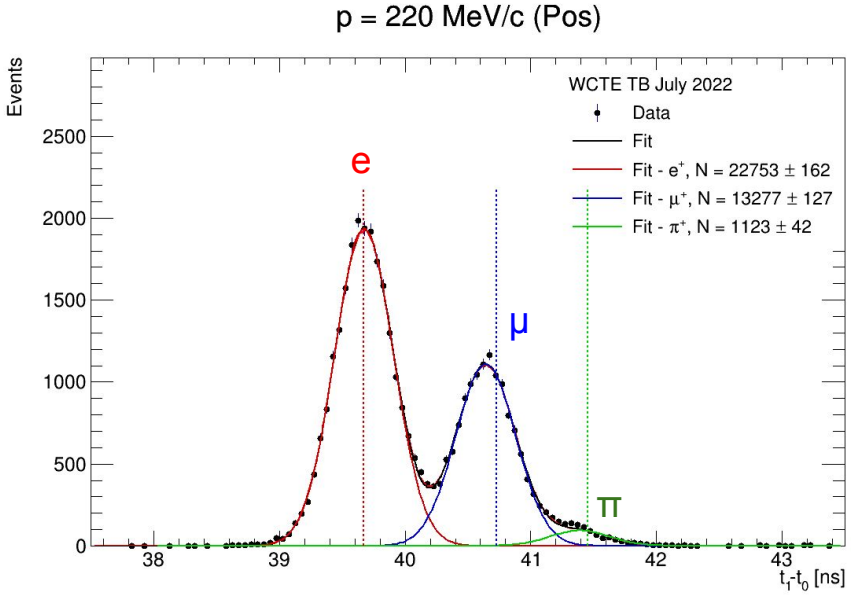
Time-of-flight scintillators, 349cm apart.



Results – Example TOF distributions 2022

- Theoretical vertical lines for the μ and π TOF shifts after the fitted e peak.
- TOF peak positions: a possible beam bias between Pos and Neg momenta?





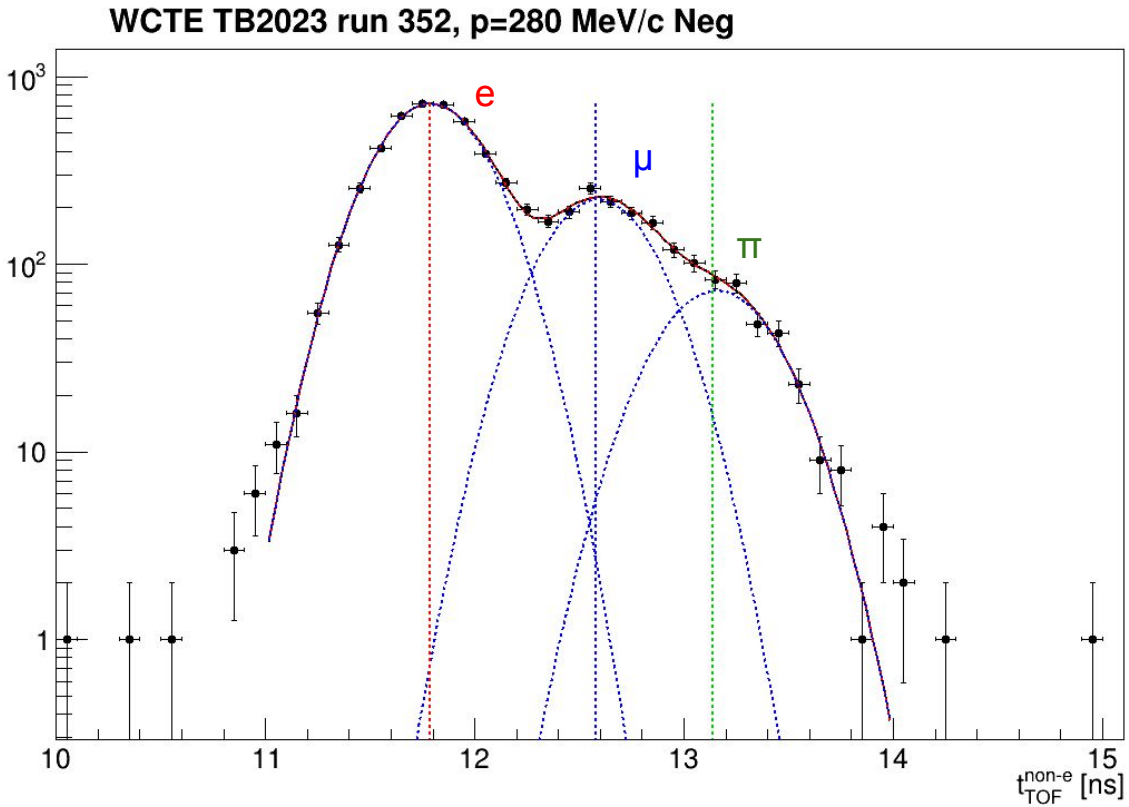
Doba letu

- Computing time… pc = 280 MeV
- Kdy doletí elektron, mion, pion.
- $m_e c^2 = 0.511$ MeV
- $m_\mu c^2 = 105.6$ MeV
- $m_\pi c^2 = 139.6$ MeV

Preliminary tof fit, ~no evidence for bias in neg. *p* 2023

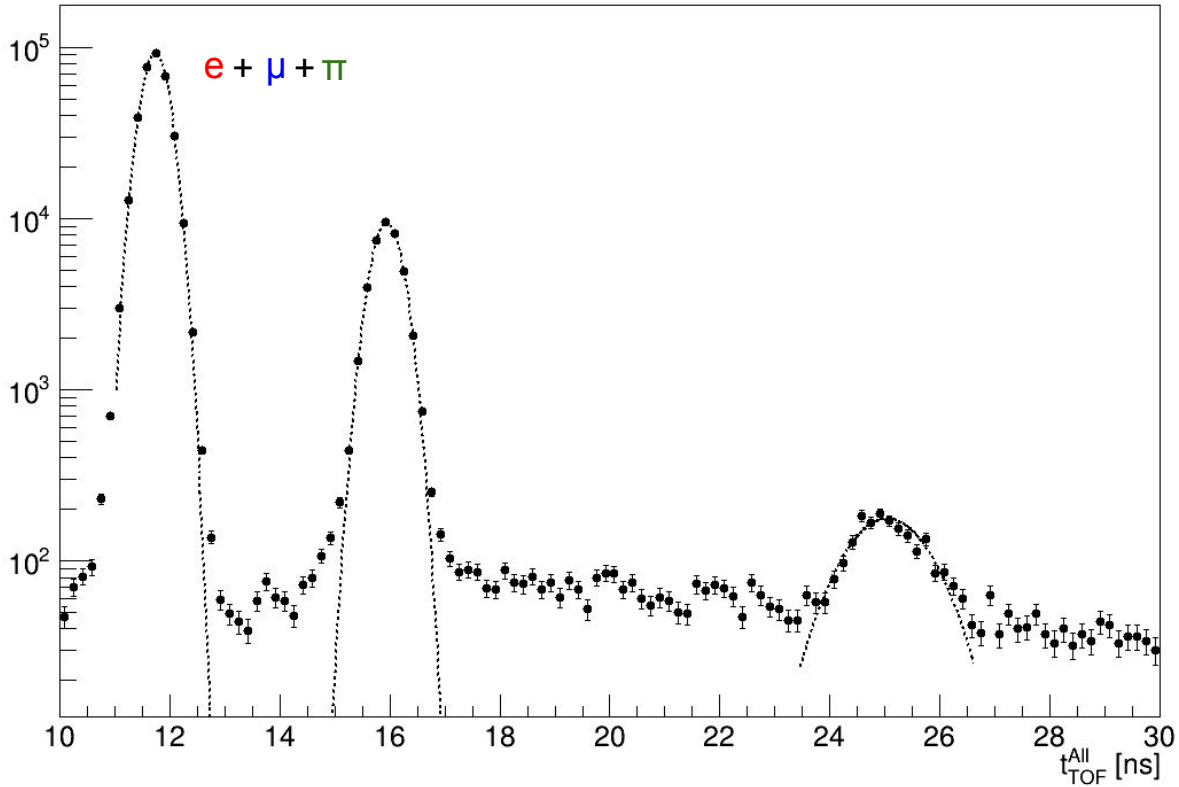
Jiri

- require simple [ACT2+ACT3]/2 < 3. to remove most of electrons



Results – Example TOF distributions 2023, 1000 MeV/c

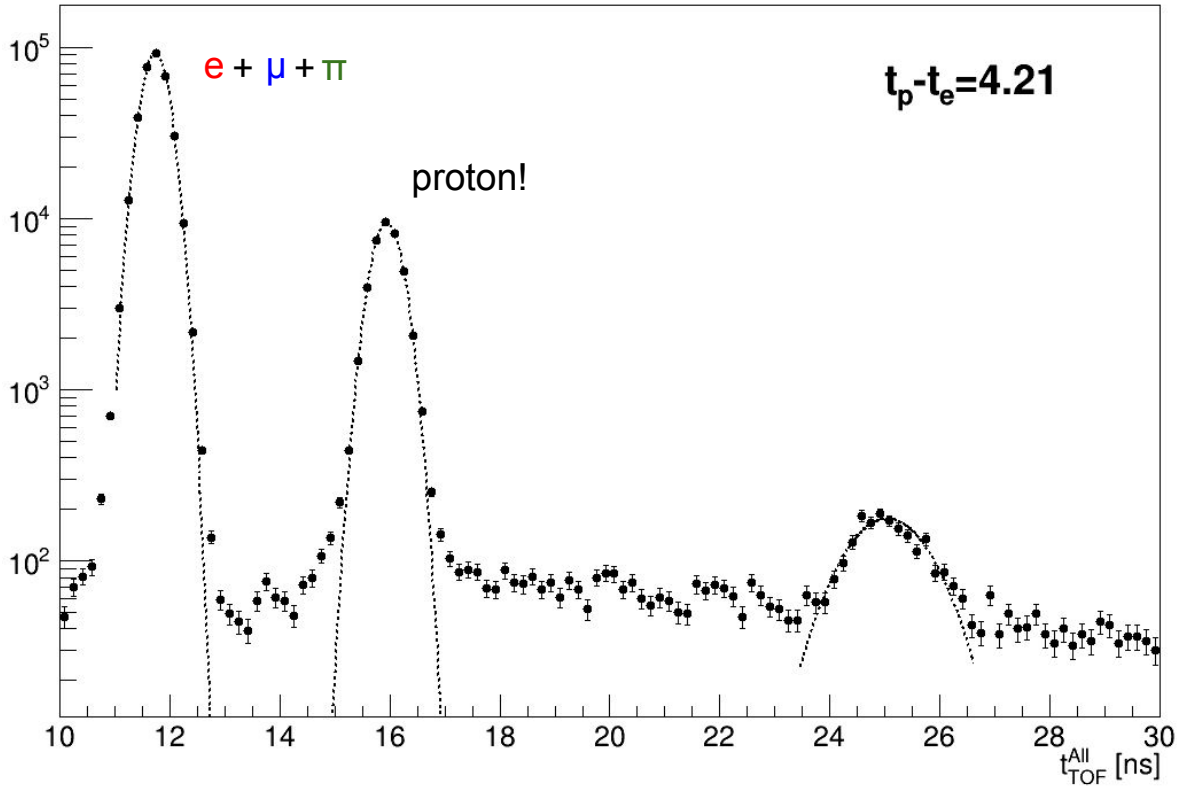
- ?



Doba letu

- Computing time… pc = 1000 MeV
- Kdy doletí elektron, mion, pion…?
- $m_e c^2 = 0.511$ MeV
- $m_\mu c^2 = 105.6$ MeV
- $m_\pi c^2 = 139.6$ MeV
- $m_p c^2 = 938.27$ MeV

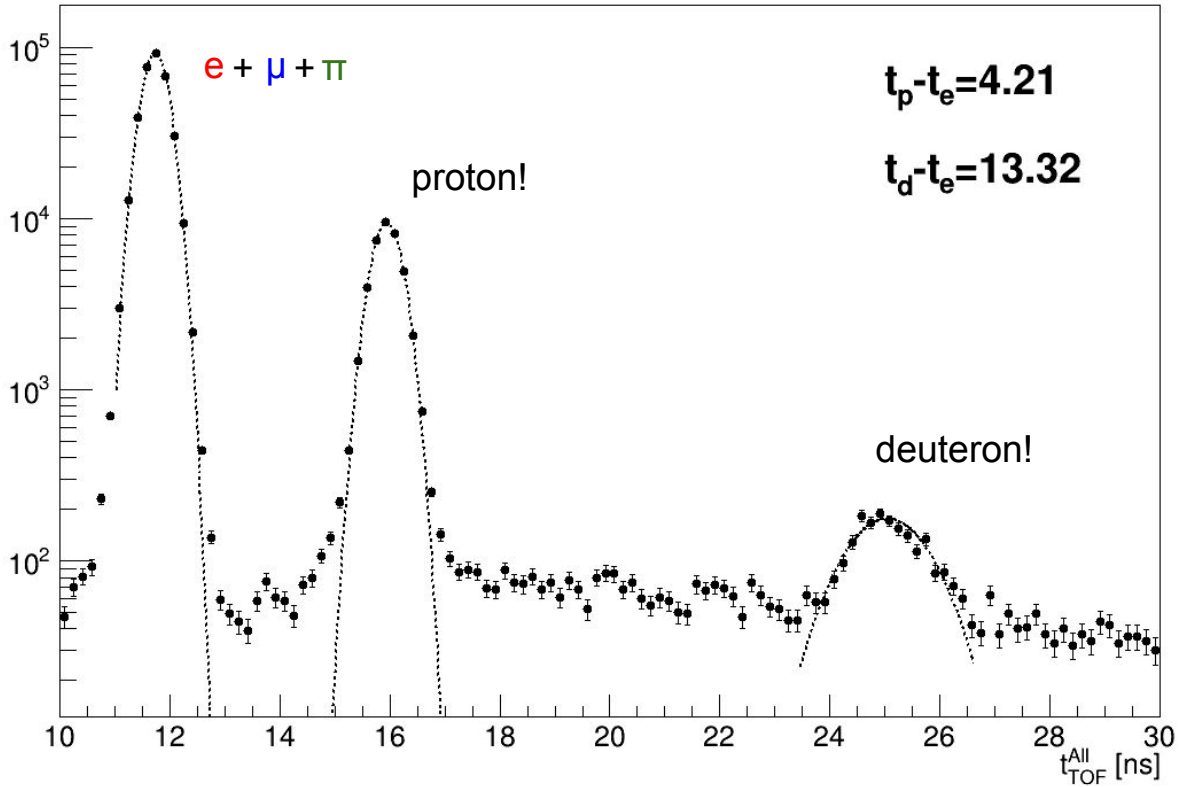
Results – Example TOF distributions 2023, 1000 MeV/c



Doba letu

- Computing time… pc = 1000 MeV
- Kdy doletí elektron, mion, pion…?
- $m_e c^2 = 0.511$ MeV
- $m_\mu c^2 = 105.6$ MeV
- $m_\pi c^2 = 139.6$ MeV
- $m_p c^2 = 938.27$ MeV
- $m_n c^2 = 939.565$ MeV
- $m_D c^2 = 1875.613$ MeV

Results – Example TOF distributions 2023, 1000 MeV/c

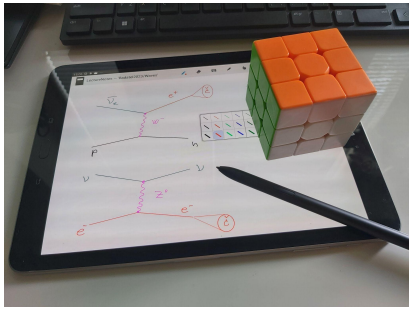


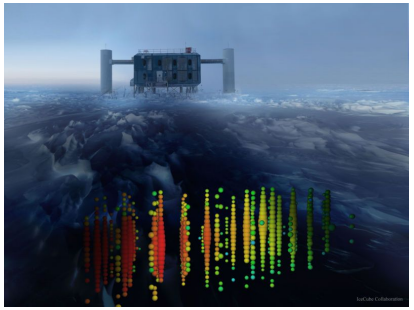
You found your particle!

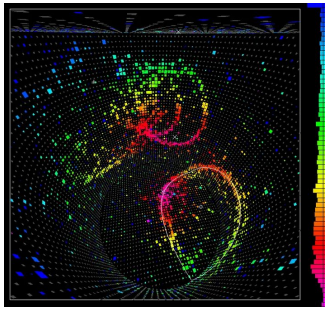
The End

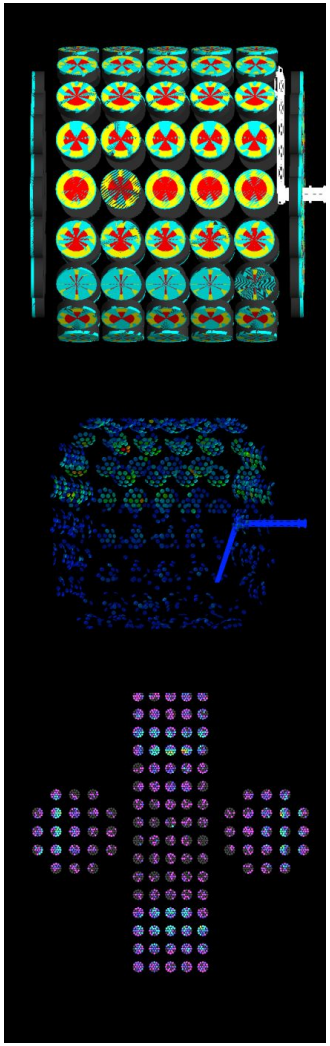
Backup

- Čerenkovovo záření
  - rozpad protonu, detekce neutrin, identifikace částic.
- Neutrinos not dead!
- Nejlehčí hmotné částice v mikrosvětě
- Nejméně interagující.
- Neutrina
  - ze Slunce
  - z atmosféry (z kosmického záření)
  - Astrofyzikální, kosmologická
  - Urychlovačová, reaktorová, geofyzikální…
- Studium oscilací neutrin a (narušení) symetrie mezi částicemi a antičásticemi.
- Neutrinová astronomie
  - nejen čekání na další supernovu:)

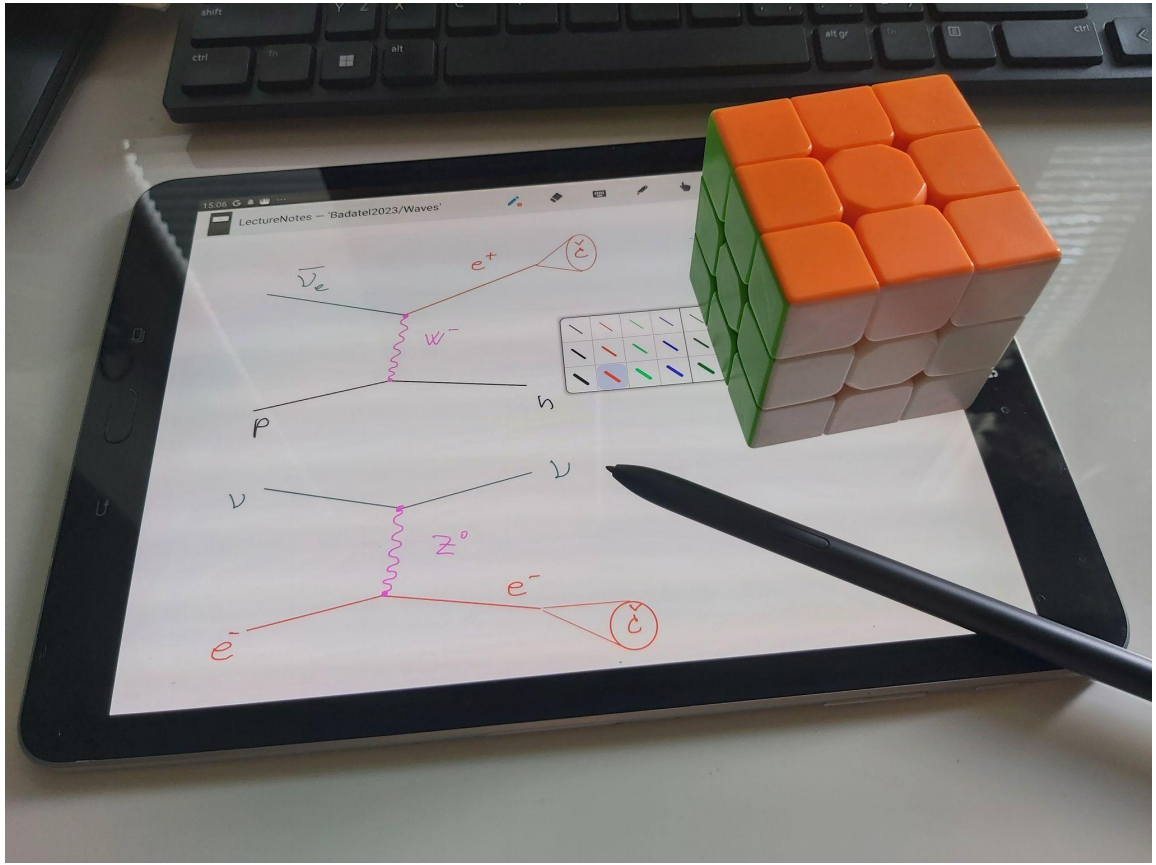




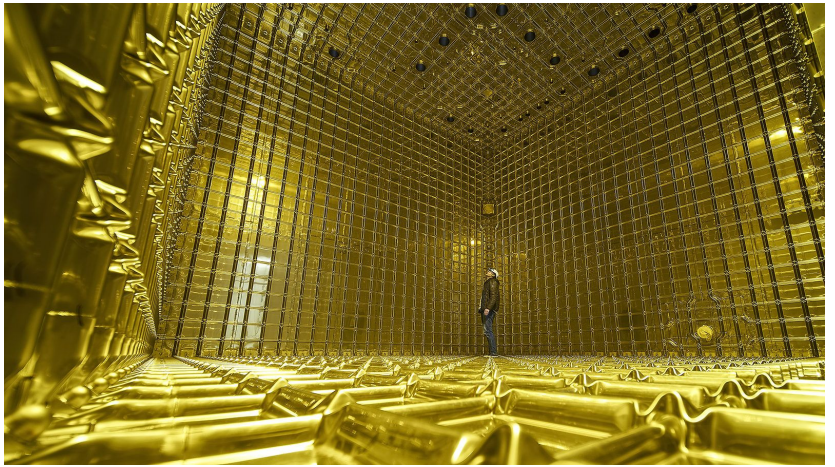


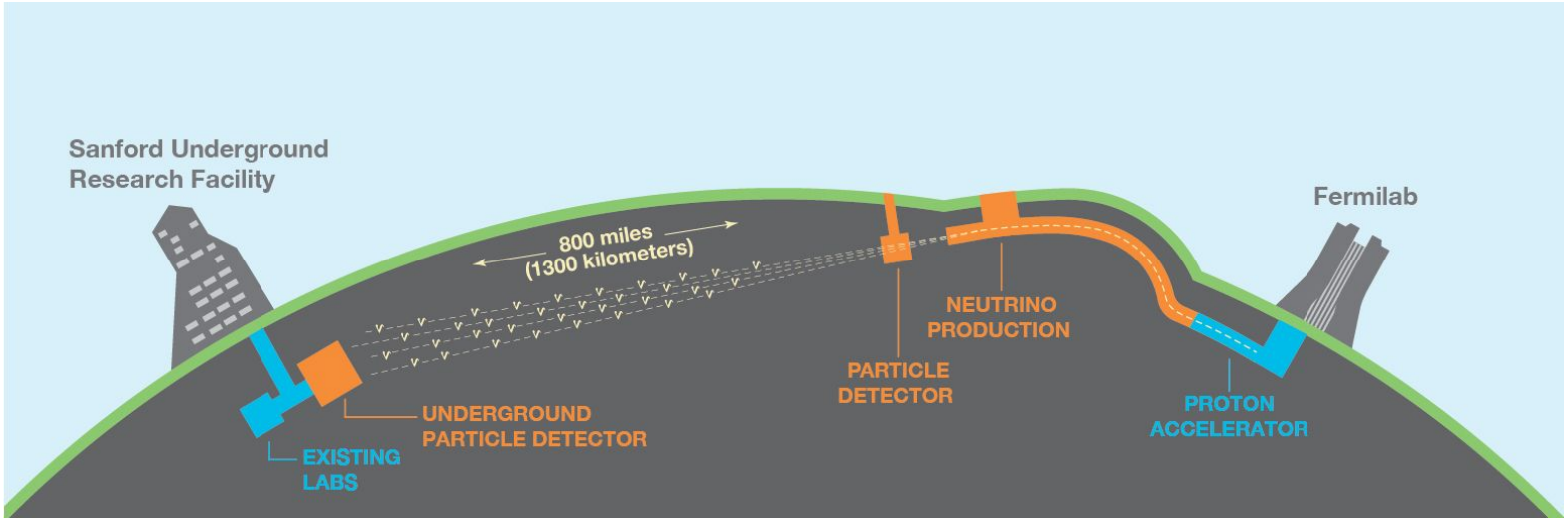


Backup

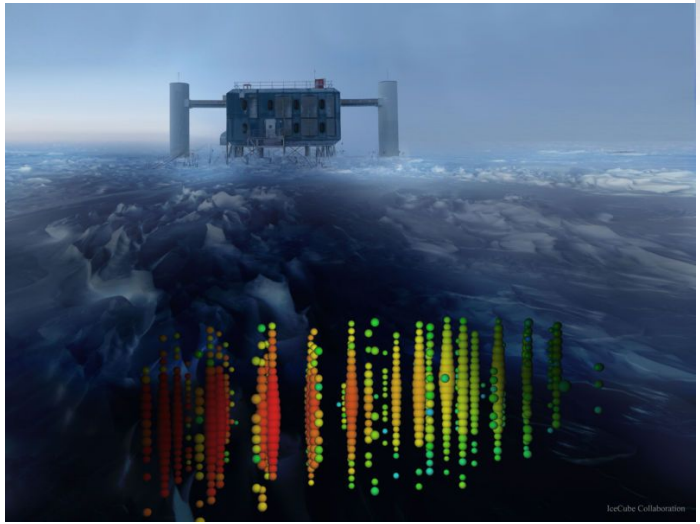

LectureNotes — 'Badatel2023/Waves'
$\bar{\nu}_e$
e^+
$\bar{c}$
W^-
p
h
ν
ν
Z^0
e^-
e^-
$\bar{c}$

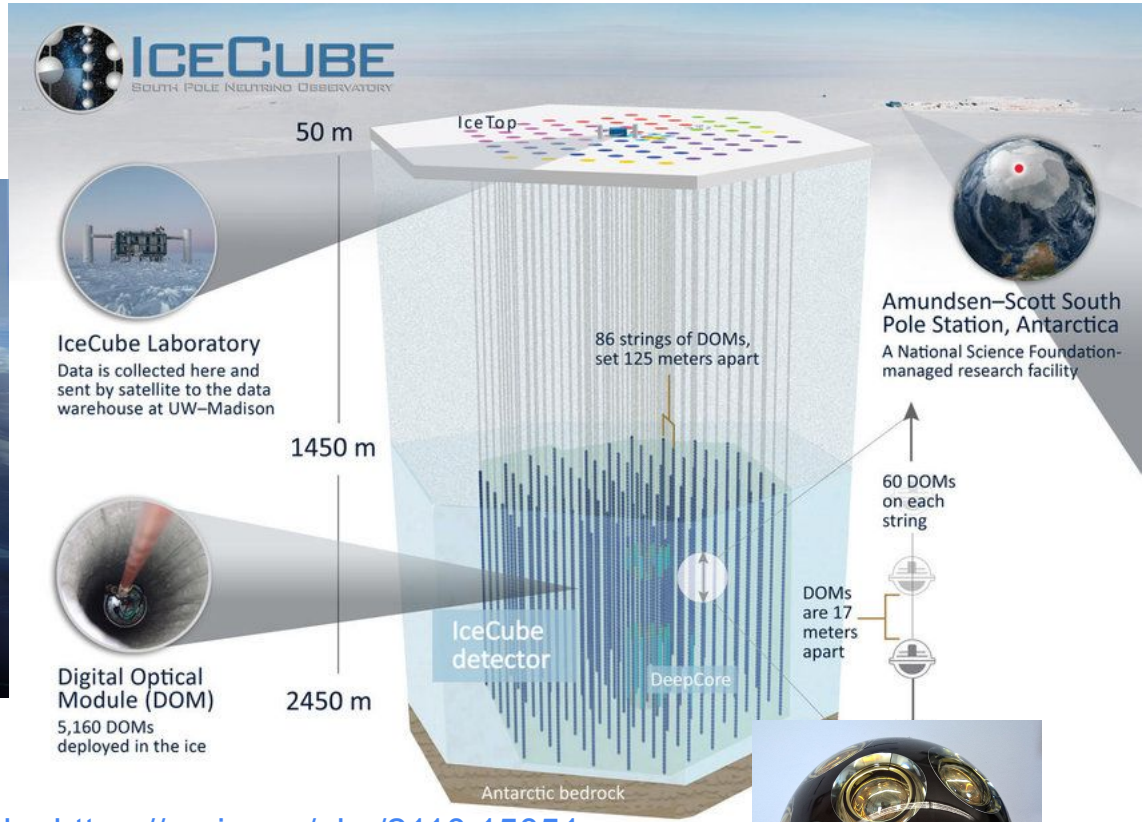
DUNE, ProtoDUNE
a neutrina z urychlovače ve americké laboratoři FERMILAB





IceCUBE





Francis Halzen

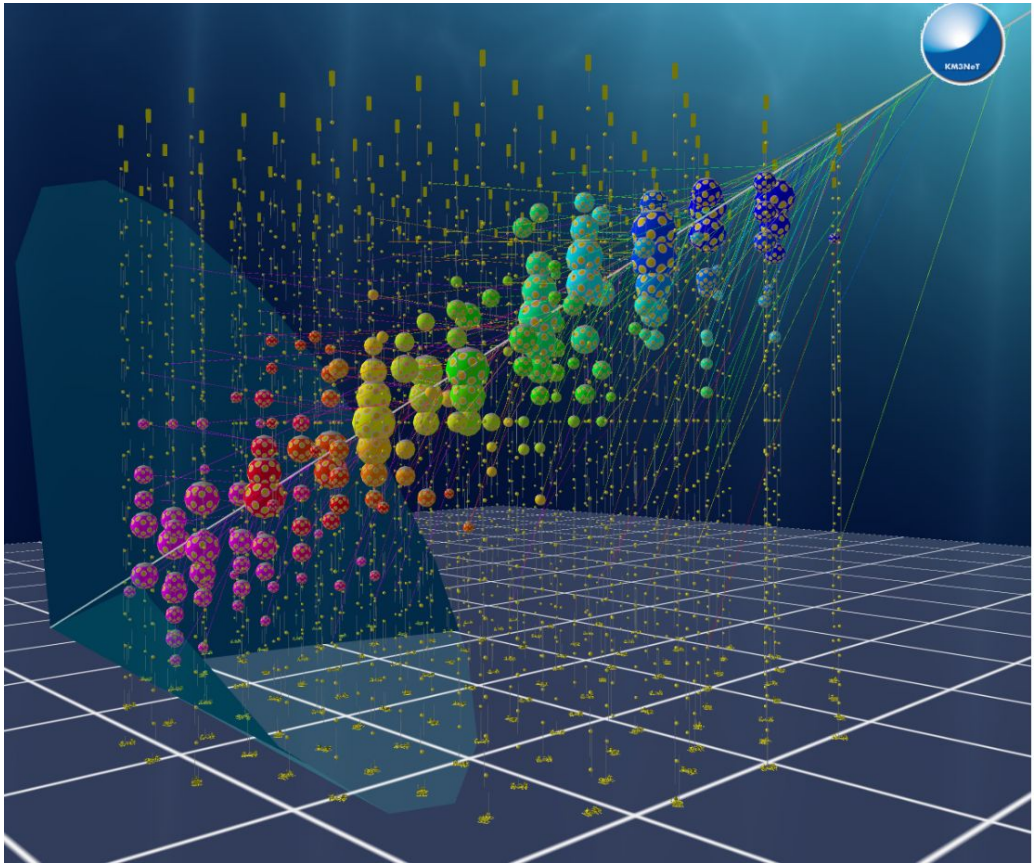
http://jointlab.upol.cz/kvita/idpacs2022.html :: https://arxiv.org/abs/2110.15051

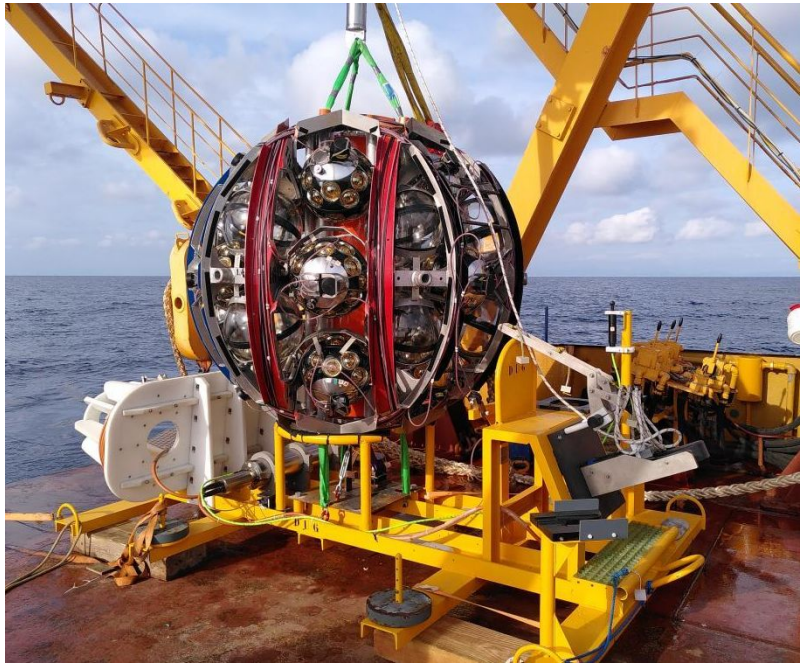
https://www.nature.com/articles/s41586-021-03256-1

https://sciencesprings.wordpress.com/tag/u-wisconsin-icecube-collaboration/

https://www.researchgate.net/figure/Prototypes-of-an-mDOM-left-and-an-D-Egg-right-for-the-IceCube-Upgrade_fig8_351199668

KM3NeT





Neutrina:)

Original - Photocopie of PLC 0393

Abschrift/15.12.30 PM

Offener Brief an die Gruppe der Radioaktiven bei der Gauvereins-Tagung zu Tübingen.

Abschrift

Physikalisches Institut
der Eidg. Technischen Hochschule
Zürich

Zürich, 4. Dez. 1930
Gloriastrasse

Liebe Radioaktive Damen und Herren,

Wie der Ueberbringer dieser Zeilen, den ich huldvollst ansuhören bitte, Ihnen des näheren auseinandersetzen wird, bin ich angesichts der "falschen" Statistik der N- und Li-6 Kerne, sowie des kontinuierlichen beta-Spektrums auf einen verzweifelten Ausweg verfallen um den "Wechselsatz" (1) der Statistik und den Energiesatz zu retten. Nämlich die Möglichkeit, es könnten elektrisch neutrale Teilchen, die ich Neutronen nennen will, in den Kernen existieren, welche den Spin 1/2 haben und das Ausschliessungsprinsip befolgen und sich von Lichtquanten ausserdem noch dadurch unterscheiden, dass sie nicht mit Lichtgeschwindigkeit laufen. Die Masse der Neutronen müsste von derselben Grössenordnung wie die Elektronenmasse sein und jedenfalls nicht grösser als 0,01 Protonenmasse.- Das kontinuierliche beta- Spektrum wäre dann verständlich unter der Annahme, dass beim beta-Zerfall mit dem Elektron jeweils noch ein Neutron emittiert wird, derart, dass die Summe der Energien von Neutron und Elektron konstant ist.

Nun handelt es sich weiter darum, welche Kräfte auf die Neutronen wirken. Das wahrscheinlichste Modell für das Neutron scheint mir aus wellenmechanischen Gründen (näheres weiss der Ueberbringer dieser Zeilen) dieses zu sein, dass das ruhende Neutron ein magnetischer Dipol von einem gewissen Moment μ ist. Die Experimente verlangen wohl, dass die ionisierende Wirkung eines solchen Neutrons nicht grösser sein kann, als die eines gamma-Strahls und darf dann μ wohl nicht grösser sein als $e \cdot (10^{-13}$ cm).

Ich traue mich vorläufig aber nicht, etwas über diese Idee zu publizieren und wende mich erst vertrauensvoll an Euch, liebe Radioaktive, mit der Frage, wie es um den experimentellen Nachweis eines solchen Neutrons stände, wenn dieses ein ebensolches oder etwa 10mal grösseres Durchdringungsvermögen besitzen würde, wie ein gamma-Strahl.

Ich gebe zu, dass mein Ausweg vielleicht von vornherein wenig wahrscheinlich erscheinen wird, weil man die Neutronen, wenn sie existieren, wohl schon längst gesehen hätte. Aber nur wer wagt, gewinnt und der Ernst der Situation beim kontinuierliche beta-Spektrum wird durch einen Ausspruch meines verehrten Vorgängers im Amte, Herrn Debye, beleuchtet, der mir kürzlich in Brüssel gesagt hat: "O, daran soll man am besten gar nicht denken, sowie an die neuen Steuern." Darum soll man jeden Weg zur Rettung ernstlich diskutieren.- Also, liebe Radioaktive, prüfet, und richtet.- Leider kann ich nicht persönlich in Tübingen erscheinen, da ich infolge eines in der Nacht vom 6. zum 7 Dez. in Zürich stattfindenden Balles hier unabkömmlich bin.- Mit vielen Grüssen an Euch, sowie an Herrn Back, Euer untertänigster Diener

gez. W. Pauli

Wolfgang Pauli

Abschrift/15.12.30

Offener Brief an die Gruppe der Radioaktiven bei der Gauvereins-Tagung zu Tübingen.

Abschrift

Physikalisches Institut
der Eidg. Technischen Hochschule
Zürich

Zürich, 4. Dez. 1930
Gloriastrasse

Liebe Radioaktive Damen und Herren,

Wie der Ueberbringer dieser Zeilen, den ich huldvollst anzuhören bitte, Ihnen des näheren auseinandersetzen wird, bin ich angesichts der "falschen" Statistik der N- und Li-6 Kerne, sowie des kontinuierlichen beta-Spektrums auf einen verzweifelten Ausweg verfallen um den "Wechselsatz" (1) der Statistik und den Energiesatz zu retten. Nämlich die Möglichkeit, es könnten elektrisch neutrale Teilchen, die ich Neutronen nennen will, in den Kernen existieren, welche den Spin 1/2 haben und das Ausschliessungsprinzip befolgen und sich von Lichtquanten ausserdem noch dadurch unterscheiden, dass sie nicht mit Lichtgeschwindigkeit laufen. Die Masse der Neutronen müsste von derselben Grössenordnung wie die Elektronenmasse sein und jedenfalls nicht grösser als 0,01 Protonenmasse.- Das kontinuierliche beta- Spektrum wäre dann verständlich unter der Annahme, dass beim beta-Zerfall mit dem Elektron jeweils noch ein Neutron emittiert wird, derart, dass die Summe der Energien von Neutron und Elektron konstant ist.

Nun handelt es sich weiter darum, welche Kräfte auf die Neutronen wirken. Das wahrscheinlichste Modell für das Neutron scheint mir aus wellenmechanischen Gründen (näheres weiss der Ueberbringer dieser Zeilen) dieses zu sein, dass das ruhende Neutron ein magnetischer Dipol von einem gewissen Moment μ ist. Die Experimente verlangen wohl, dass die ionisierende Wirkung eines solchen Neutrons nicht grösser sein kann, als die eines gamma-Strahls und darf dann μ wohl nicht grösser sein als $e \cdot (10^{-13}$ cm).

Ich traue mich vorläufig aber nicht, etwas über diese Idee zu publizieren und wende mich erst vertrauensvoll an Euch, liebe Radioaktive, mit der Frage, wie es um den experimentellen Nachweis eines solchen Neutrons stände, wenn dieses ein ebensolches oder etwa 10mal grösseres Durchdringungsvermögen besitzen würde, wie ein gamma-Strahl.

Ich gebe zu, dass mein Ausweg vielleicht von vornherein wenig wahrscheinlich erscheinen wird, weil man die Neutronen, wenn sie existieren, wohl schon längst gesehen hätte. Aber nur wer wagt, gewinnt und der Ernst der Situation beim kontinuierliche beta-Spektrum wird durch einen Ausspruch meines verehrten Vorgängers im Amte, Herrn Debye, beleuchtet, der mir kürzlich in Brüssel gesagt hat: "O, daran soll man am besten gar nicht denken, sowie an die neuen Steuern." Darum soll man jeden Weg zur Rettung ernstlich diskutieren.- Also, liebe Radioaktive, prüfet, und richtet.- Leider kann ich nicht persönlich in Tübingen erscheinen, da ich infolge eines in der Nacht vom 6. zum 7 Dez. in Zürich stattfindenden Balles hier unabkömmlich bin.- Mit vielen Grüssen an Euch, sowie an Herrn Back, Euer untertänigster Diener

gez. W. Pauli

http://library.cern/archives/history_CERN/historical_images/month-88-years-ago

Neutrina:)

COSMIC GALL

Every second, hundreds of billions of these neutrinos pass through each square inch of our bodies, coming from above during the day and from below at night, when the sun is shining on the other side of the earth!

—From "An Explanatory Statement on Elementary Particle Physics," by M. A. Ruderman and A. H. Rosenfeld, in American Scientist.

Neutrinos, they are very small.
  They have no charge and have no mass
And do not interact at all.
The earth is just a silly ball
  To them, through which they simply pass,
Like dustmaids down a drafty hall
  Or photons through a sheet of glass.
  They snub the most exquisite gas,
Ignore the most substantial wall,
  Cold shoulder steel and sounding brass,
Insult the stallion in his stall,
  And, scorning barriers of class,
Infiltrate you and me! Like tall
And painless guillotines, they fall
  Down through our heads into the grass.
At night, they enter at Nepal
  And pierce the lover and his lass
From underneath the bed - you call
  It wonderful; I call it crass.

— *John Updike*

Wolfgang Pauli

https://twitter.com/mcnees/status/1334877571079671810

Neutrina:)

Offener Brief an die Gruppe der Radioaktiven bei der Gauvereins-Tagung zu Tübingen.

Abschrift

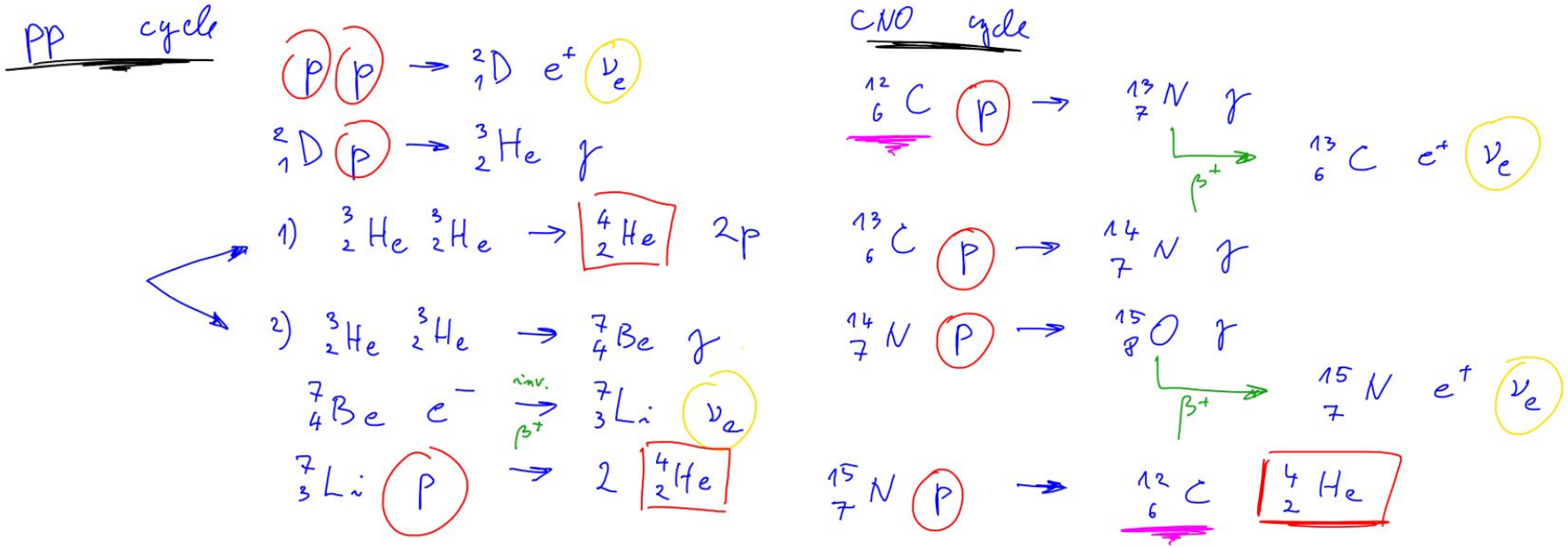
Physikalisches Institut
der Eidg. Technischen Hochschule
Zürich

Zürich, 4. Dez. 1930
Gloriastrasse

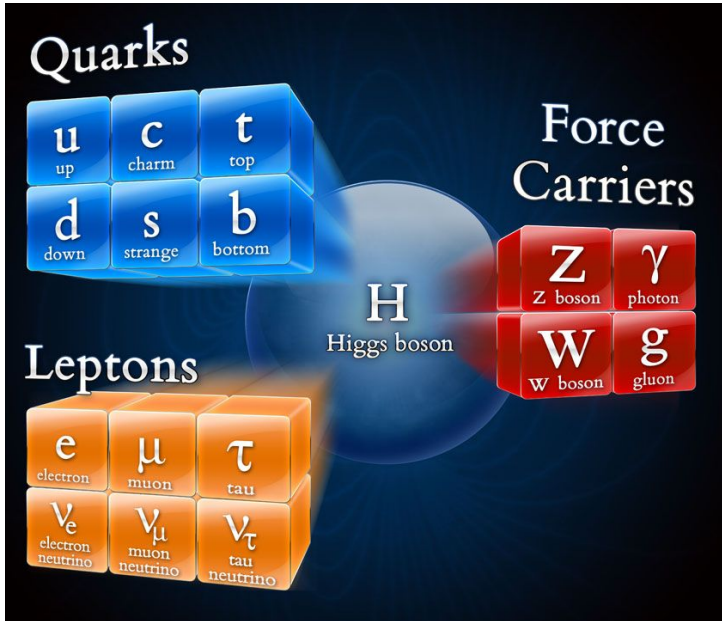
Liebe Radioaktive Damen und Herren,

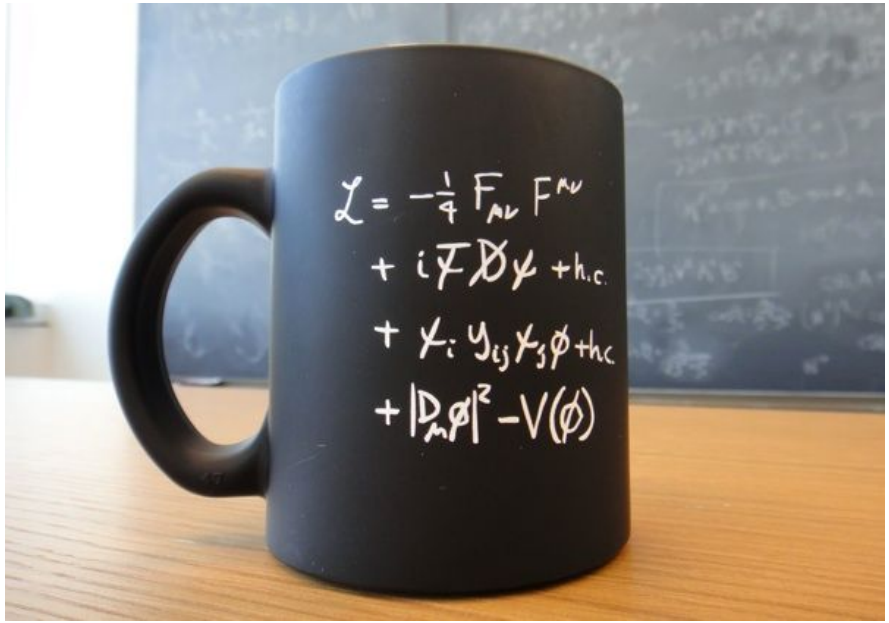
Wie der Ueberbringer dieser Zeilen, den ich huldvollst anzuhören bitte, Ihnen des näheren auseinandersetzen wird, bin ich angesichts der "falschen" Statistik der N- und Li-6 Kerne, sowie

Neutrina:)

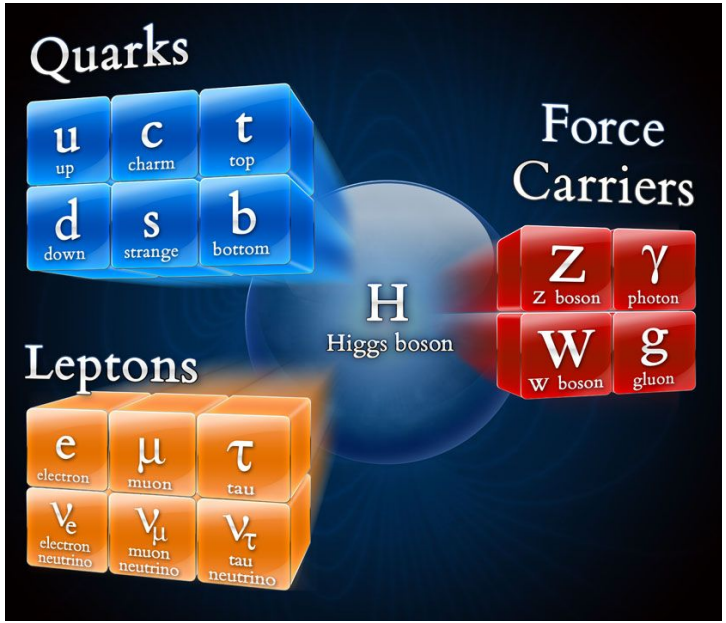

pp cycle
p p → ²₁D e⁺ νe
²₁D p → ³₂He γ
1) ³₂He ³₂He → ⁴₂He 2p
2) ³₂He ³₂He → ⁷₄Be γ
⁷₄Be e⁻ → ⁷₃Li νe
inv. β⁺
⁷₃Li p → 2 ⁴₂He
CNO cycle
¹²₆C p → ¹³₇N γ
β⁺
¹³₆C e⁺ νe
¹³₆C p → ¹⁴₇N γ
¹⁴₇N p → ¹⁵₈O γ
β⁺
¹⁵₇N e⁺ νe
¹⁵₇N p → ¹²₆C ⁴₂He

Neutrina:)


Quarks
u up
c charm
t top
d down
s strange
b bottom
Leptons
e electron
μ muon
τ tau
νe electron neutrino
νμ muon neutrino
ντ tau neutrino
H Higgs boson
Force Carriers
Z Z boson
γ photon
W W boson
g gluon


$\mathcal{L} = -\frac{1}{4} F_{\mu\nu} F^{\mu\nu}$
$+ i\bar{\psi} \not{D} \psi + h.c.$
$+ \psi_i y_{ij} \psi_j \phi + h.c.$
$+ |D_\mu \phi|^2 - V(\phi)$

Neutrina:)



$$\mathcal{L}_{SM} = -\tfrac{1}{2}\partial_\nu g^a_\mu \partial_\nu g^a_\mu - g_s f^{abc}\partial_\mu g^a_\nu g^b_\mu g^c_\nu - \tfrac{1}{4}g_s^2 f^{abc}f^{ade}g^b_\mu g^c_\nu g^d_\mu g^e_\nu - \partial_\nu W^+_\mu \partial_\nu W^-_\mu - M^2 W^+_\mu W^-_\mu - \tfrac{1}{2}\partial_\nu Z^0_\mu \partial_\nu Z^0_\mu - \tfrac{1}{2c_w^2}M^2 Z^0_\mu Z^0_\mu - \tfrac{1}{2}\partial_\mu A_\nu \partial_\mu A_\nu - igc_w(\partial_\nu Z^0_\mu(W^+_\mu W^-_\nu - W^+_\nu W^-_\mu) - Z^0_\nu(W^+_\mu \partial_\nu W^-_\mu - W^-_\mu \partial_\nu W^+_\mu) + Z^0_\mu(W^+_\nu \partial_\mu W^-_\mu - W^-_\nu \partial_\nu W^+_\mu)) - igs_w(\partial_\nu A_\mu(W^+_\mu W^-_\nu - W^+_\nu W^-_\mu) - A_\nu(W^+_\mu \partial_\nu W^-_\mu - W^-_\mu \partial_\nu W^+_\mu) + A_\mu(W^+_\nu \partial_\mu W^-_\mu - W^-_\mu \partial_\nu W^+_\mu)) - \tfrac{1}{2}g^2 W^+_\mu W^-_\mu W^+_\nu W^-_\nu + \tfrac{1}{2}g^2 W^+_\mu W^-_\nu W^+_\mu W^-_\nu + g^2 c_w^2(Z^0_\mu W^+_\mu Z^0_\nu W^-_\nu - Z^0_\mu Z^0_\mu W^+_\nu W^-_\nu) + g^2 s_w^2(A_\mu W^+_\mu A_\nu W^-_\nu - A_\mu A_\mu W^+_\nu W^-_\nu) + g^2 s_w c_w(A_\mu Z^0_\nu(W^+_\mu W^-_\nu - W^+_\nu W^-_\mu) - 2A_\mu Z^0_\mu W^+_\nu W^-_\nu) - \tfrac{1}{2}\partial_\mu H \partial_\mu H - 2M^2\alpha_h H^2 - \partial_\mu \phi^+ \partial_\mu \phi^- - \tfrac{1}{2}\partial_\mu \phi^0 \partial_\mu \phi^0 - \beta_h\left(\tfrac{2M^2}{g^2} + \tfrac{2M}{g}H + \tfrac{1}{2}(H^2 + \phi^0\phi^0 + 2\phi^+\phi^-)\right) + \tfrac{2M^4}{g^2}\alpha_h - g\alpha_h M\left(H^3 + H\phi^0\phi^0 + 2H\phi^+\phi^-\right) - \tfrac{1}{8}g^2\alpha_h\left(H^4 + (\phi^0)^4 + 4(\phi^+\phi^-)^2 + 4(\phi^0)^2\phi^+\phi^- + 4H^2\phi^+\phi^- + 2(\phi^0)^2H^2\right) - gMW^+_\mu W^-_\mu H - \tfrac{1}{2}g\tfrac{M}{c_w^2}Z^0_\mu Z^0_\mu H - \tfrac{1}{2}ig\left(W^+_\mu(\phi^0\partial_\mu\phi^- - \phi^-\partial_\mu\phi^0) - W^-_\mu(\phi^0\partial_\mu\phi^+ - \phi^+\partial_\mu\phi^0)\right) + \tfrac{1}{2}g\left(W^+_\mu(H\partial_\mu\phi^- - \phi^-\partial_\mu H) + W^-_\mu(H\partial_\mu\phi^+ - \phi^+\partial_\mu H)\right) + \tfrac{1}{2}g\tfrac{1}{c_w}(Z^0_\mu(H\partial_\mu\phi^0 - \phi^0\partial_\mu H) + M\left(\tfrac{1}{c_w}Z^0_\mu\partial_\mu\phi^0 + W^+_\mu\partial_\mu\phi^- + W^-_\mu\partial_\mu\phi^+\right) - ig\tfrac{s_w^2}{c_w}MZ^0_\mu(W^+_\mu\phi^- - W^-_\mu\phi^+) + igs_wMA_\mu(W^+_\mu\phi^- - W^-_\mu\phi^+) - ig\tfrac{1-2c_w^2}{2c_w}Z^0_\mu(\phi^+\partial_\mu\phi^- - \phi^-\partial_\mu\phi^+) + igs_wA_\mu(\phi^+\partial_\mu\phi^- - \phi^-\partial_\mu\phi^+) - \tfrac{1}{4}g^2W^+_\mu W^-_\mu\left(H^2 + (\phi^0)^2 + 2\phi^+\phi^-\right) - \tfrac{1}{8}g^2\tfrac{1}{c_w^2}Z^0_\mu Z^0_\mu\left(H^2 + (\phi^0)^2 + 2(2s_w^2 - 1)^2\phi^+\phi^-\right) - \tfrac{1}{2}g^2\tfrac{s_w^2}{c_w}Z^0_\mu\phi^0(W^+_\mu\phi^- + W^-_\mu\phi^+) - \tfrac{1}{2}ig^2\tfrac{s_w^2}{c_w}Z^0_\mu H(W^+_\mu\phi^- - W^-_\mu\phi^+) + \tfrac{1}{2}g^2s_wA_\mu\phi^0(W^+_\mu\phi^- + W^-_\mu\phi^+) + \tfrac{1}{2}ig^2s_wA_\mu H(W^+_\mu\phi^- - W^-_\mu\phi^+) - g^2\tfrac{s_w}{c_w}(2c_w^2 - 1)Z^0_\mu A_\mu\phi^+\phi^- - g^2s_w^2A_\mu A_\mu\phi^+\phi^- + \tfrac{1}{2}ig_s\lambda^a_{ij}(\bar{q}^a_i\gamma^\mu q^a_j)g^a_\mu - \bar{e}^\lambda(\gamma\partial + m^\lambda_e)e^\lambda - \bar{\nu}^\lambda(\gamma\partial + m^\lambda_\nu)\nu^\lambda - \bar{u}^\lambda_j(\gamma\partial + m^\lambda_u)u^\lambda_j - \bar{d}^\lambda_j(\gamma\partial + m^\lambda_d)d^\lambda_j + igs_wA_\mu\left(-(\bar{e}^\lambda\gamma^\mu e^\lambda) + \tfrac{2}{3}(\bar{u}^\lambda_j\gamma^\mu u^\lambda_j) - \tfrac{1}{3}(\bar{d}^\lambda_j\gamma^\mu d^\lambda_j)\right) + \tfrac{ig}{4c_w}Z^0_\mu\{(\bar{\nu}^\lambda\gamma^\mu(1+\gamma^5)\nu^\lambda) + (\bar{e}^\lambda\gamma^\mu(4s_w^2 - 1 - \gamma^5)e^\lambda) + (\bar{d}^\lambda_j\gamma^\mu(\tfrac{4}{3}s_w^2 - 1 - \gamma^5)d^\lambda_j) + (\bar{u}^\lambda_j\gamma^\mu(1 - \tfrac{8}{3}s_w^2 + \gamma^5)u^\lambda_j)\} + \tfrac{ig}{2\sqrt{2}}W^+_\mu\left((\bar{\nu}^\lambda\gamma^\mu(1+\gamma^5)U^{lep}_{\lambda\kappa}e^\kappa) + (\bar{u}^\lambda_j\gamma^\mu(1+\gamma^5)C_{\lambda\kappa}d^\kappa_j)\right) + \tfrac{ig}{2\sqrt{2}}W^-_\mu\left((\bar{e}^\kappa U^{lep\dagger}_{\kappa\lambda}\gamma^\mu(1+\gamma^5)\nu^\lambda) + (\bar{d}^\kappa_j C^\dagger_{\kappa\lambda}\gamma^\mu(1+\gamma^5)u^\lambda_j)\right) + \tfrac{ig}{2M\sqrt{2}}\phi^+\left(-m^\kappa_e(\bar{\nu}^\lambda U^{lep}_{\lambda\kappa}(1-\gamma^5)e^\kappa) + m^\lambda_u(\bar{\nu}^\lambda U^{lep}_{\lambda\kappa}(1+\gamma^5)e^\kappa) + \tfrac{ig}{2M\sqrt{2}}\phi^-\left(m^\lambda_e(\bar{e}^\lambda U^{lep\dagger}_{\lambda\kappa}(1+\gamma^5)\nu^\kappa) - m^\kappa_\nu(\bar{e}^\lambda U^{lep\dagger}_{\lambda\kappa}(1-\gamma^5)\nu^\kappa)\right) - \tfrac{g}{2}\tfrac{m^\lambda_e}{M}H(\bar{\nu}^\lambda\nu^\lambda) - \tfrac{g}{2}\tfrac{m^\lambda_e}{M}H(\bar{e}^\lambda e^\lambda) + \tfrac{ig}{2}\tfrac{m^\lambda_\nu}{M}\phi^0(\bar{\nu}^\lambda\gamma^5\nu^\lambda) - \tfrac{ig}{2}\tfrac{m^\lambda_e}{M}\phi^0(\bar{e}^\lambda\gamma^5 e^\lambda) - \tfrac{1}{4}\bar{\nu}_\lambda M^R_{\lambda\kappa}(1-\gamma_5)\bar{\nu}_\kappa - \tfrac{1}{4}\bar{\nu}_\lambda M^R_{\lambda\kappa}(1-\gamma_5)\bar{\nu}_\kappa + \tfrac{ig}{2M\sqrt{2}}\phi^+\left(-m^\kappa_d(\bar{u}^\lambda_j C_{\lambda\kappa}(1-\gamma^5)d^\kappa_j) + m^\lambda_u(\bar{u}^\lambda_j C_{\lambda\kappa}(1+\gamma^5)d^\kappa_j)\right) + \tfrac{ig}{2M\sqrt{2}}\phi^-\left(m^\lambda_d(\bar{d}^\lambda_j C^\dagger_{\lambda\kappa}(1+\gamma^5)u^\kappa_j) - m^\kappa_u(\bar{d}^\lambda_j C^\dagger_{\lambda\kappa}(1-\gamma^5)u^\kappa_j)\right) - \tfrac{g}{2}\tfrac{m^\lambda_u}{M}H(\bar{u}^\lambda_j u^\lambda_j) + \tfrac{g}{2}\tfrac{m^\lambda_d}{M}H(\bar{d}^\lambda_j d^\lambda_j) + \tfrac{ig}{2}\tfrac{m^\lambda_u}{M}\phi^0(\bar{u}^\lambda_j\gamma^5 u^\lambda_j) - \tfrac{ig}{2}\tfrac{m^\lambda_d}{M}\phi^0(\bar{d}^\lambda_j\gamma^5 d^\lambda_j) + \bar{G}^a\partial^2 G^a + g_s f^{abc}\partial_\mu\bar{G}^a G^b g^c_\mu + \bar{X}^+(\partial^2 - M^2)X^+ + \bar{X}^-(\partial^2 - M^2)X^- + \bar{X}^0(\partial^2 - \tfrac{M^2}{c_w^2})X^0 + \bar{Y}\partial^2 Y + igc_wW^+_\mu(\partial_\mu\bar{X}^0X^- - \partial_\mu\bar{X}^+X^0) + igs_wW^+_\mu(\partial_\mu\bar{Y}X^- - \partial_\mu\bar{X}^+Y) + igc_wW^-_\mu(\partial_\mu\bar{X}^-X^0 - \partial_\mu\bar{X}^0X^+) + igs_wW^-_\mu(\partial_\mu\bar{X}^-Y - \partial_\mu\bar{Y}X^+) + igc_wZ^0_\mu(\partial_\mu\bar{X}^+X^+ - \partial_\mu\bar{X}^-X^-) + igs_wA_\mu(\partial_\mu\bar{X}^+X^+ - \partial_\mu\bar{X}^-X^-) - \tfrac{1}{2}gM\left(\bar{X}^+X^+H + \bar{X}^-X^-H + \tfrac{1}{c_w^2}\bar{X}^0X^0H\right) + \tfrac{1-2c_w^2}{2c_w}igM\left(\bar{X}^+X^0\phi^+ - \bar{X}^-X^0\phi^-\right) + \tfrac{1}{2c_w}igM\left(\bar{X}^0X^-\phi^+ - \bar{X}^0X^+\phi^-\right) + igMs_w\left(\bar{X}^0X^-\phi^+ - \bar{X}^0X^+\phi^-\right) + \tfrac{1}{2}igM\left(\bar{X}^+X^+\phi^0 - \bar{X}^-X^-\phi^0\right).$$

SN1987A 24.2.1987 :: 170kly

SN1987A 24.2.1987 :: 170kly

SN1987A 24.2.1987 :: 170kly

https://en.wikipedia.org/wiki/SN_1987A

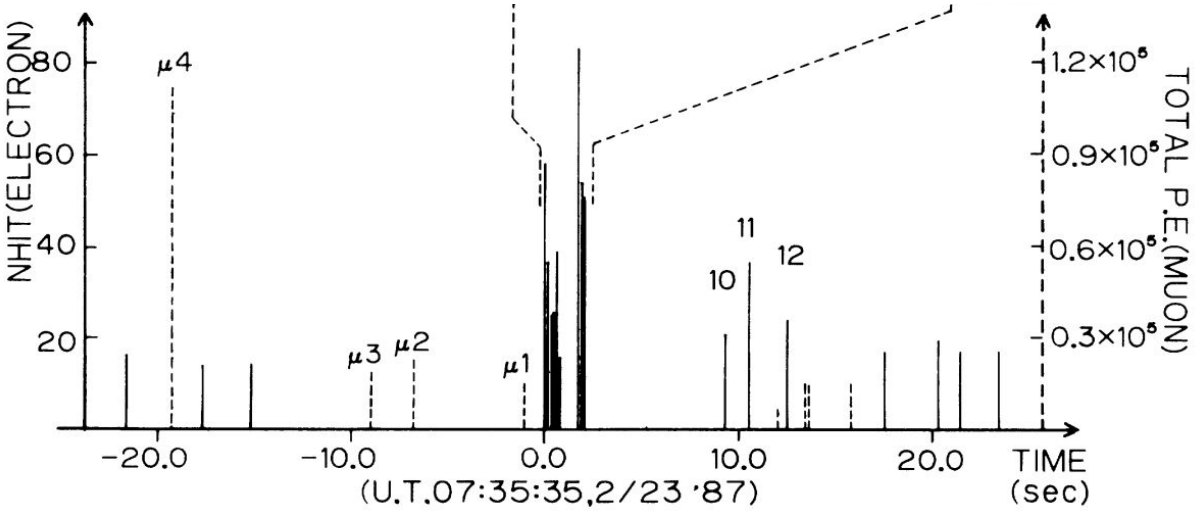
https://www-sk.icrr.u-tokyo.ac.jp/en/news/detail/324

Experiment Kamiokande(-II) :: 3kt H_20

VOLUME 58, NUMBER 14 PHYSICAL REVIEW LETTERS 6 APRIL 1987

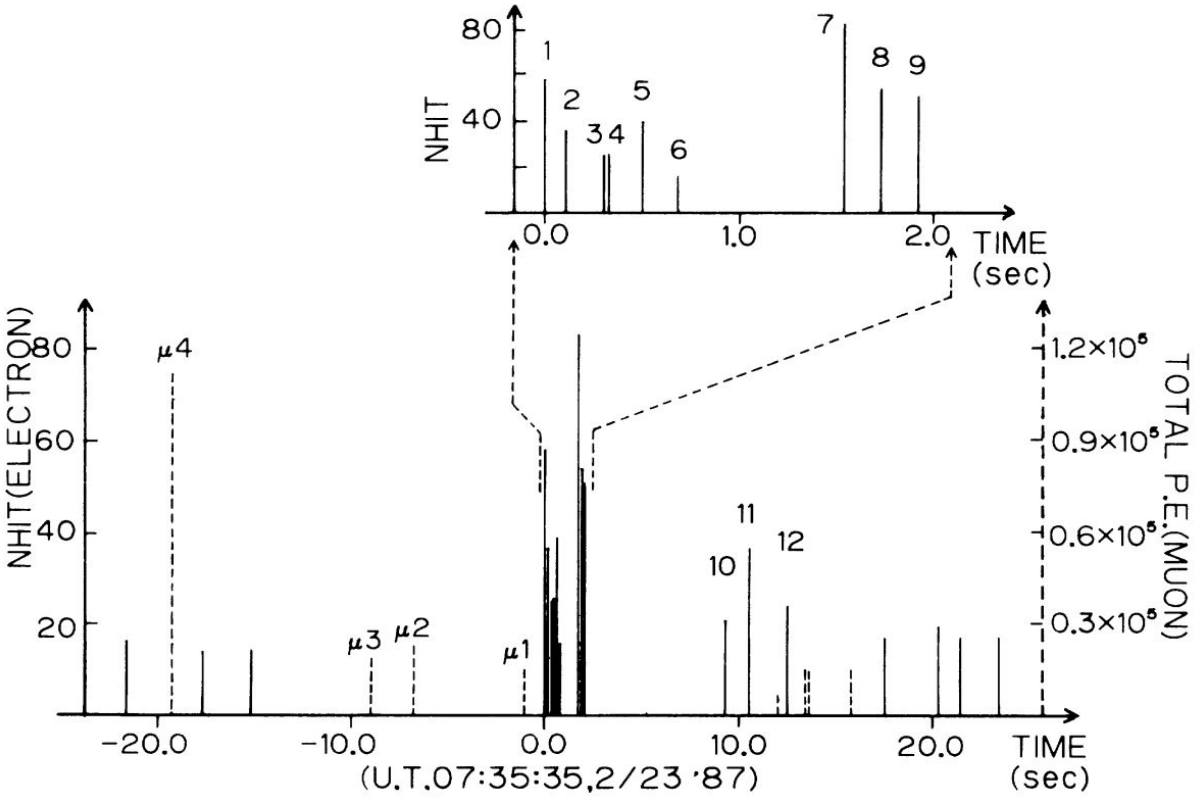
Observation of a Neutrino Burst from the Supernova SN1987A

K. Hirata,[(a)] T. Kajita,[(a)] M. Koshiba,[(a,b)] M. Nakahata,[(b)] Y. Oyama,[(b)]
N. Sato,[(c)] A. Suzuki,[(b)] M. Takita,[(b)] and Y. Totsuka[(a,c)]
University of Tokyo, Tokyo 113, Japan



Experiment Kamiokande(-II) :: 3kt H_20

- Počátky neutrinové astronomie! :-)



Experiment Kamiokande(-II) :: 3kt H_20

- Počátky neutrinové astronomie! :-)
- 1 Erg is a unit of energy equal to 10^{-7} Joules (100 nJ)

form volume distribution. Additional support is provided by the correlation in angle of the first two observed events with the direction to SN1987A. The event burst occurred roughly 18 h prior to the first optical sighting.[1] Correcting for energy-dependent detection efficiency, and assuming that nine of the twelve events are due to $\bar{\nu}_e p^+ \rightarrow e^+ n$, we obtain an integral flux of 1.0×10^{10} $\bar{\nu}_e$ cm^{-2} for the burst, where the $\bar{\nu}_e$ energy (the observed electron energy plus 1.3 MeV) is above 8.8 MeV. This, in turn, leads to the $\bar{\nu}_e$ output of SN1987A of 8×10^{52} ergs for an assumed average energy of 15 MeV.

This observation is the first direct observation in neutrino astronomy, and coincides remarkably well with the current model of supernova collapse and neutron-star formation.[6] In that model an aged, massive star, having exhausted its nuclear fuel, undergoes a supernova explosion. In supernovae of Type II almost all of the gravitational binding energy of the resultant neutron star, $\sim3\times10^{53}$ ergs, is radiated within a few seconds in the form of 10^{58} neutrinos of all flavors with average energy in the vicinity of 10–15 MeV.

Experiment Kamiokande(-II) :: 3kt H_20

- Počátky neutrinové astronomie! :-)
- 1 Erg is a unit of energy equal to 10^{-7} Joules (100 nJ)

form volume distribution. Additional support is provided by the correlation in angle of the first two observed events with the direction to SN1987A. The event burst occurred roughly 18 h prior to the first optical sighting.[1] Correcting for energy-dependent detection efficiency, and assuming that nine of the twelve events are due to $\bar{\nu}_e p^+ \rightarrow e^+ n$, we obtain an integral flux of 1.0×10^{10} $\bar{\nu}_e$ cm^{-2} for the burst, where the $\bar{\nu}_e$ energy (the observed electron energy plus 1.3 MeV) is above 8.8 MeV. This, in turn, leads to the $\bar{\nu}_e$ output of SN1987A of 8×10^{52} ergs for an assumed average energy of 15 MeV.

This observation is the first direct observation in neutrino astronomy, and coincides remarkably well with the current model of supernova collapse and neutron-star formation.[6] In that model an aged, massive star, having exhausted its nuclear fuel, undergoes a supernova explosion. In supernovae of Type II almost all of the gravitational binding energy of the resultant neutron star, $\sim3\times10^{53}$ ergs, is radiated within a few seconds in the form of 10^{58} neutrinos of all flavors with average energy in the vicinity of 10–15 MeV.

Experiment Kamiokande(-II) :: 3kt H_20

- Počátky neutrinové astronomie! :-)
- 1 Erg is a unit of energy equal to 10^{-7} Joules (100 nJ)

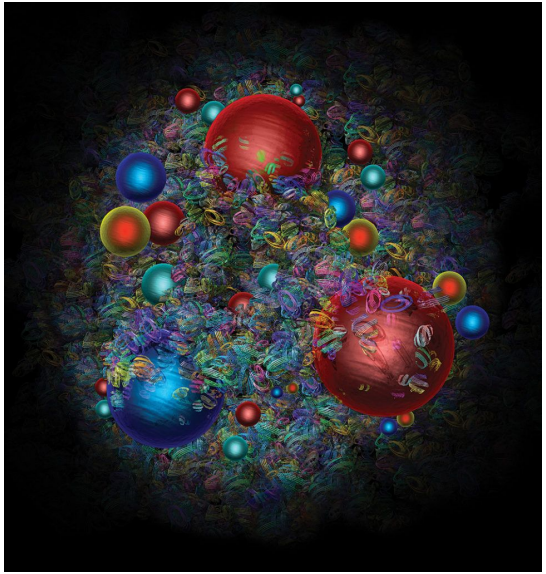
form volume distribution. Additional support is provided by the correlation in angle of the first two observed events with the direction to SN1987A. The event burst occurred roughly 18 h prior to the first optical sighting.[1] Correcting for energy-dependent detection efficiency, and assuming that nine of the twelve events are due to $\bar{\nu}_e p^+ \rightarrow e^+ n$, we obtain an integral flux of 1.0×10^{10} $\bar{\nu}_e$ cm^{-2} for the burst, where the $\bar{\nu}_e$ energy (the observed electron energy plus 1.3 MeV) is above 8.8 MeV. This, in turn, leads to the $\bar{\nu}_e$ output of SN1987A of 8×10^{52} ergs for an assumed average energy of 15 MeV.

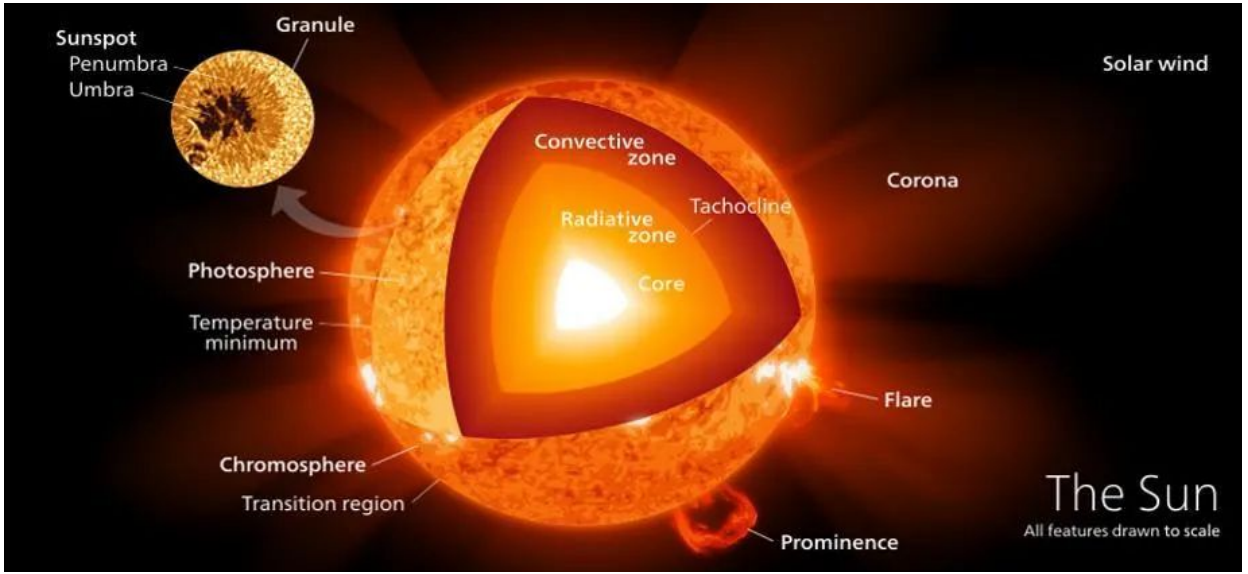
This observation is the first direct observation in neutrino astronomy, and coincides remarkably well with the current model of supernova collapse and neutron-star formation.[6] In that model an aged, massive star, having exhausted its nuclear fuel, undergoes a supernova explosion. In supernovae of Type II almost all of the gravitational binding energy of the resultant neutron star, $\sim3\times10^{53}$ ergs, is radiated within a few seconds in the form of 10^{58} neutrinos of all flavors with average energy in the vicinity of 10–15 MeV.

Proton

- Nejjednodušší a nejčastější jádro atomu ve Vesmíru
- Stabilní
- Palivo hvězd, ve kterých vznikají těžší jádra a energie

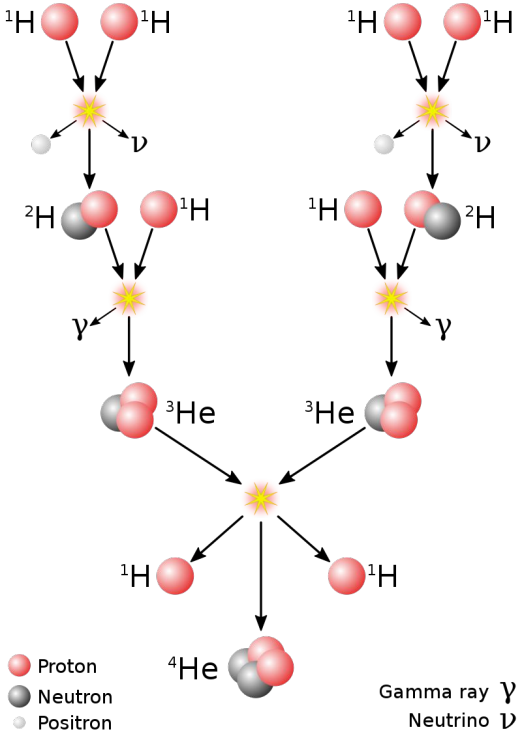
https://home.cern/news/news/physics/proton-century

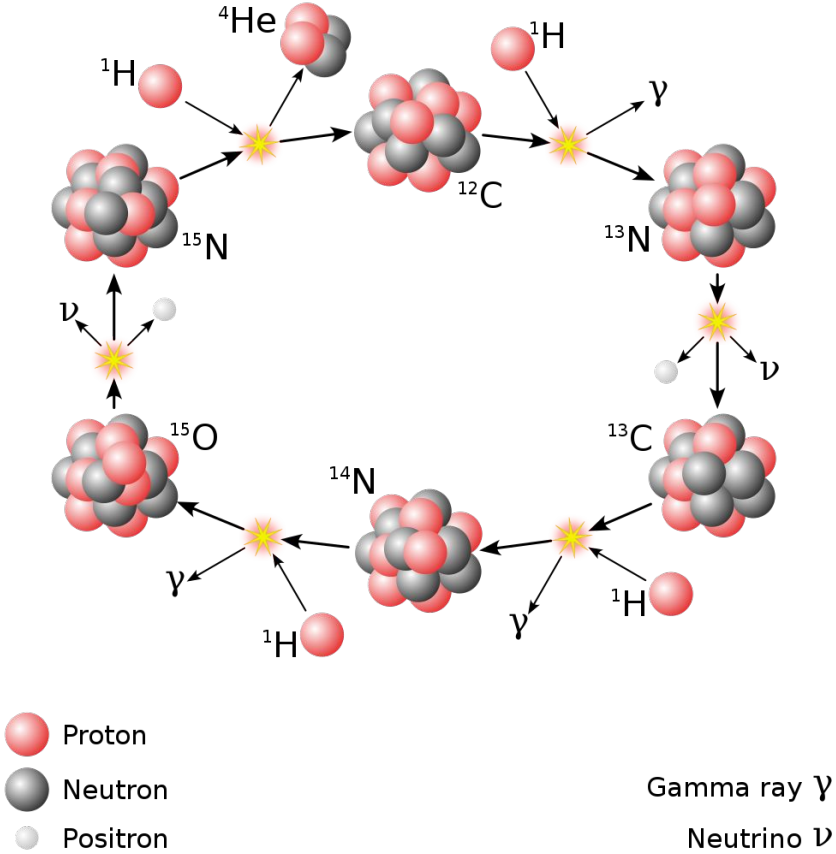




https://www.forbes.com/sites/startswithabang/2017/09/05/the-suns-energy-doesnt-come-from-fusing-hydrogen-into-helium-mostly/

Neutrina ze Slunce: pp a CNO cyklus





https://en.wikipedia.org/wiki/CNO_cycle

Fyzikální program: Oscilace neutrin

- Zde ilustrace na měření oscilací reaktorového experimentu Daya Bay a JUNO.

WW Looking at the flux of $\bar{\nu}_e$ as a function of distance and energy

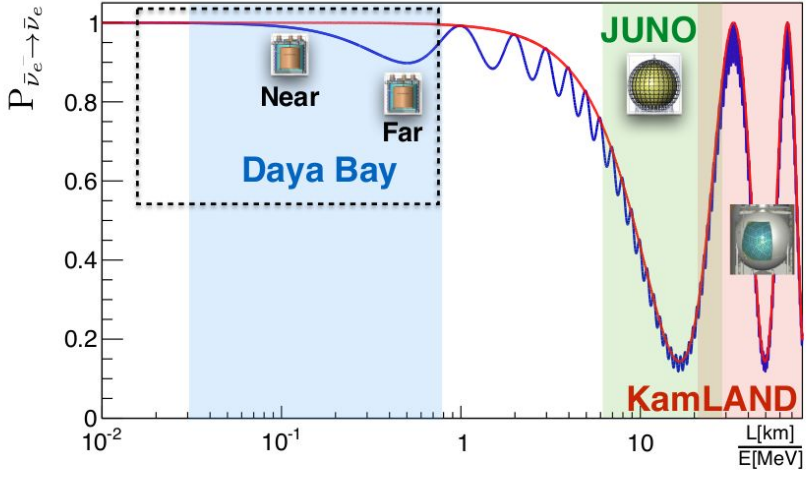
WW Some $\bar{\nu}_e$'s disappear due to neutrino oscillation

Medium baseline | Short baseline

$$P_{\bar{\nu}_e \to \bar{\nu}_e}(L,E) = 1 - \sin^2 2\theta_{12} \cos^4 \theta_{13} \sin^2 \frac{\Delta m^2_{21} L}{4E} - \sin^2 2\theta_{13} \left(\cos^2 \theta_{12} \sin^2 \frac{\Delta m^2_{31} L}{4E} + \sin^2 \theta_{12} \sin^2 \frac{\Delta m^2_{32} L}{4E} \right)$$

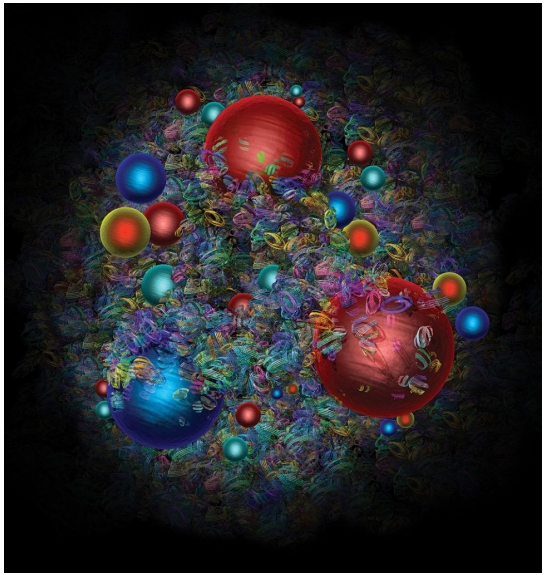
Oscilace neutrin

$$\begin{pmatrix} \nu_e \\ \nu_\mu \\ \nu_\tau \end{pmatrix} \leftrightarrow \begin{pmatrix} \nu_1 \\ \nu_2 \\ \nu_3 \end{pmatrix} \leftrightarrow \begin{pmatrix} \nu_e \\ \nu_\mu \\ \nu_\tau \end{pmatrix}$$



Rozpad protonu?

- Možný v některých teoriích za současný Standardní model mikrosvěta.
- Hledejme tedy!
- Kdyby byl pozorován, je to náznak nové fyziky!



$p^+ \rightarrow e^+ \pi^0$ $\quad \pi^0 \rightarrow \gamma\gamma$

$p^+ \rightarrow \mu^+ \pi^0$ $\quad \mu^+ \rightarrow e^+ \bar{\nu}_\mu \nu_e$

$\pi^0 \rightarrow \gamma\gamma$